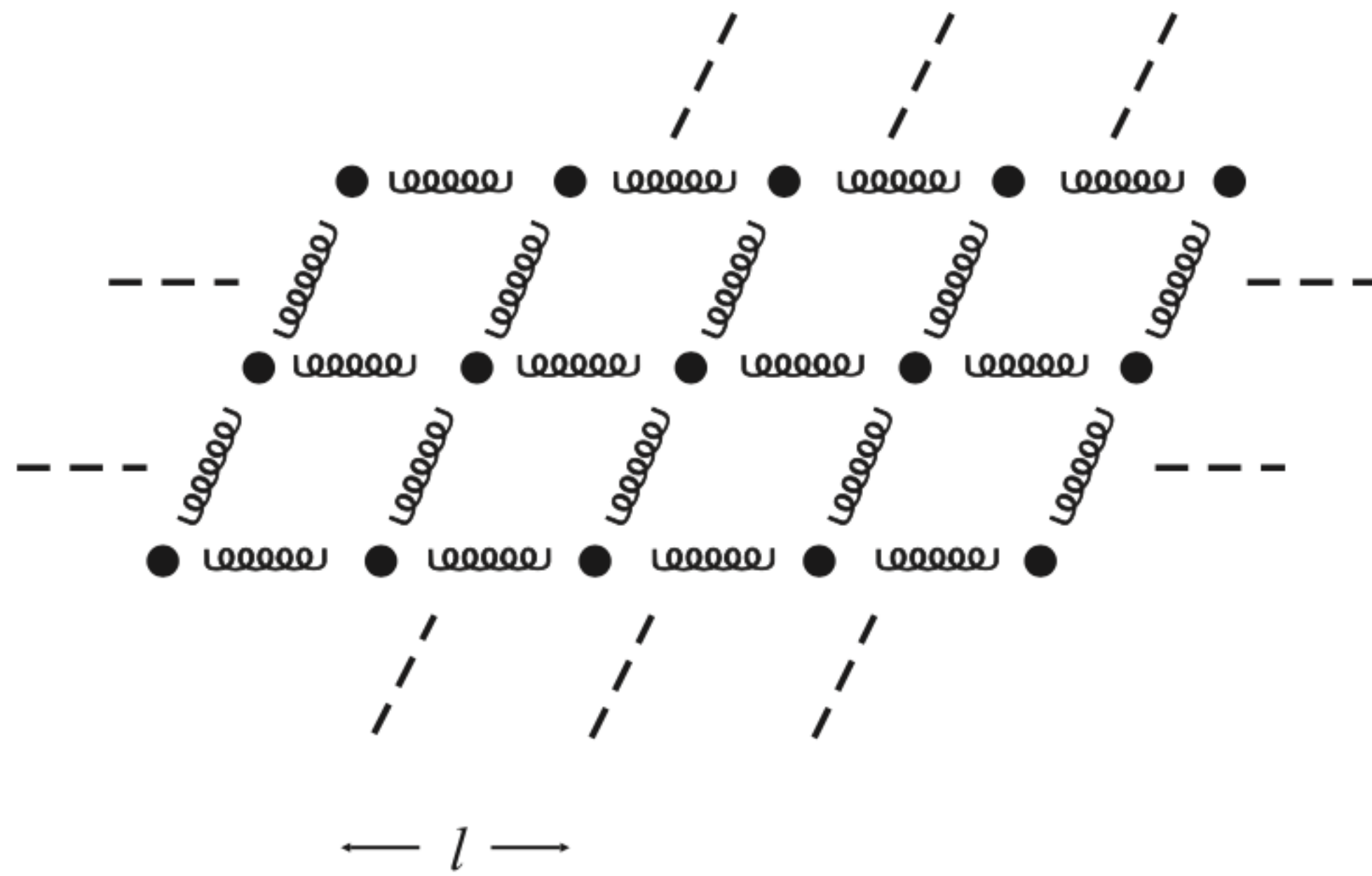


Cosmic Strings and Gravitational Waves from the Early Universe

Saturday Morning of Theoretical Physics, Nov 2024

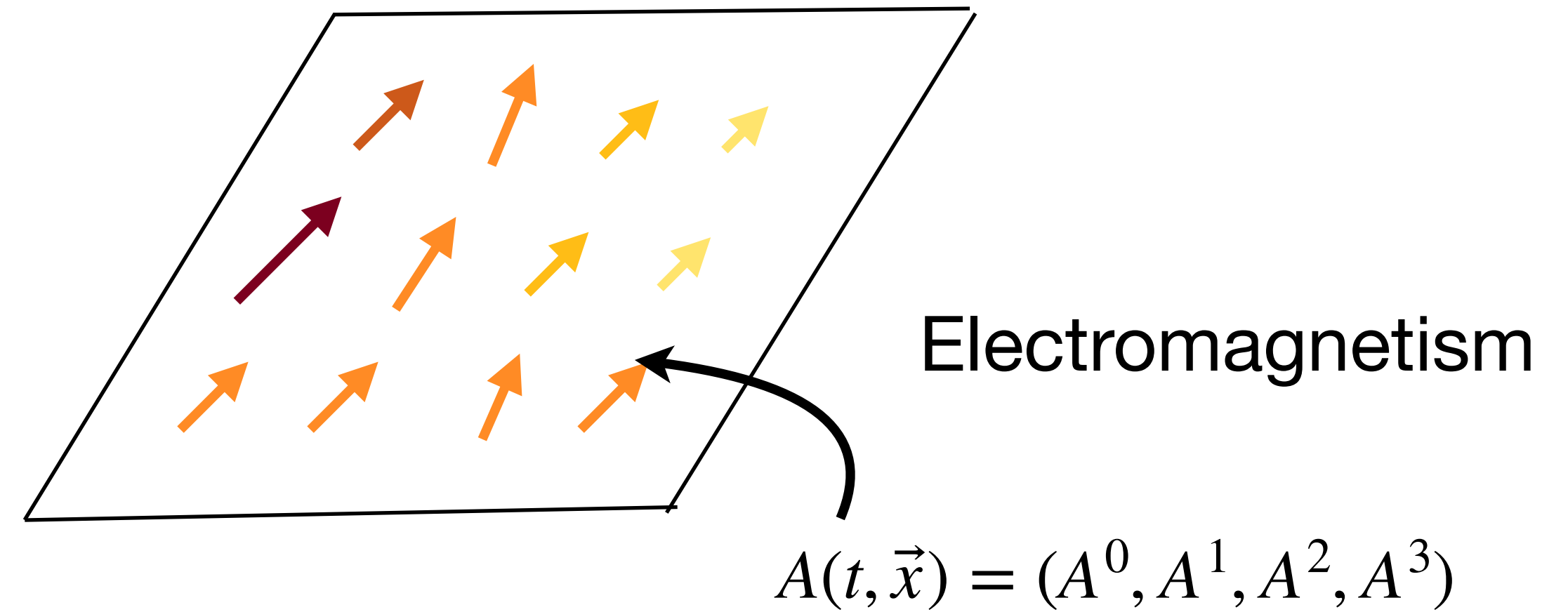
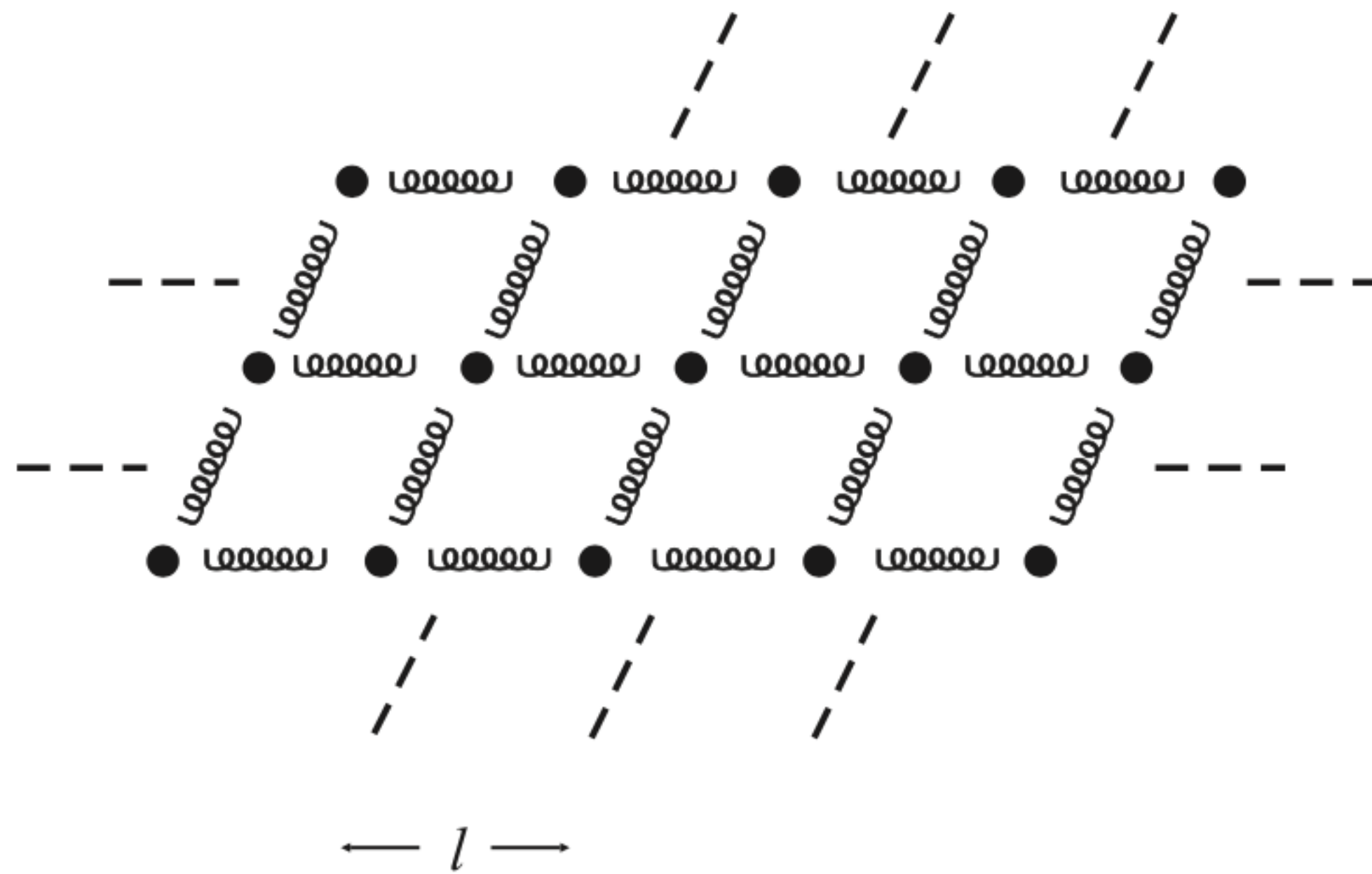
Edward Hardy

Classical Fields



[Zee, Quantum Field Theory in a Nutshell]

Classical Fields

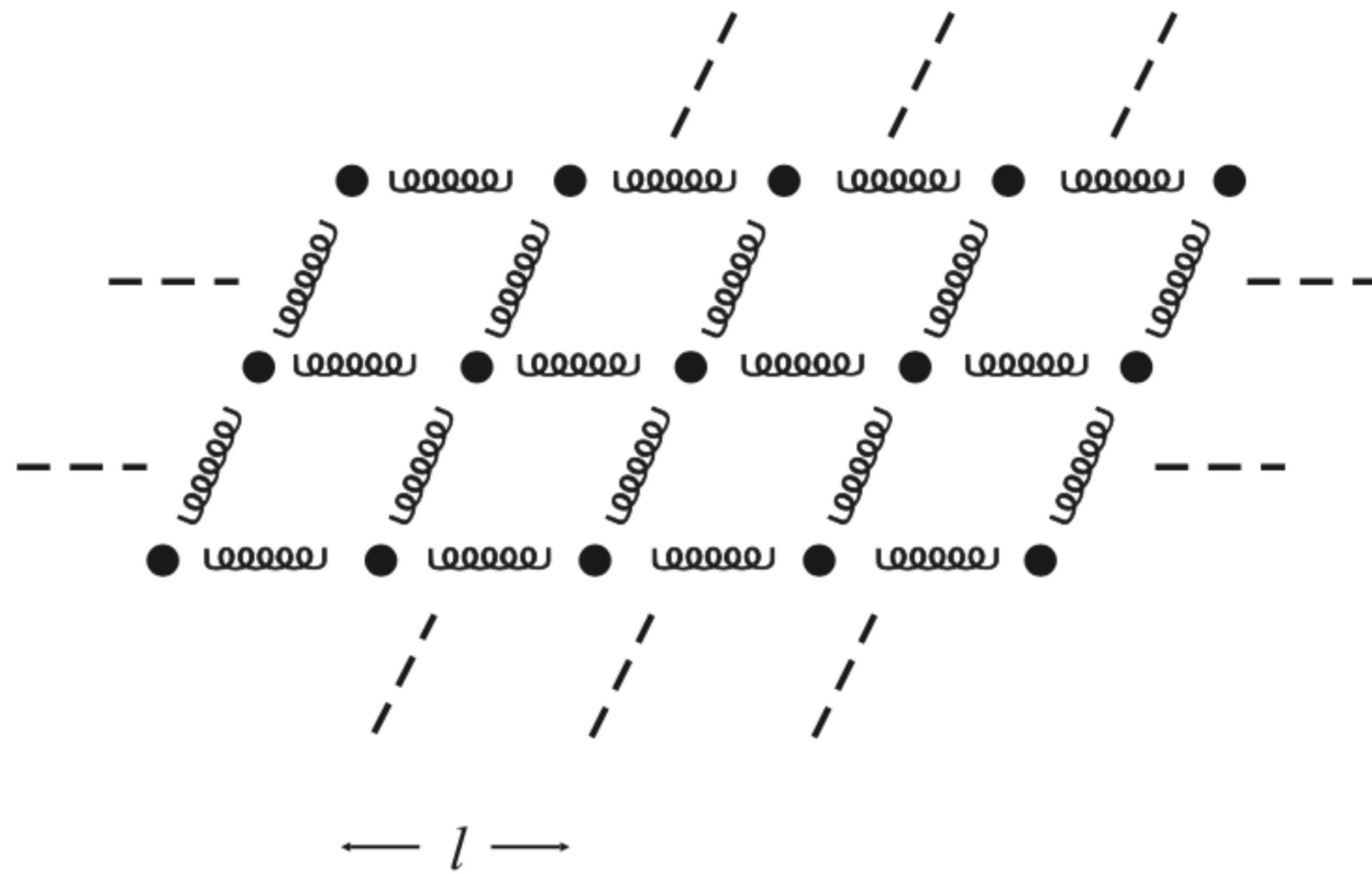


$$F_{\mu\nu} = \partial_\mu A_\nu - \partial_\nu A_\mu$$

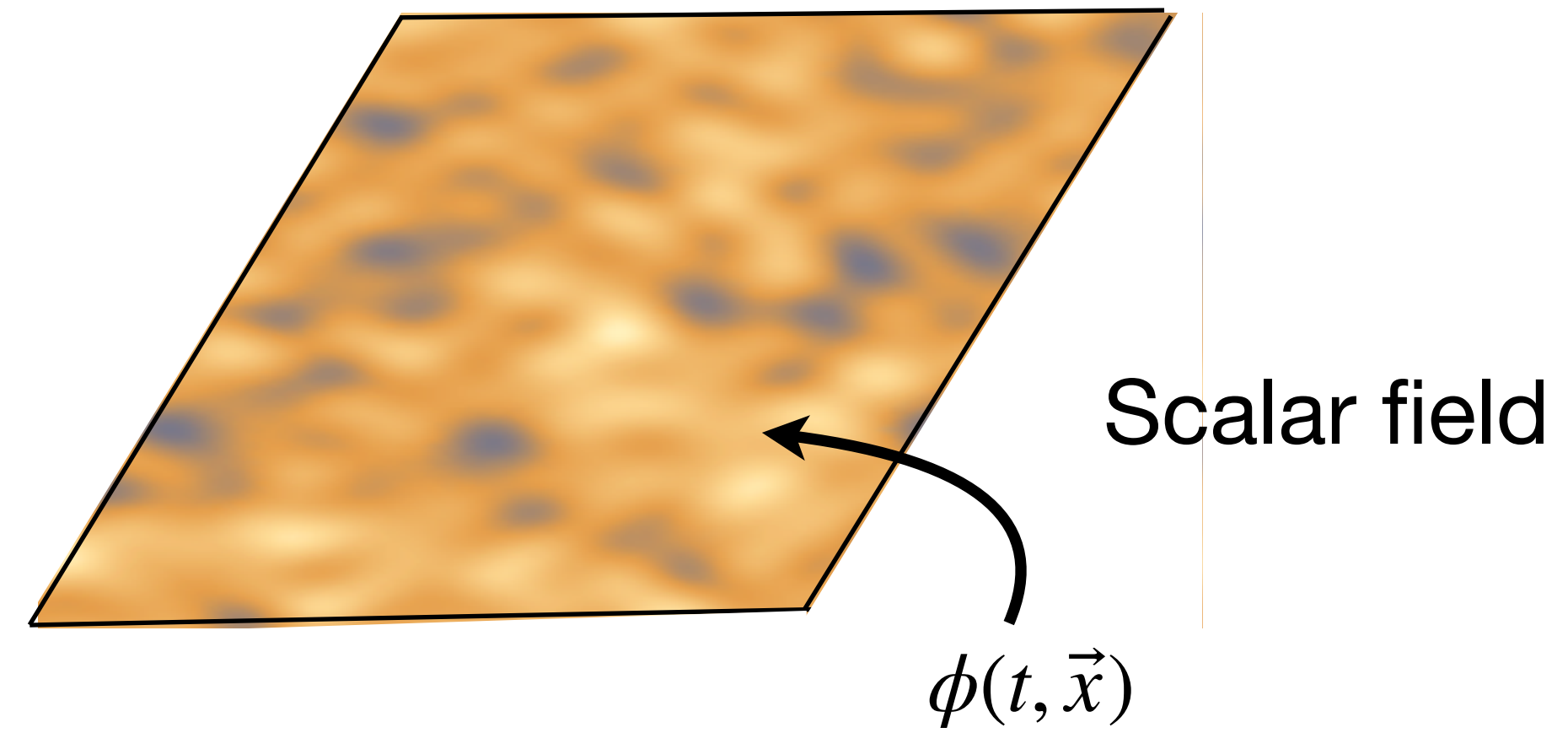
$$F^{\mu\nu} = \begin{bmatrix} 0 & -E_x/c & -E_y/c & -E_z/c \\ E_x/c & 0 & -B_z & B_y \\ E_y/c & B_z & 0 & -B_x \\ E_z/c & -B_y & B_x & 0 \end{bmatrix}$$

[Zee, Quantum Field Theory in a Nutshell]

Classical Fields



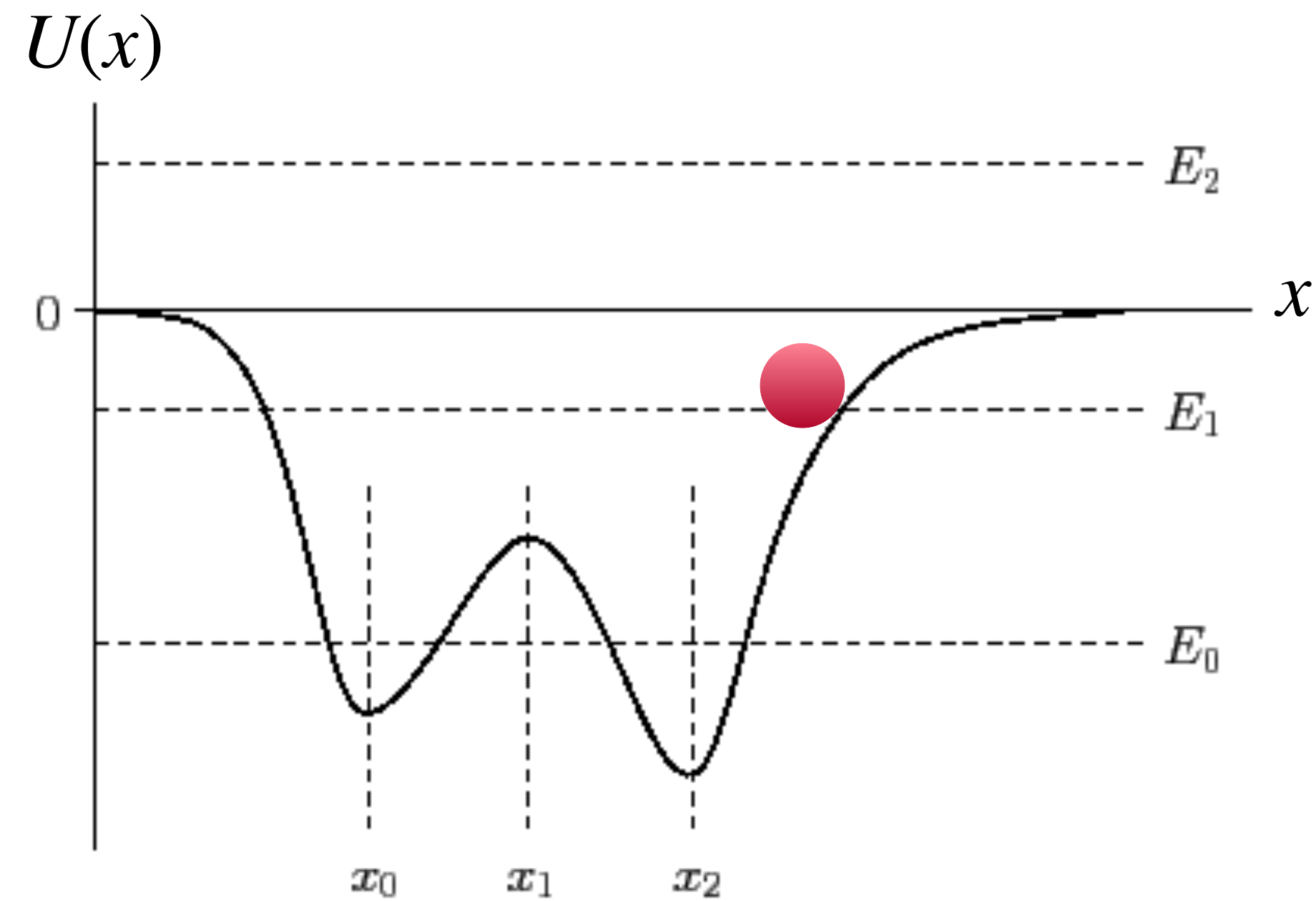
[Zee, Quantum Field Theory in a Nutshell]



Consider a new scalar field

Classical limit = large occupation number

Particle in a potential

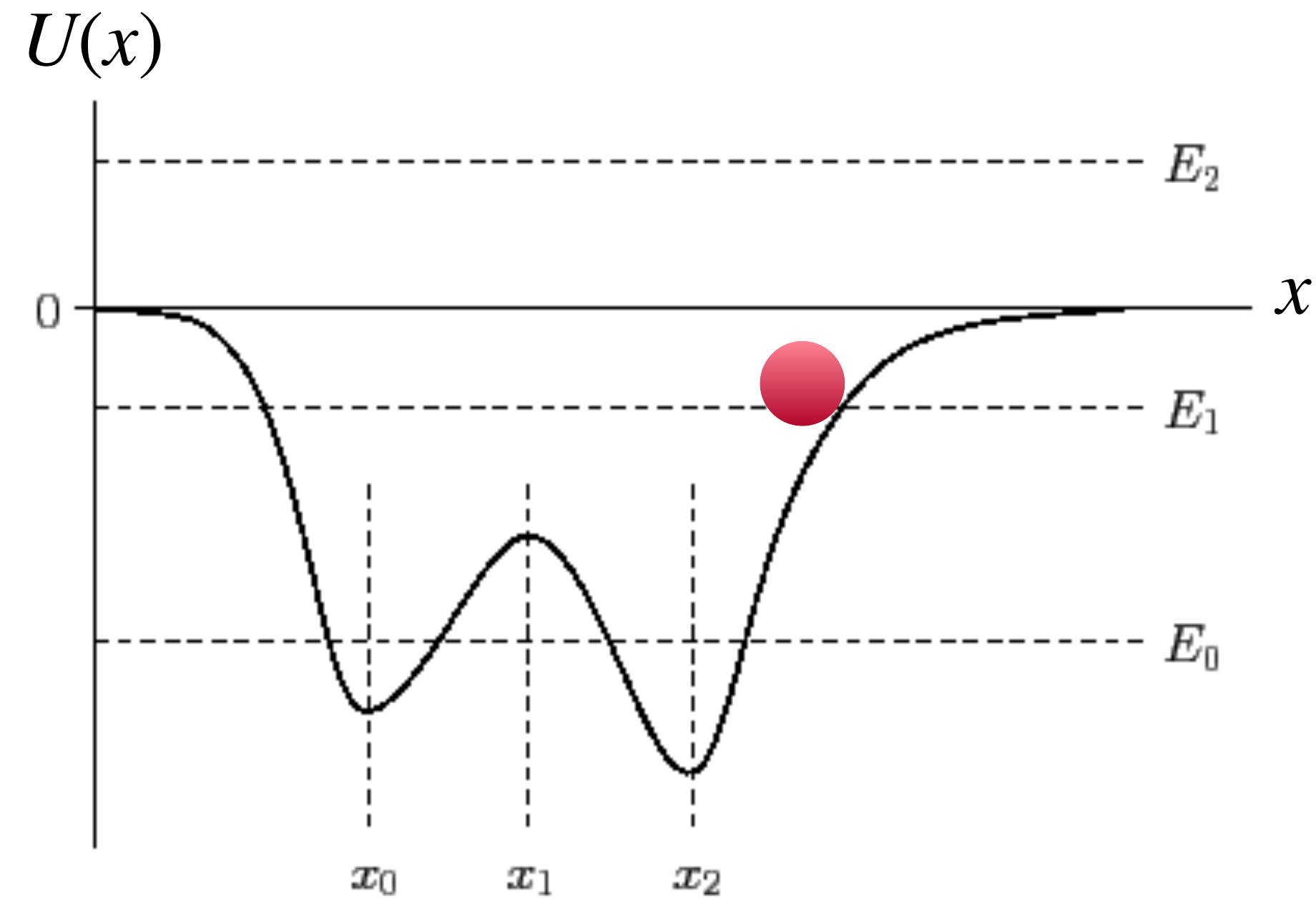


$$E = T + U(x) = \frac{1}{2}m\dot{x}^2 + U(x)$$

One degree of freedom

x

Particle in a potential



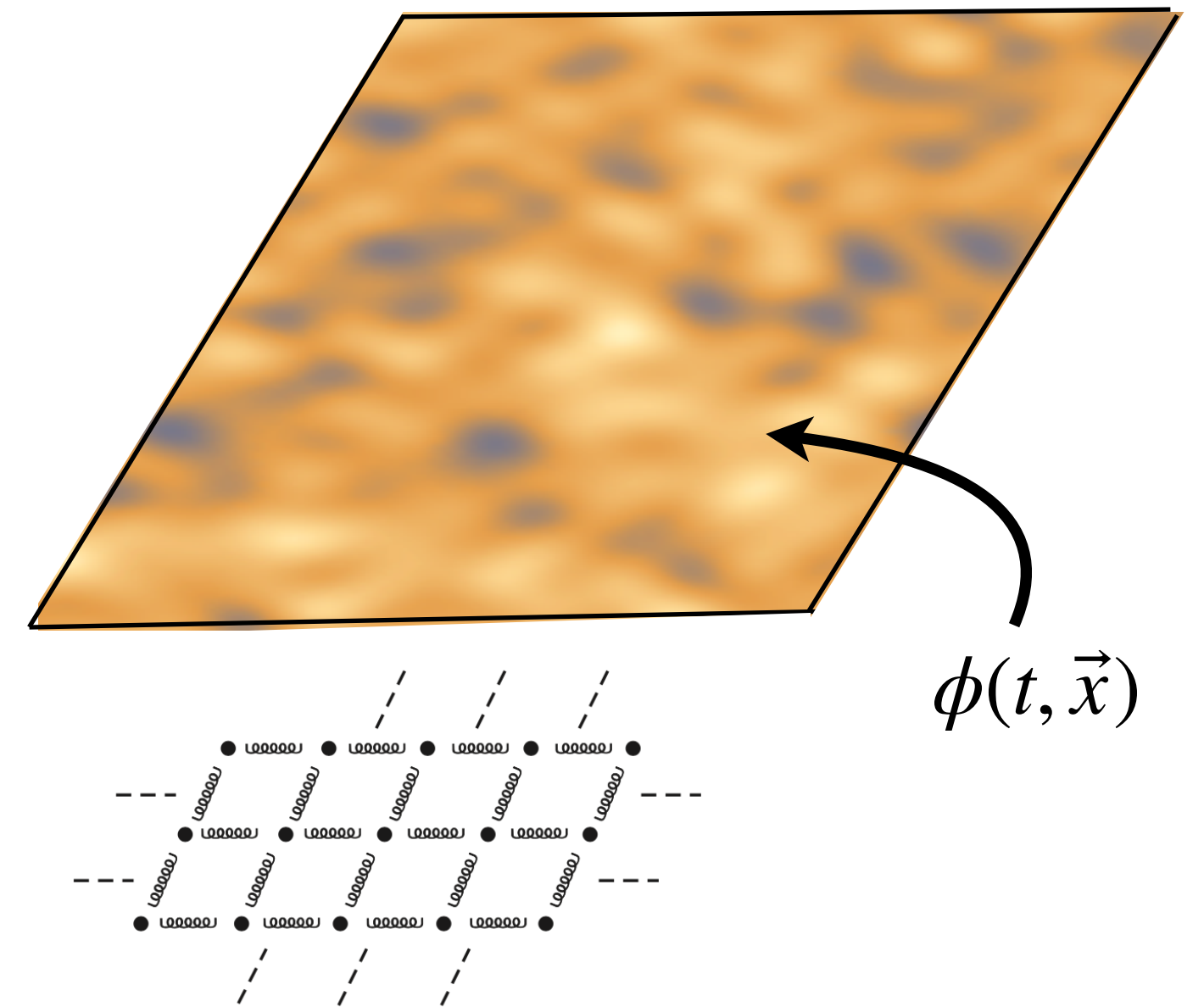
$$E = T + U(x) = \frac{1}{2}m\dot{x}^2 + U(x)$$

One degree of freedom

x



$\phi(\vec{x})$



$$\rho \sim (\partial_t \phi)^2 + (\partial_i \phi)^2 + U(\phi), \quad E = \int d^3x \rho$$

One degree of freedom at each point in space

Symmetry preserving potential

E.g. $U(\phi) = f_a^2 \phi^2 + \lambda \phi^4$

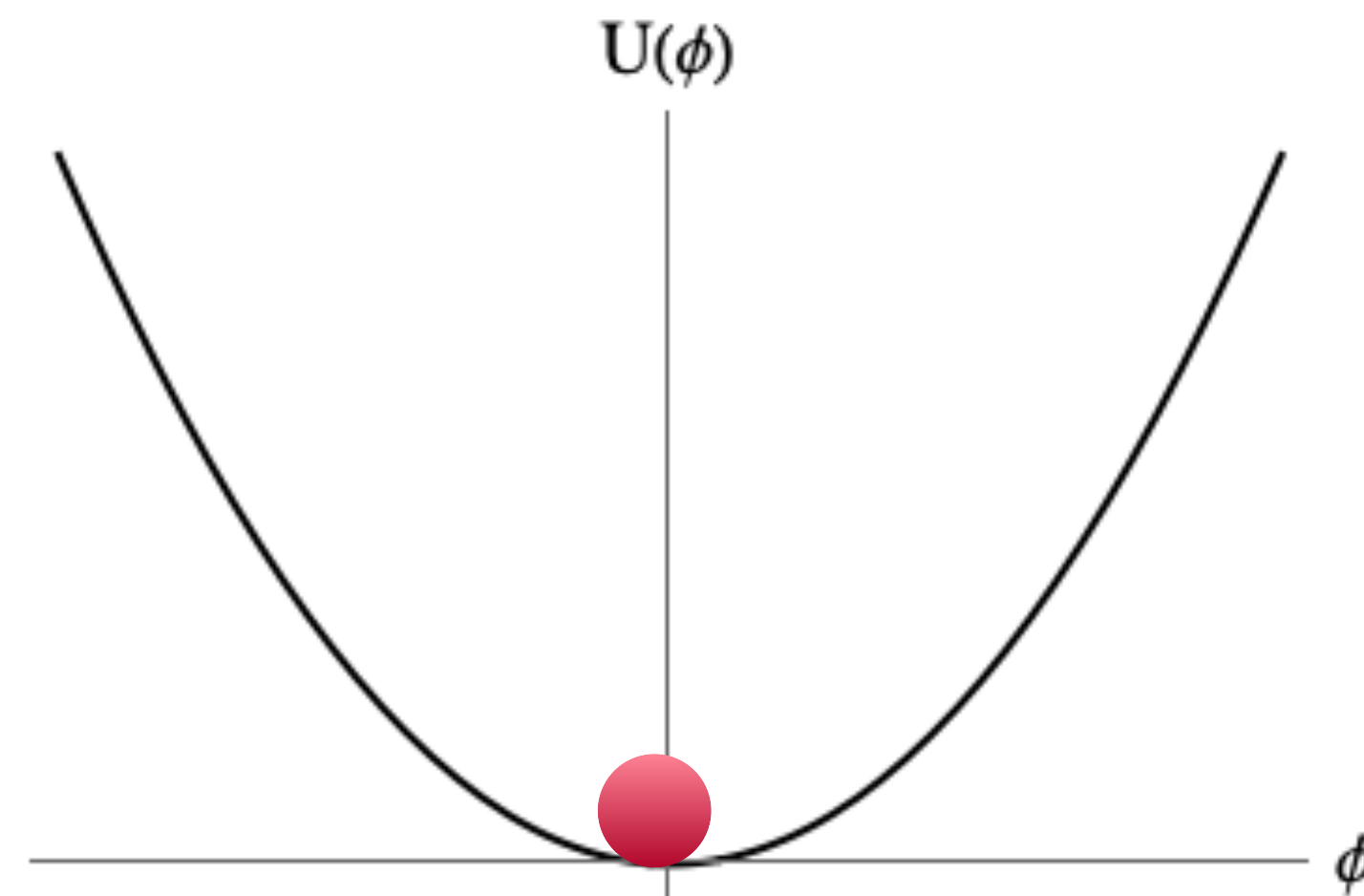
Symmetry:

$$Z_2 : \phi \rightarrow -\phi$$

$$U(-\phi) = U(\phi)$$

Preserved by the vacuum

$$\phi_{\min} = 0 = -\phi_{\min}$$



$$\rho \sim (\partial_t \phi)^2 + (\partial_i \phi)^2 + U(\phi),$$

Symmetry preserving potential

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Symmetry:

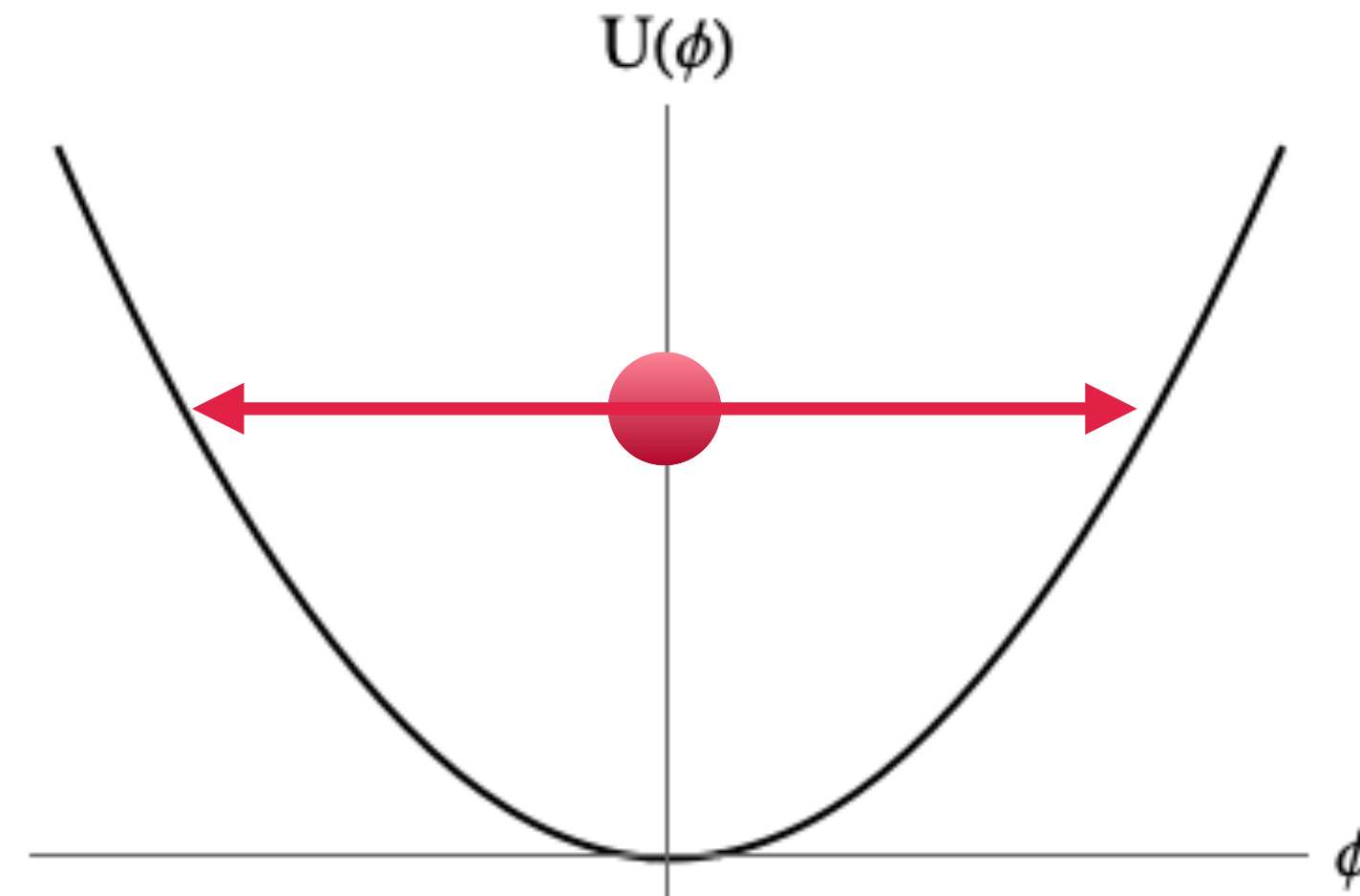
$$Z_2 : \phi \rightarrow -\phi$$

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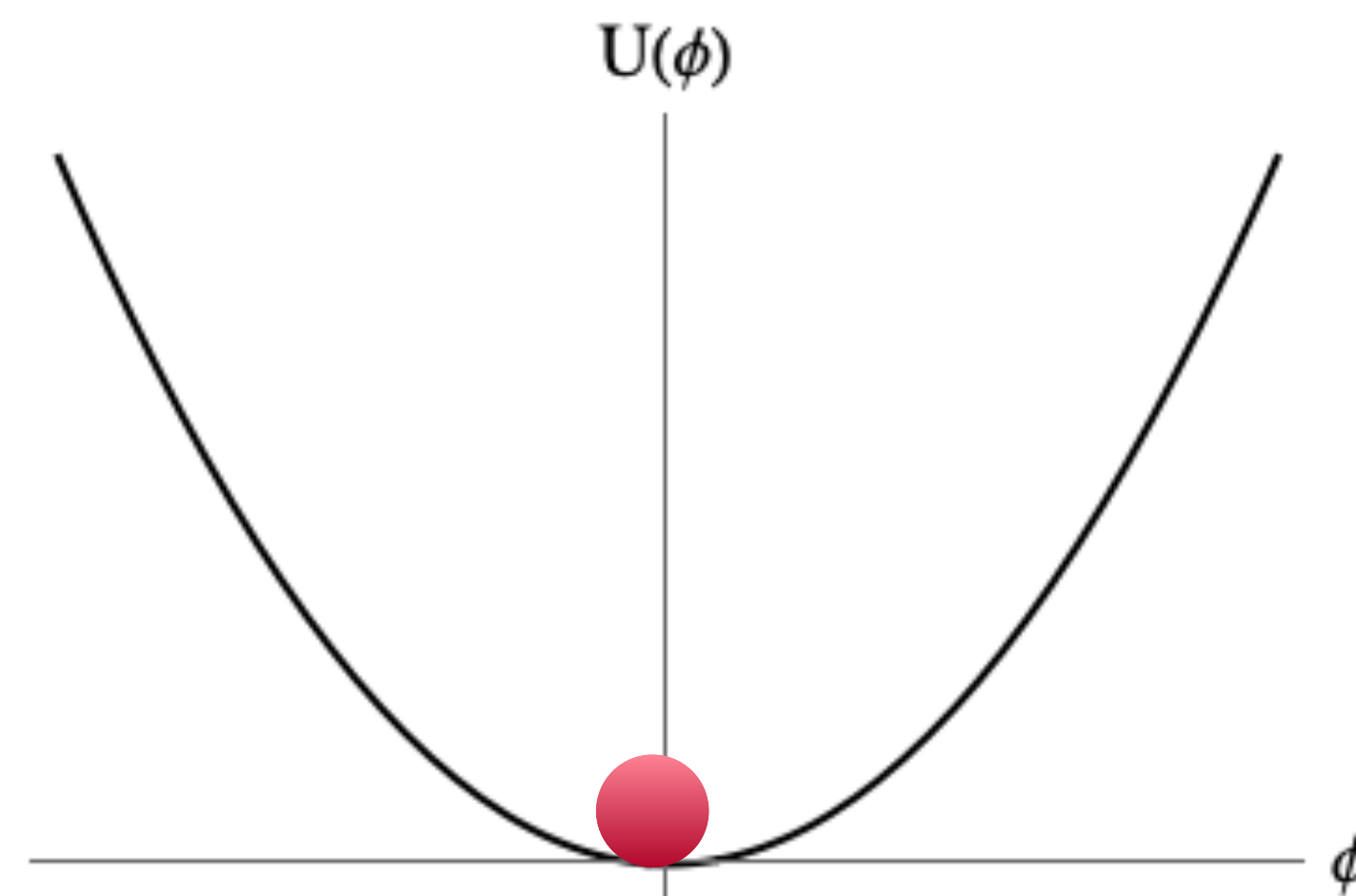
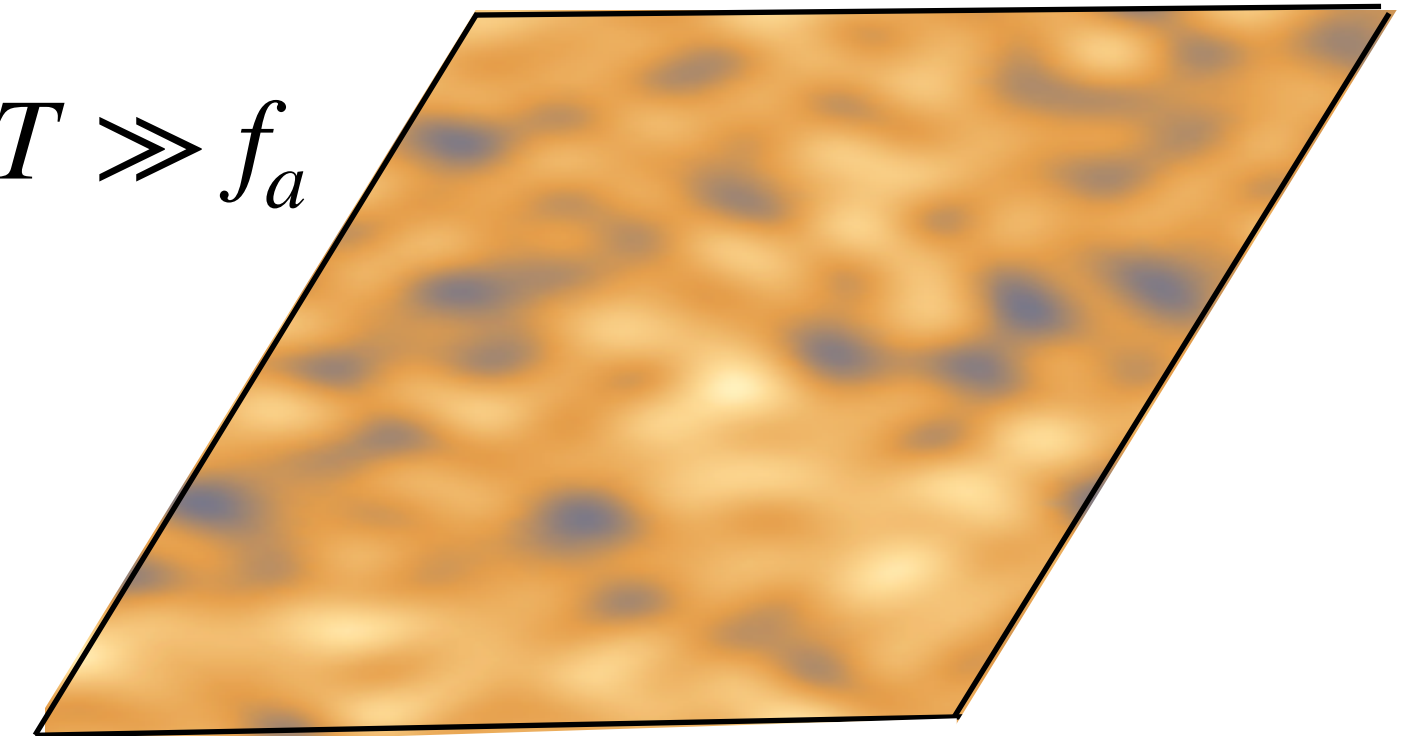
Preserved by the vacuum

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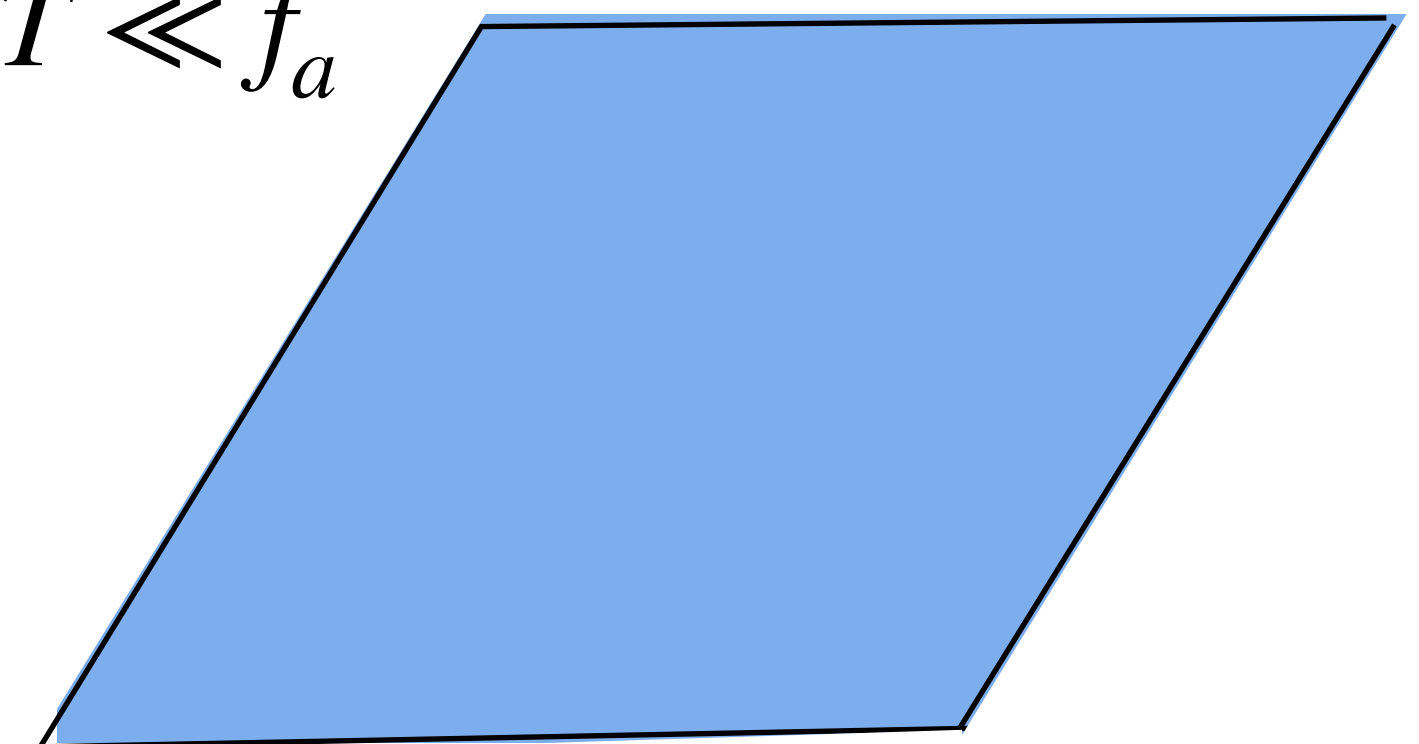
$$\rho \sim (\partial_t \phi)^2 + (\partial_i \phi)^2 + U(\phi),$$



High $T \gg f_a$



Low $T \ll f_a$



Spontaneous symmetry breaking

$$\text{E.g. } U(\phi) = -f_a^2 \phi^2 + \lambda \phi^4$$

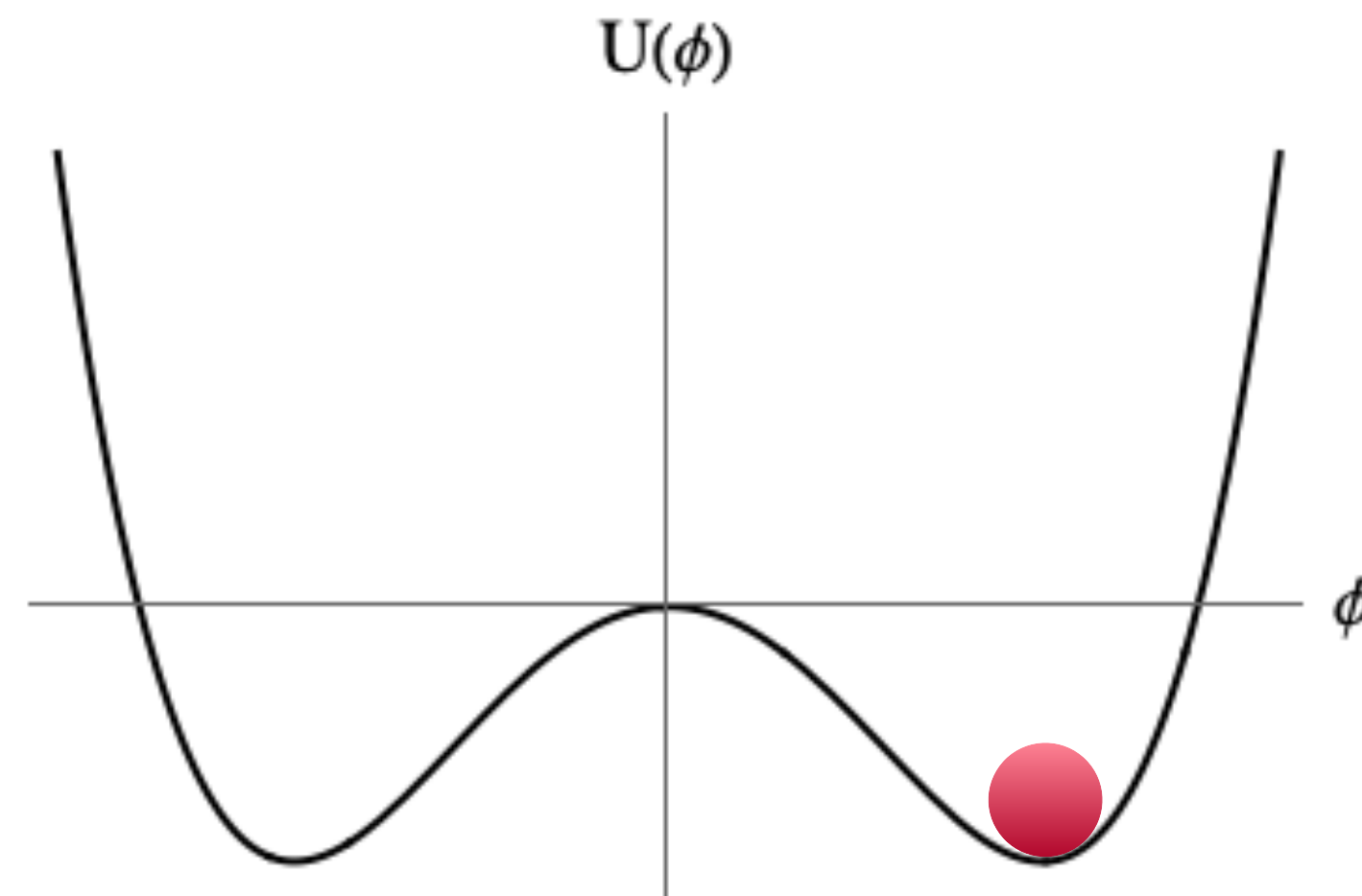
Symmetry:

$$Z_2 : \phi \rightarrow -\phi$$

$$U(-\phi) = U(\phi)$$

Broken by the vacuum

$$\phi_{\min} = \pm \frac{f_a}{\sqrt{2\lambda}} \neq -\phi_{\min}$$



$$\rho \sim (\partial_t \phi)^2 + (\partial_i \phi)^2 + U(\phi),$$

Spontaneous symmetry breaking

E.g. $U(\phi) = -f_a^2\phi^2 + \lambda\phi^4$

Symmetry:

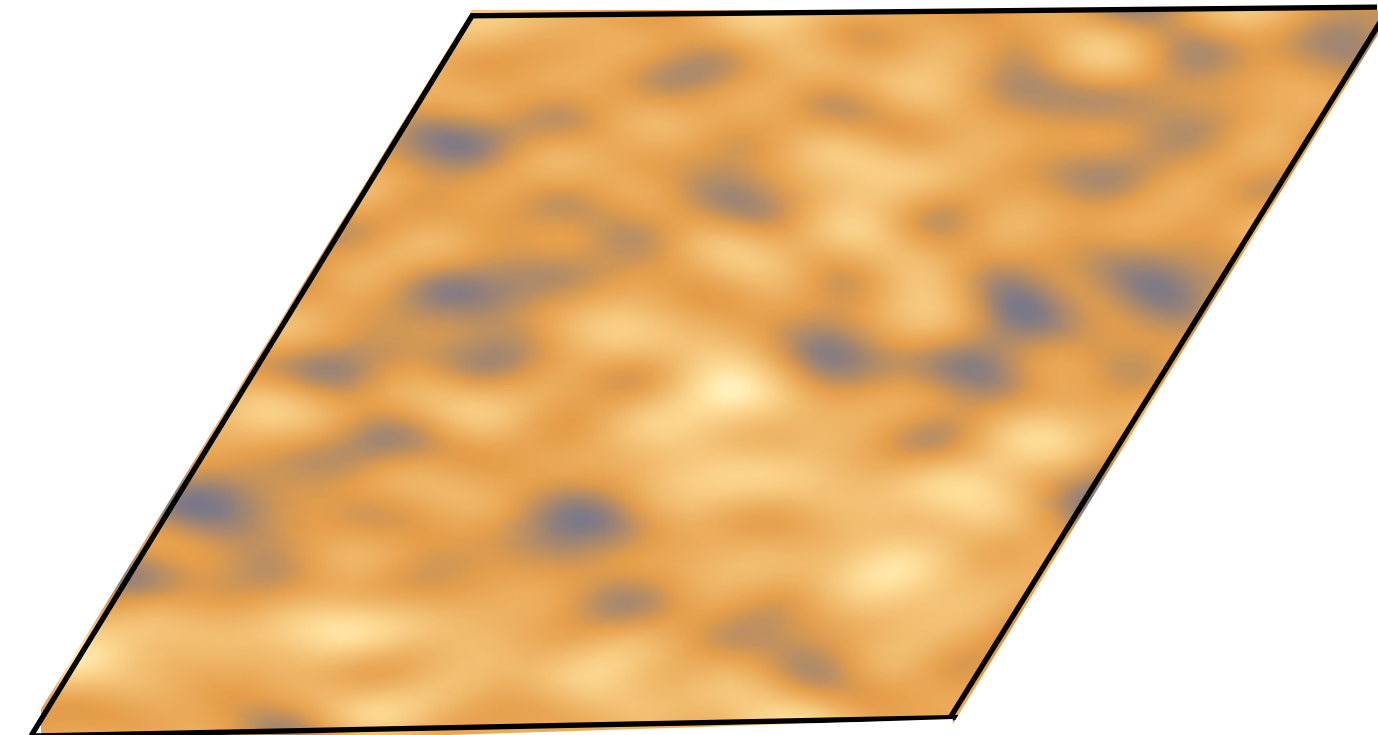
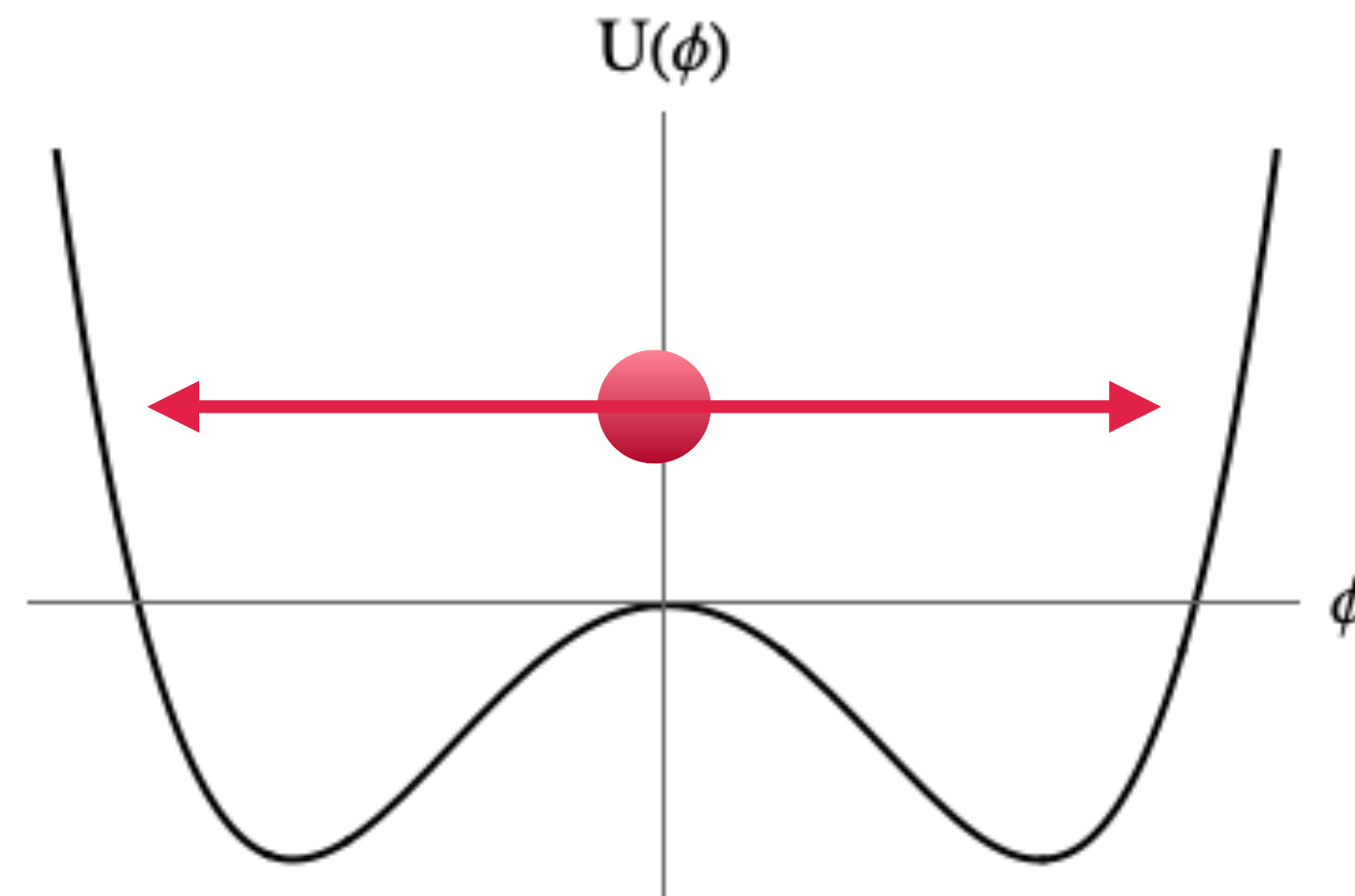
$$Z_2 : \phi \rightarrow -\phi$$

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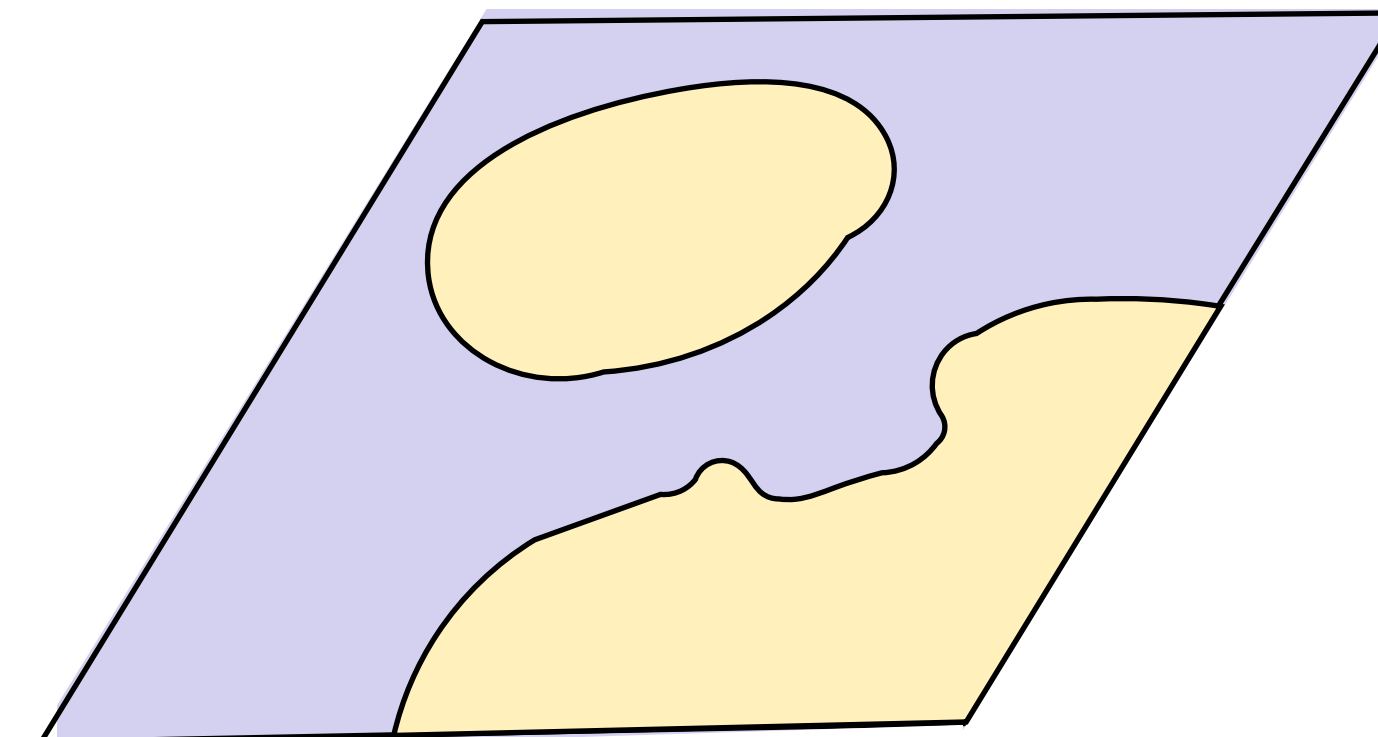
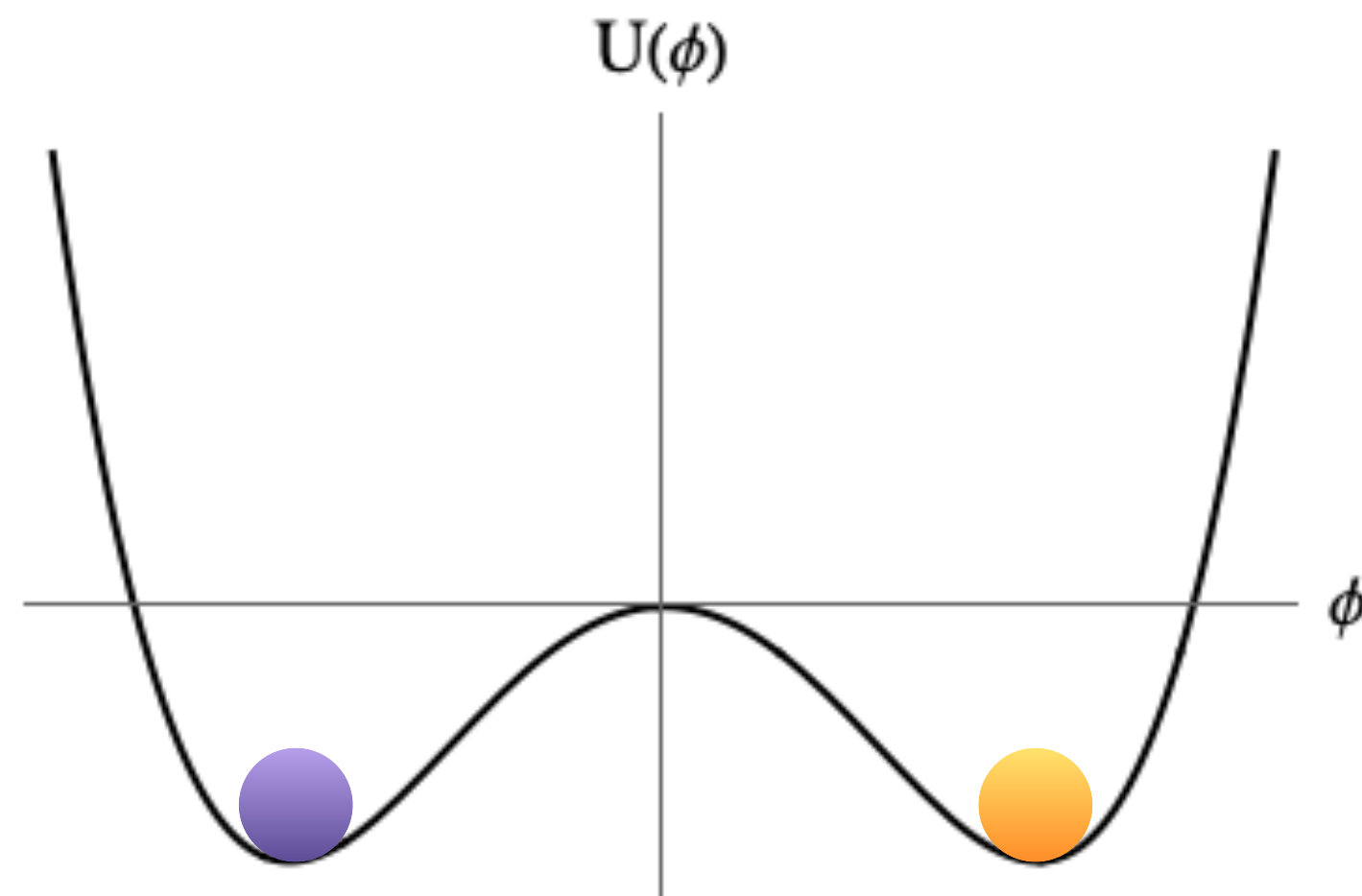
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$$\phi_{\min} = \pm \frac{f_a}{\sqrt{2\lambda}} \neq -\phi_{\min}$$

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High $T \gg f_a$



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Spontaneous symmetry breaking

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Symmetry:

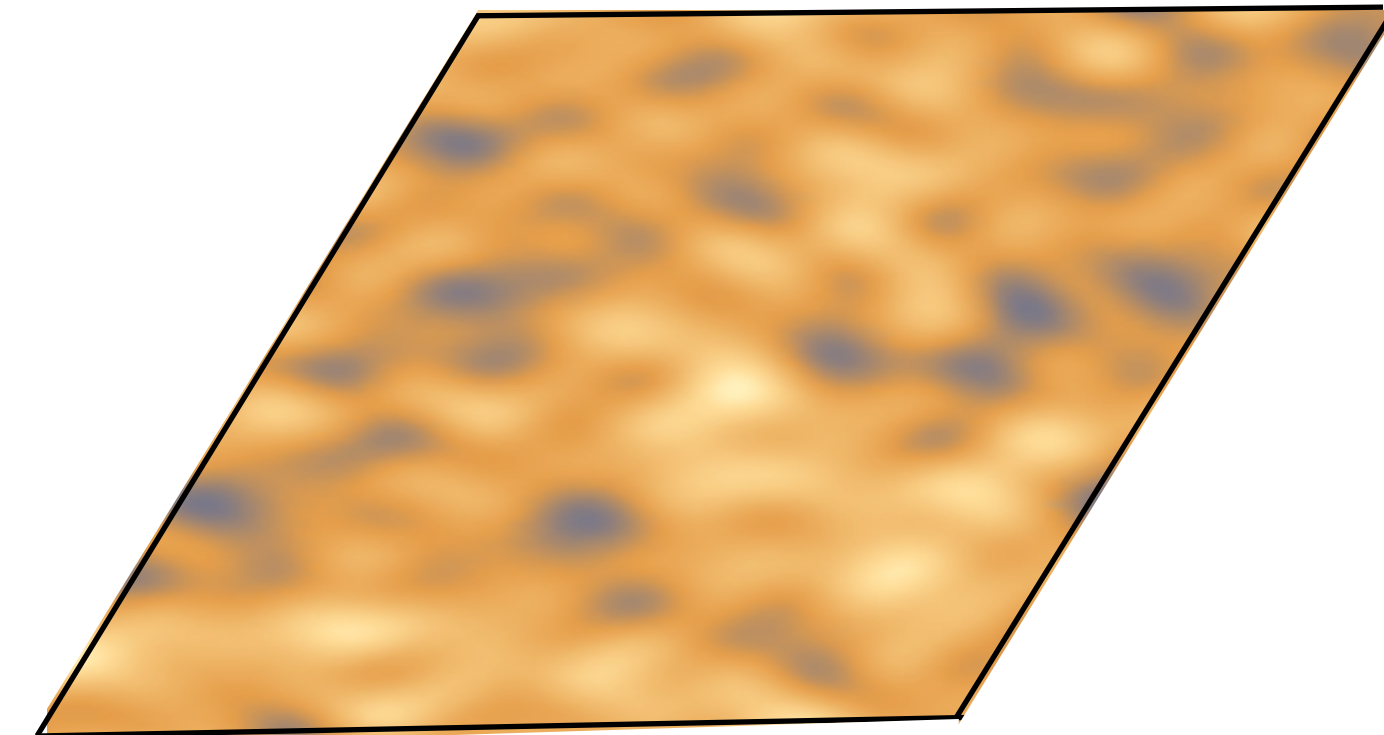
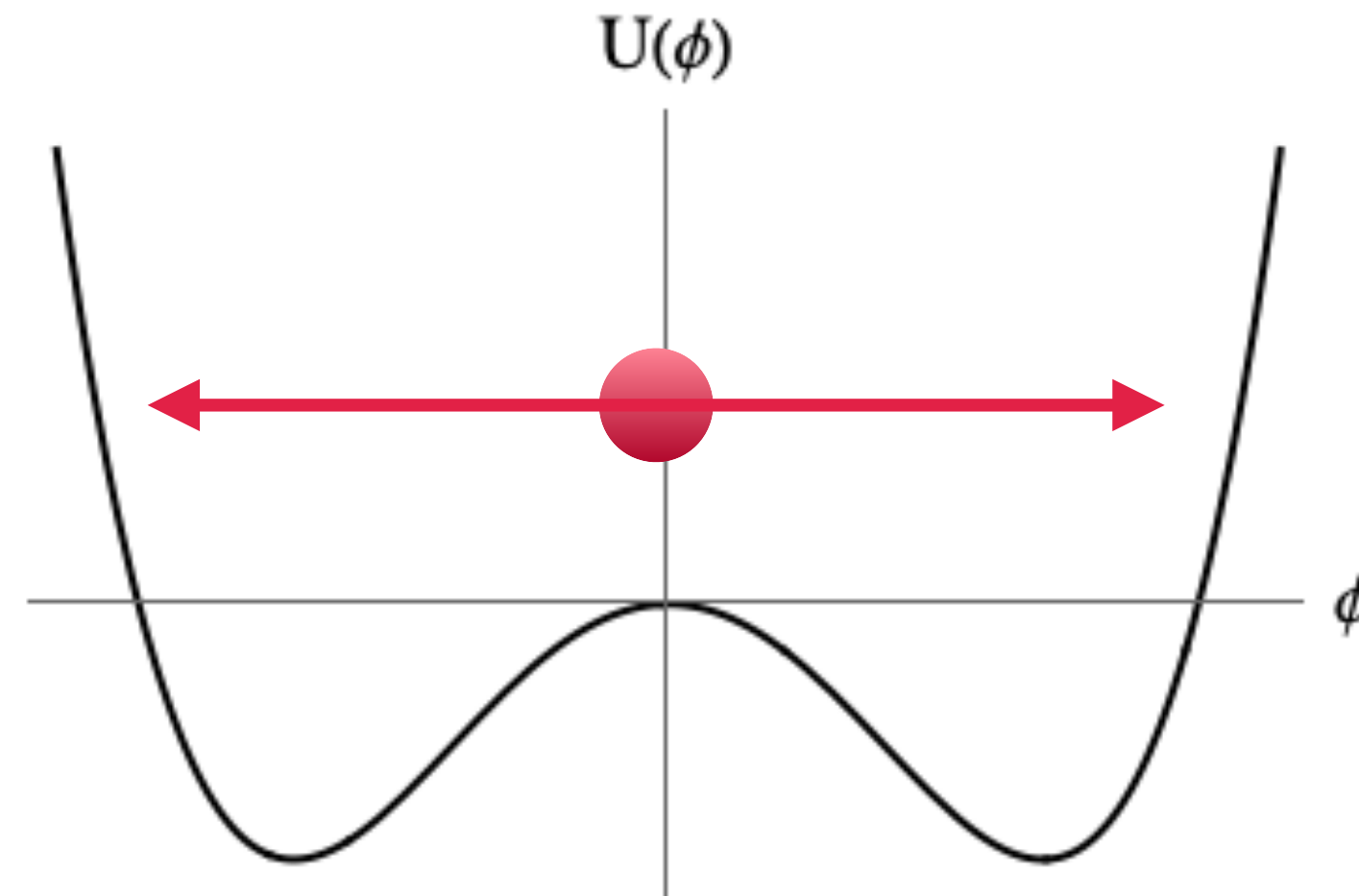
$$Z_2 : \phi \rightarrow -\phi$$

$$U(-\phi) = U(\phi)$$

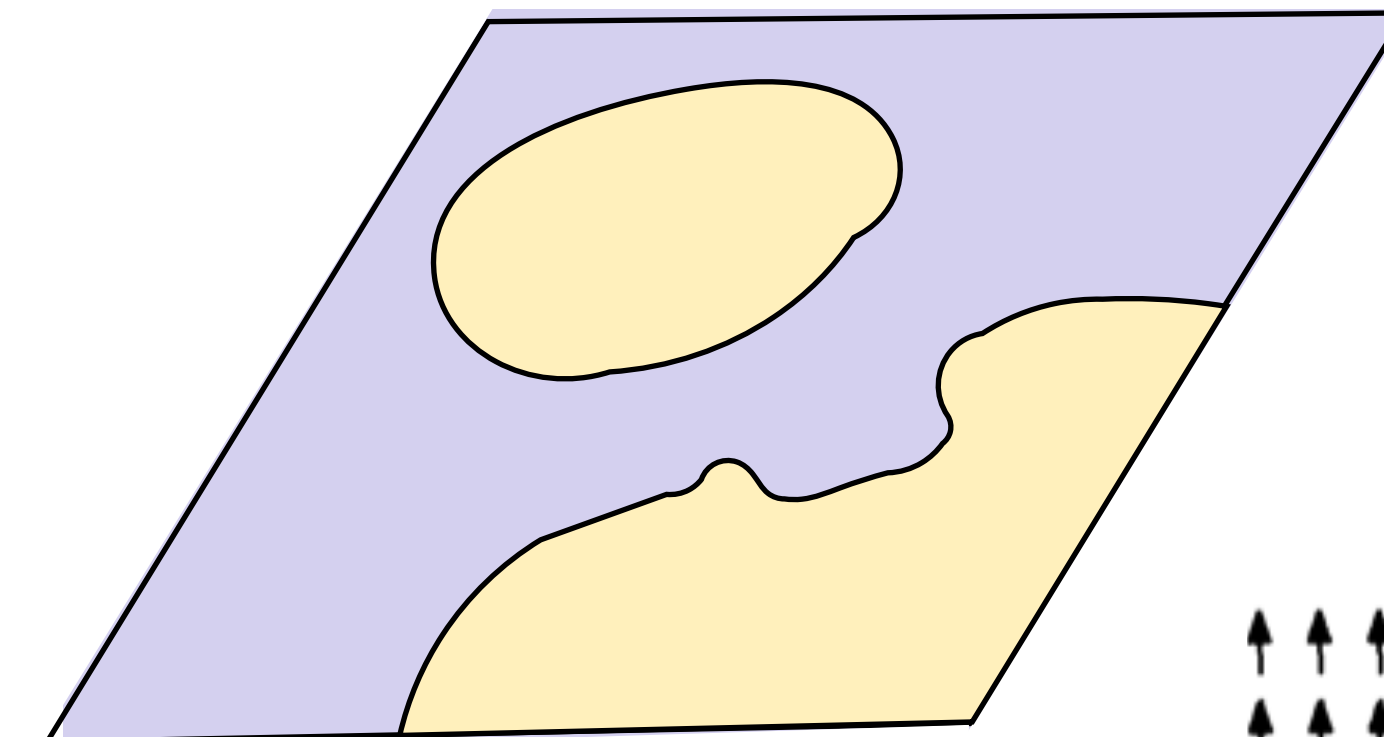
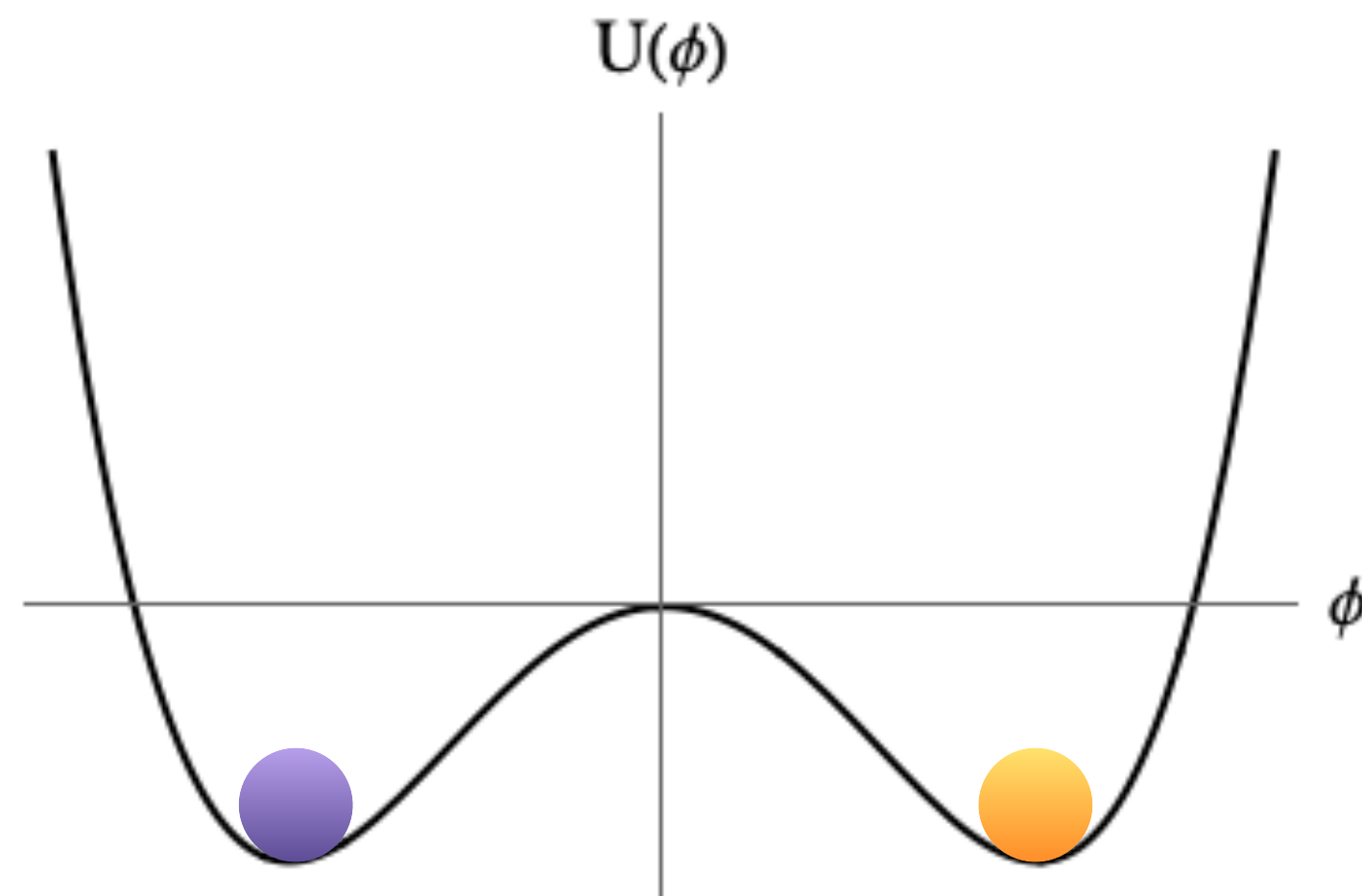
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$$\phi_{\min} = \pm \frac{f_a}{\sqrt{2\lambda}} \neq -\phi_{\min}$$

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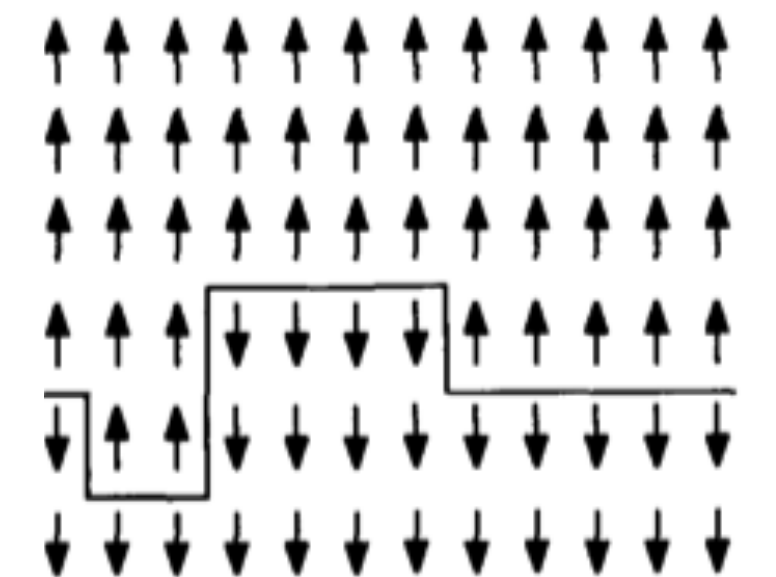


High $T \gg f_a$

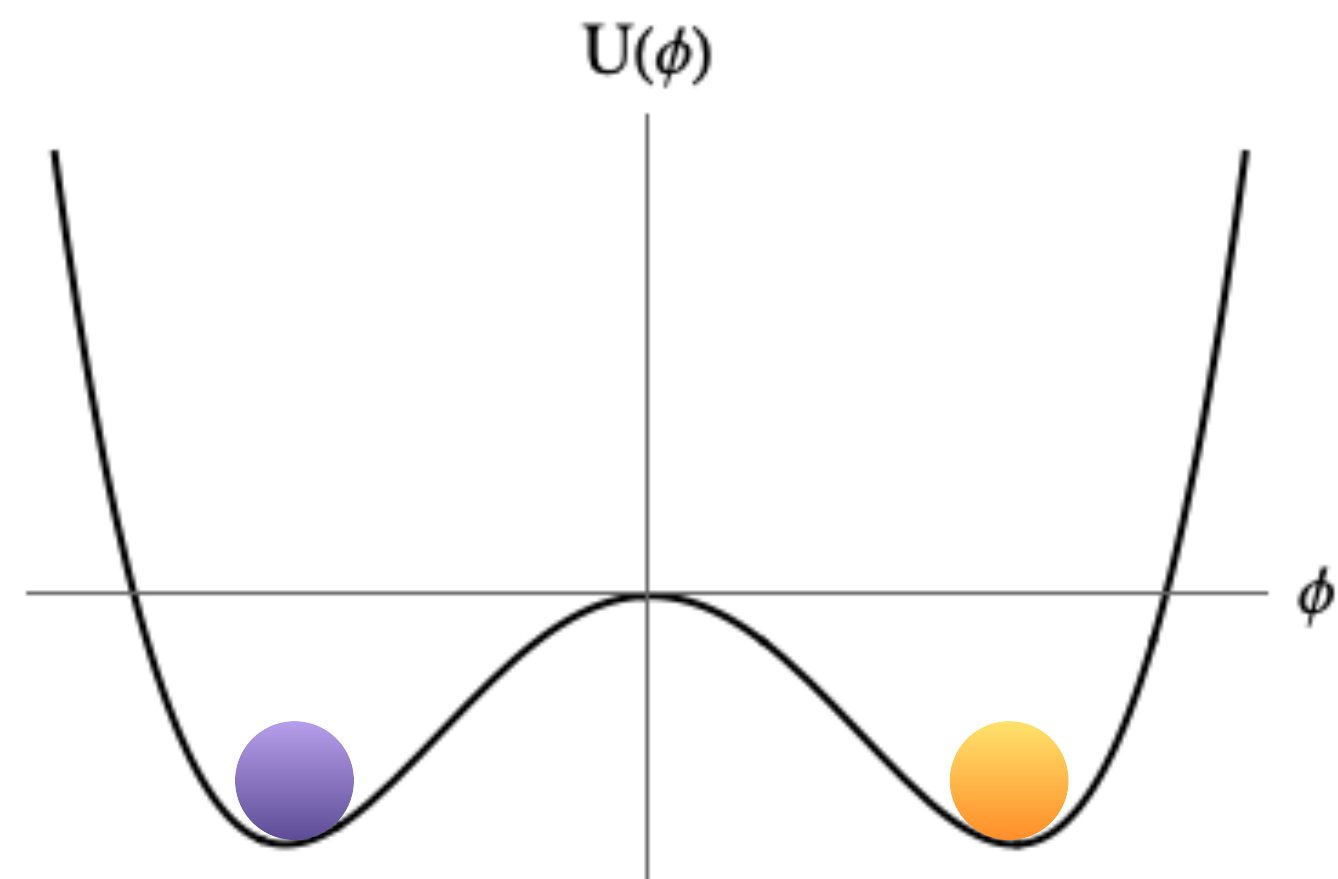
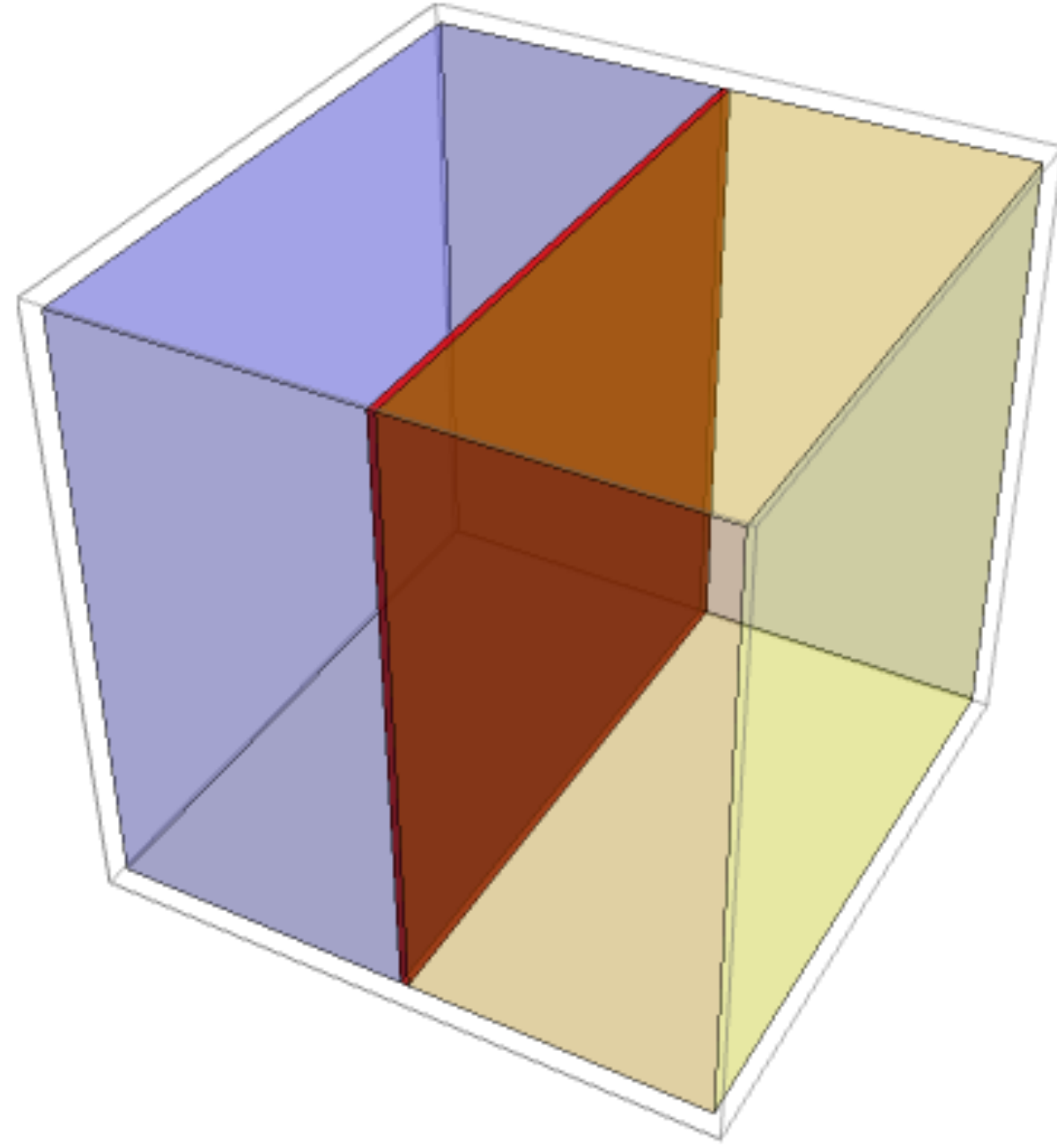


Low $T \ll f_a$

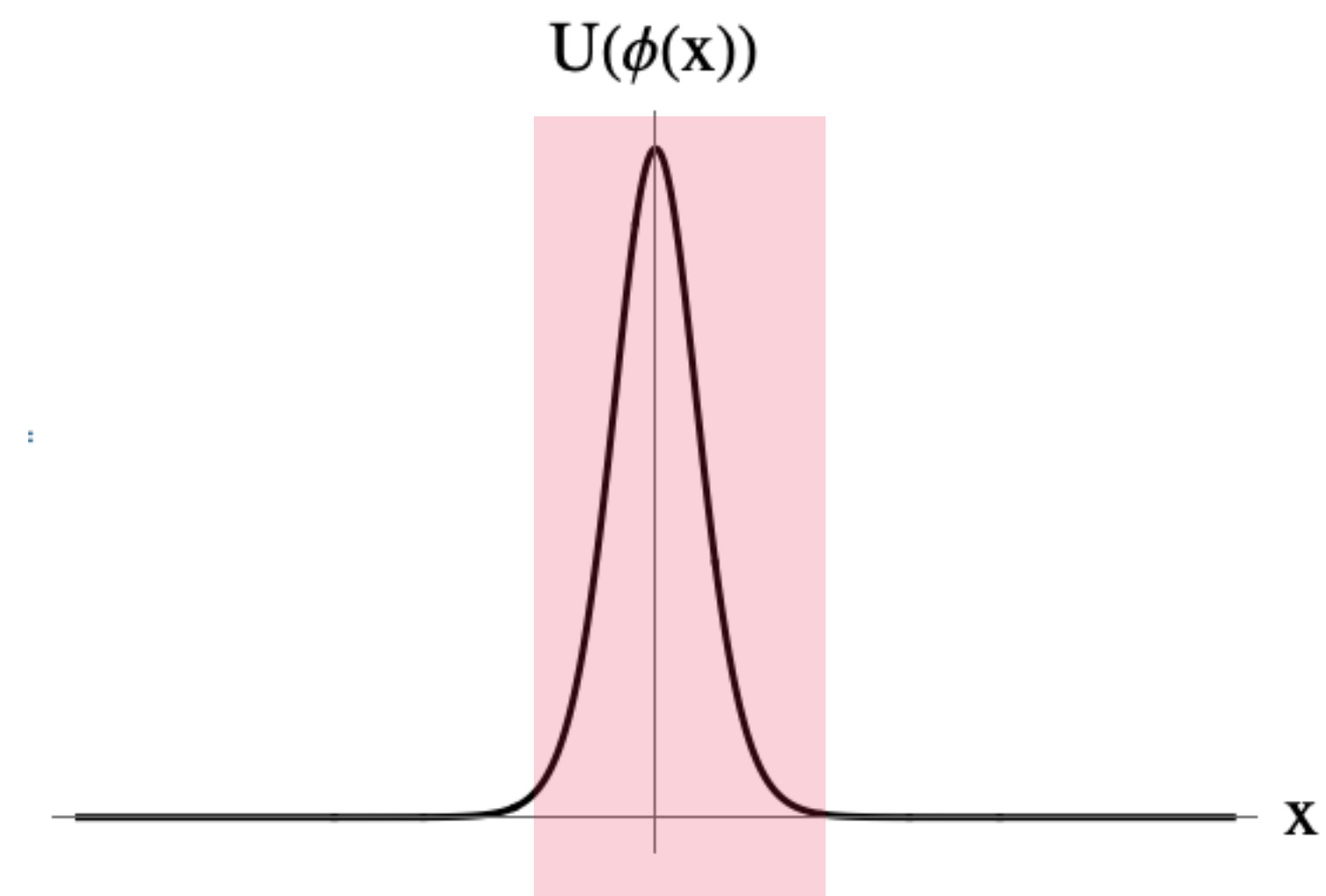
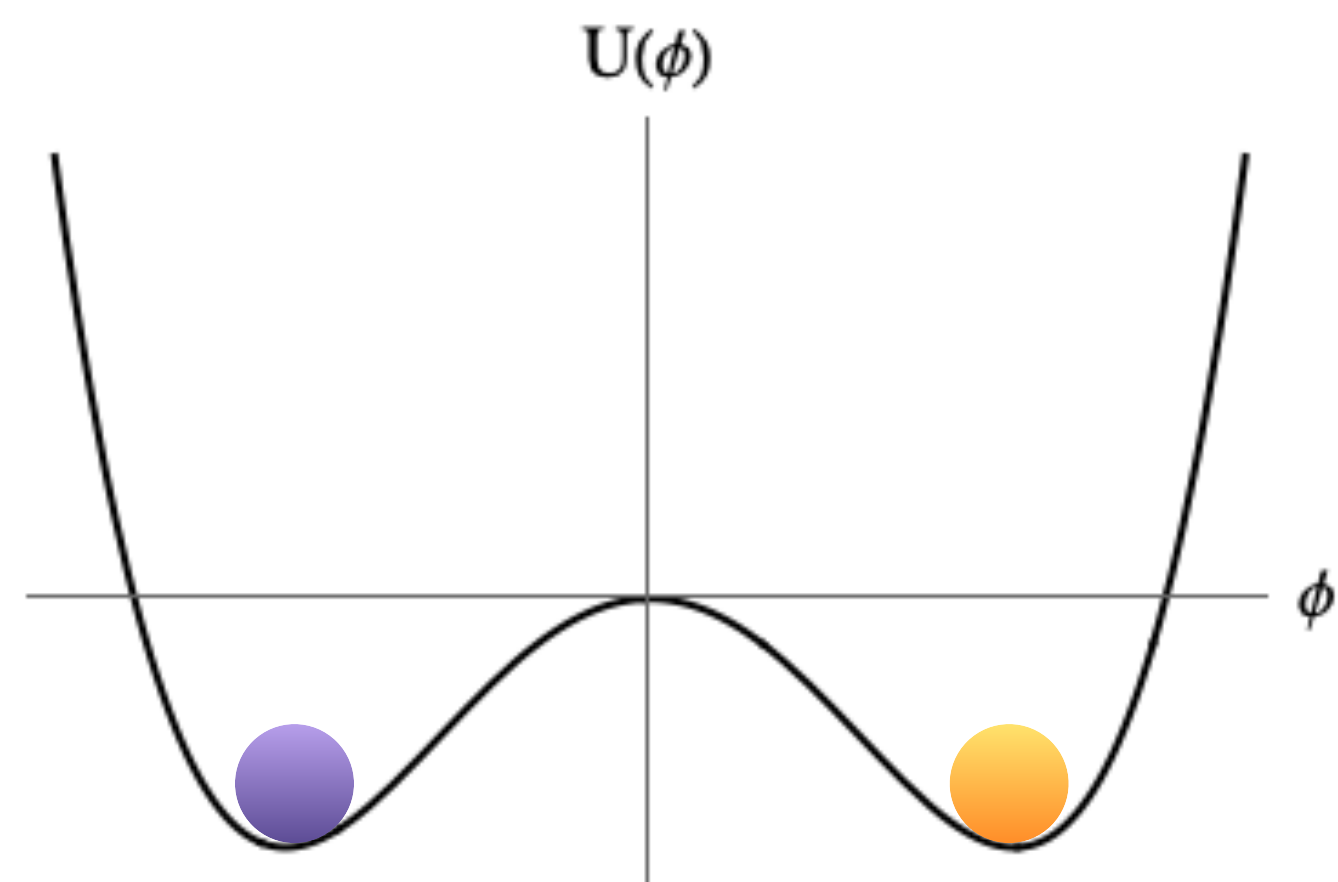
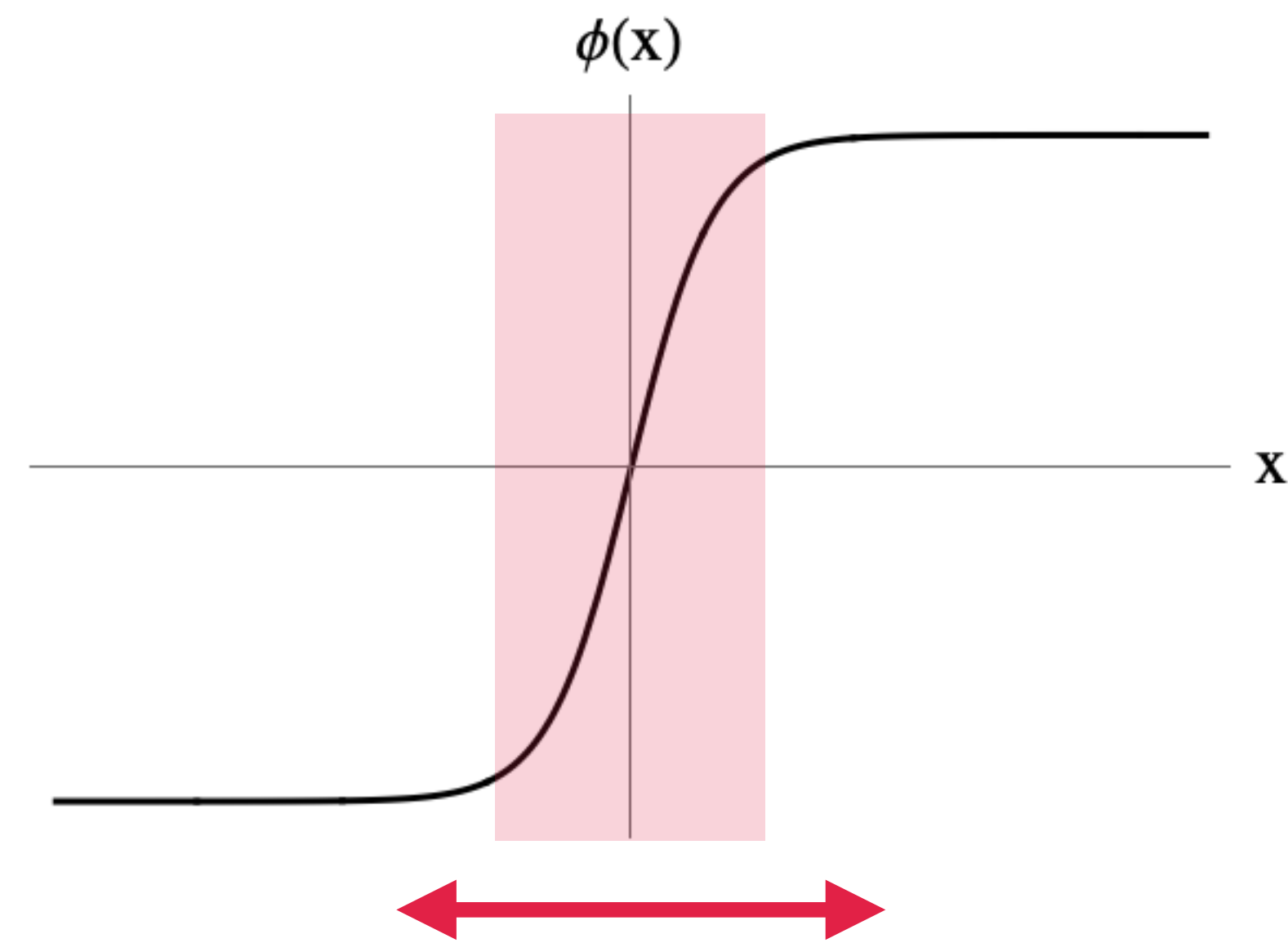
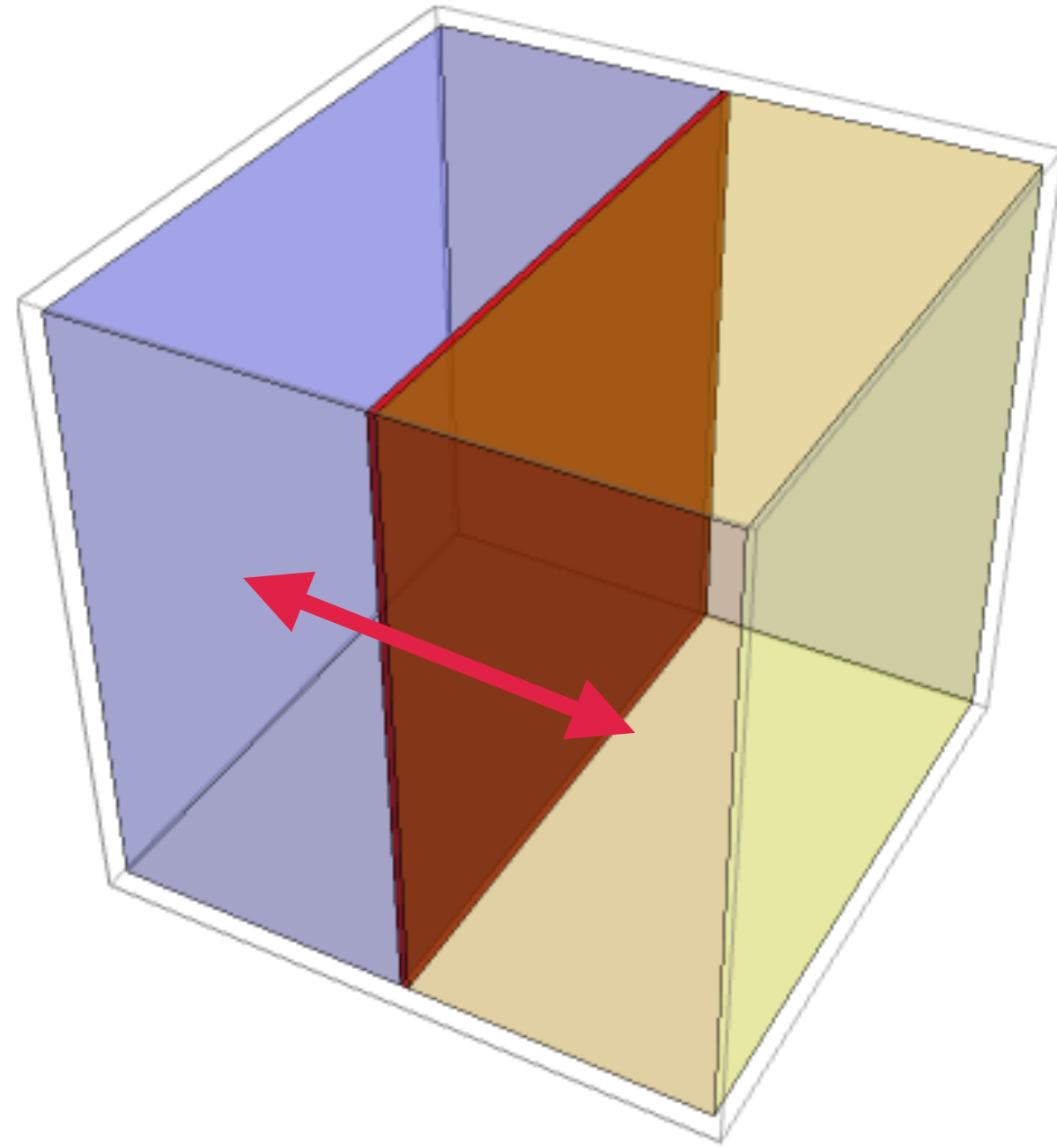
Similar to:



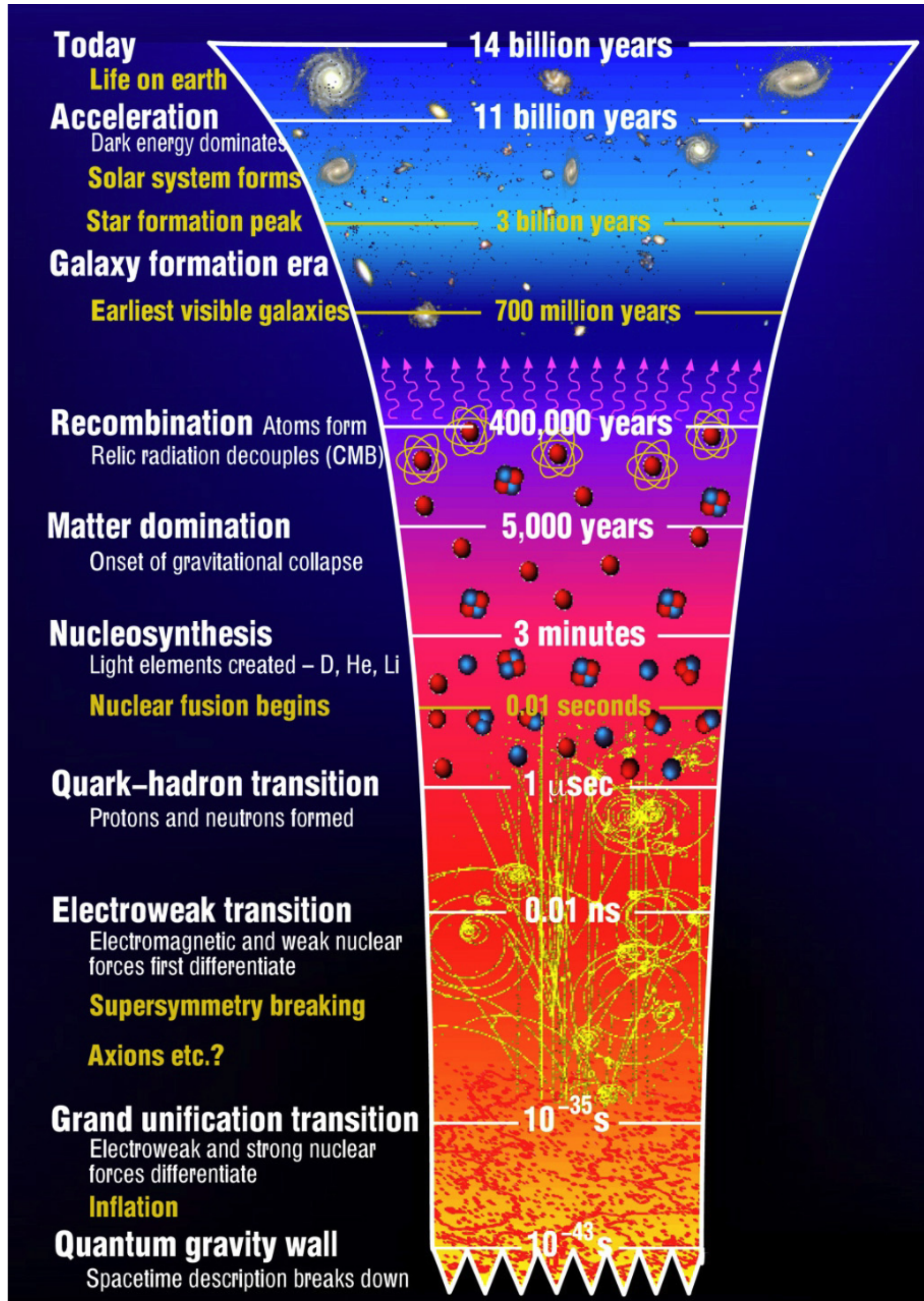
Domain walls



Domain walls



The early Universe



Temperature

10^{-4} eV

1 eV

1 MeV

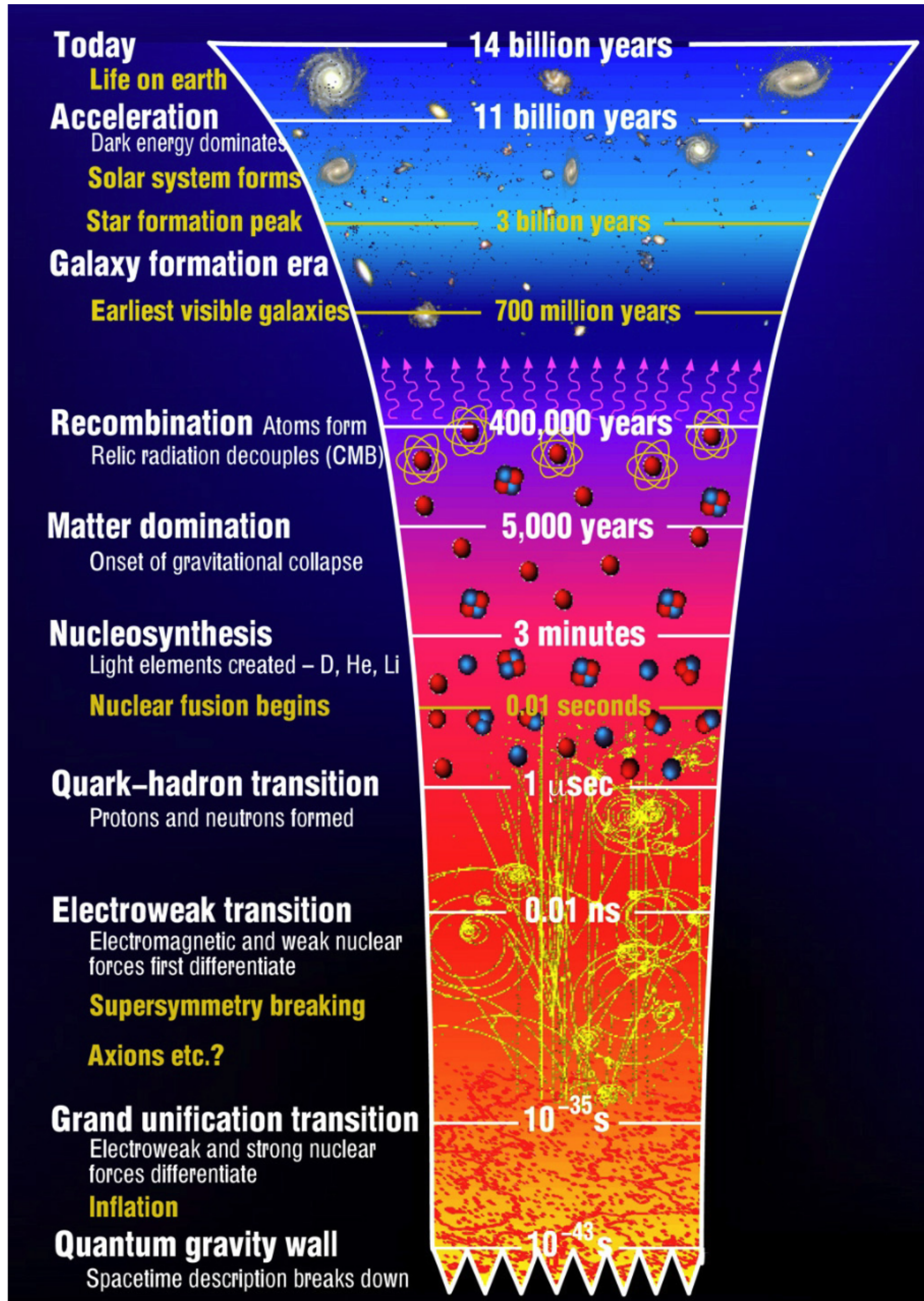
1 GeV

100 GeV
(LHC energies)

10^{16} GeV

($(\text{eV})^{-1} \sim 10^{-7} \text{ m}$ $(\text{MeV})^{-1} \sim 10^{-13} \text{ m}$ $(10^{16} \text{ GeV})^{-1} \sim 10^{-32} \text{ m}$)

The early Universe



Temperature

10^{-4} eV

1 eV

1 MeV

1 GeV

100 m

100 GeV

(LHC energies)

10^{16} GeV

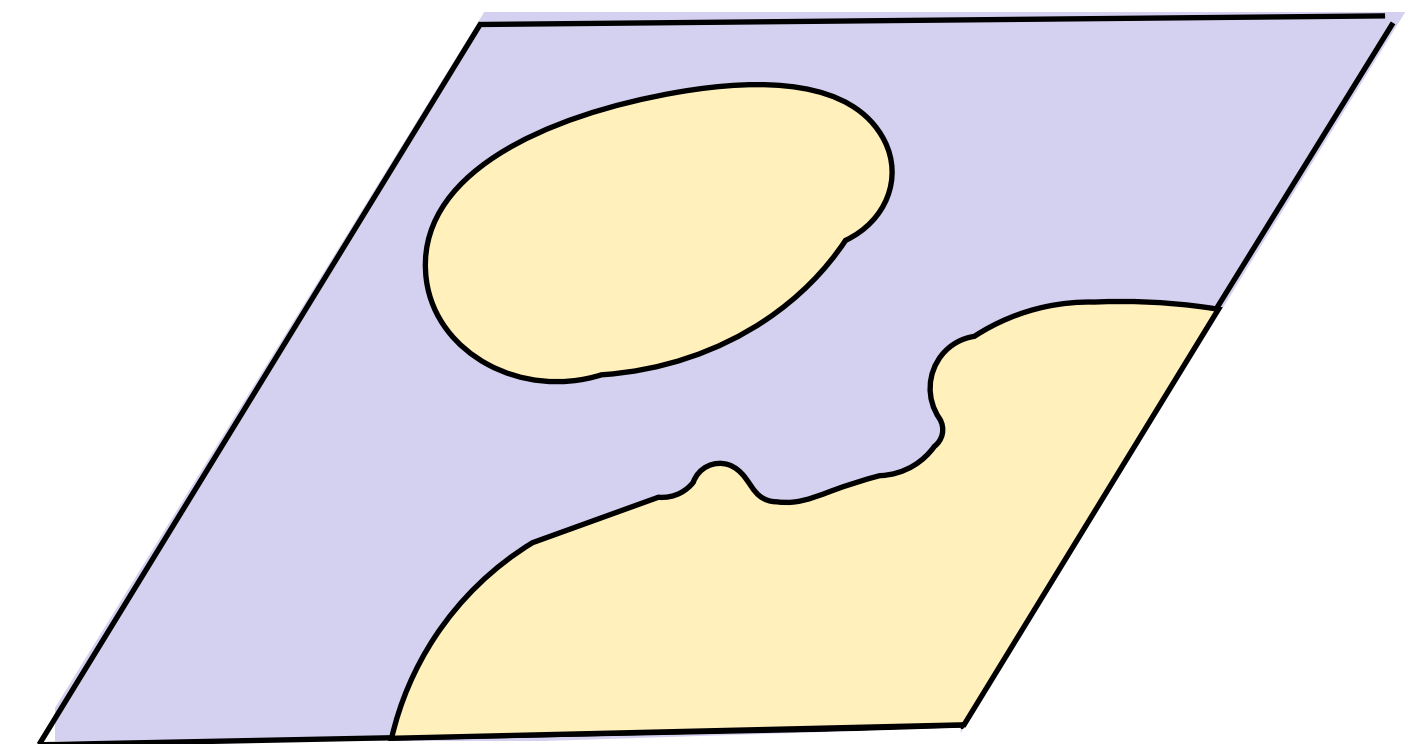
f_a

10^{-30} m

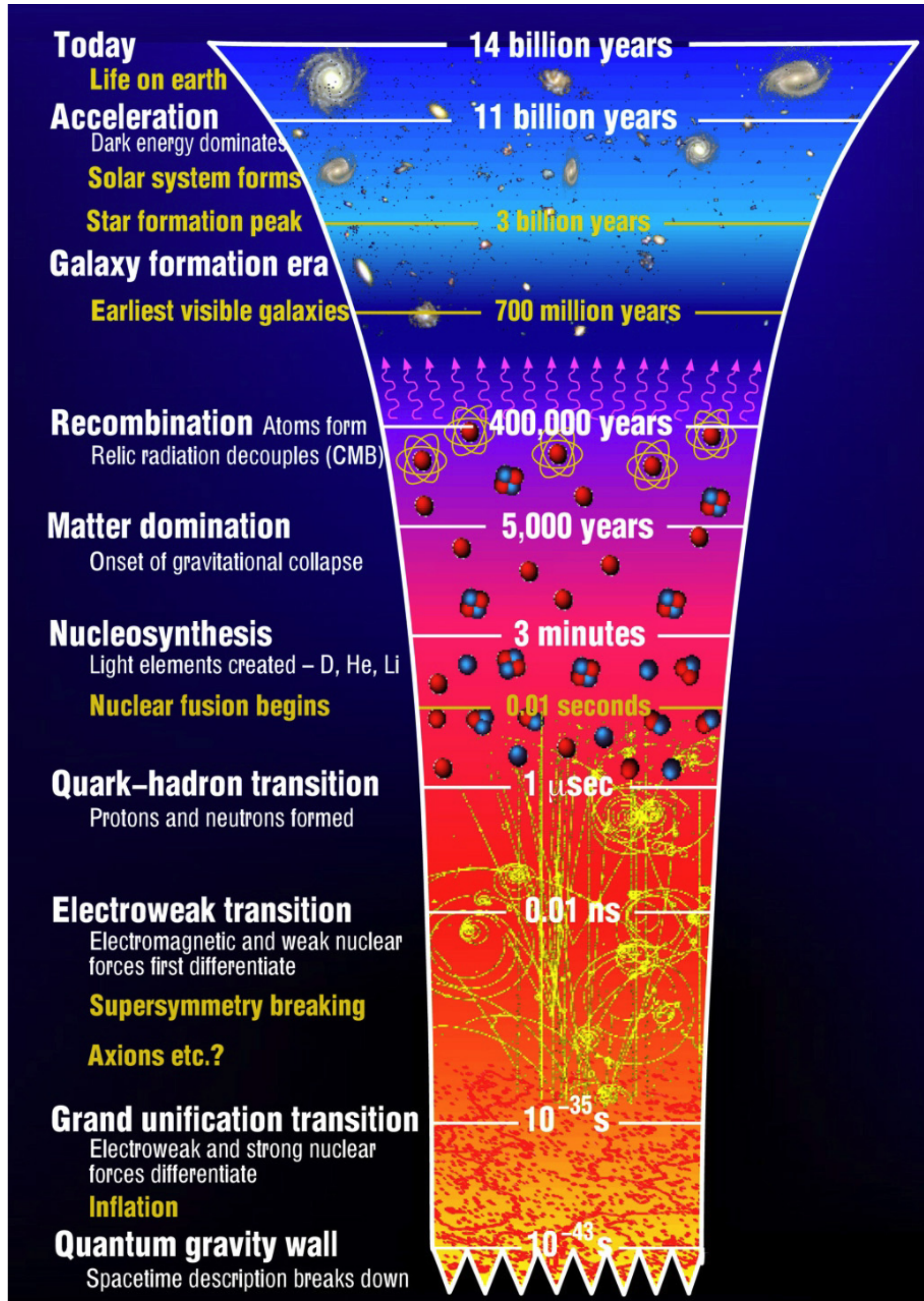
$(\text{eV})^{-1} \sim 10^{-7}$ m

$(\text{MeV})^{-1} \sim 10^{-13}$ m

$(10^{16} \text{ GeV})^{-1} \sim 10^{-32}$ m



The early Universe



Temperature

10^{-4} eV

1 eV

1 MeV

1 GeV

100 GeV

10^{16} GeV

((eV)⁻¹ ~ 10⁻⁷ m

Hubble horizon:

10^{26} m

100 m

(LHC energies)

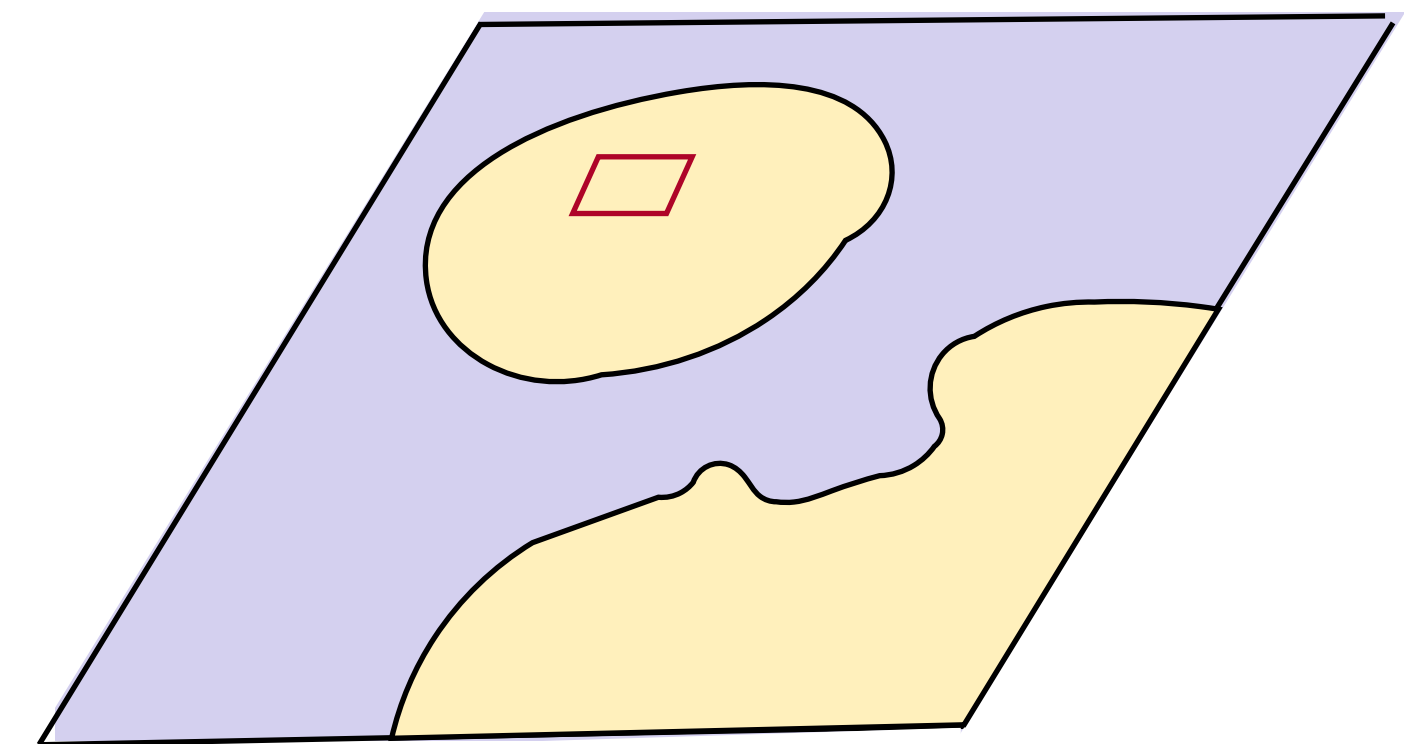
10^{-30} m

(MeV)⁻¹ ~ 10⁻¹³ m

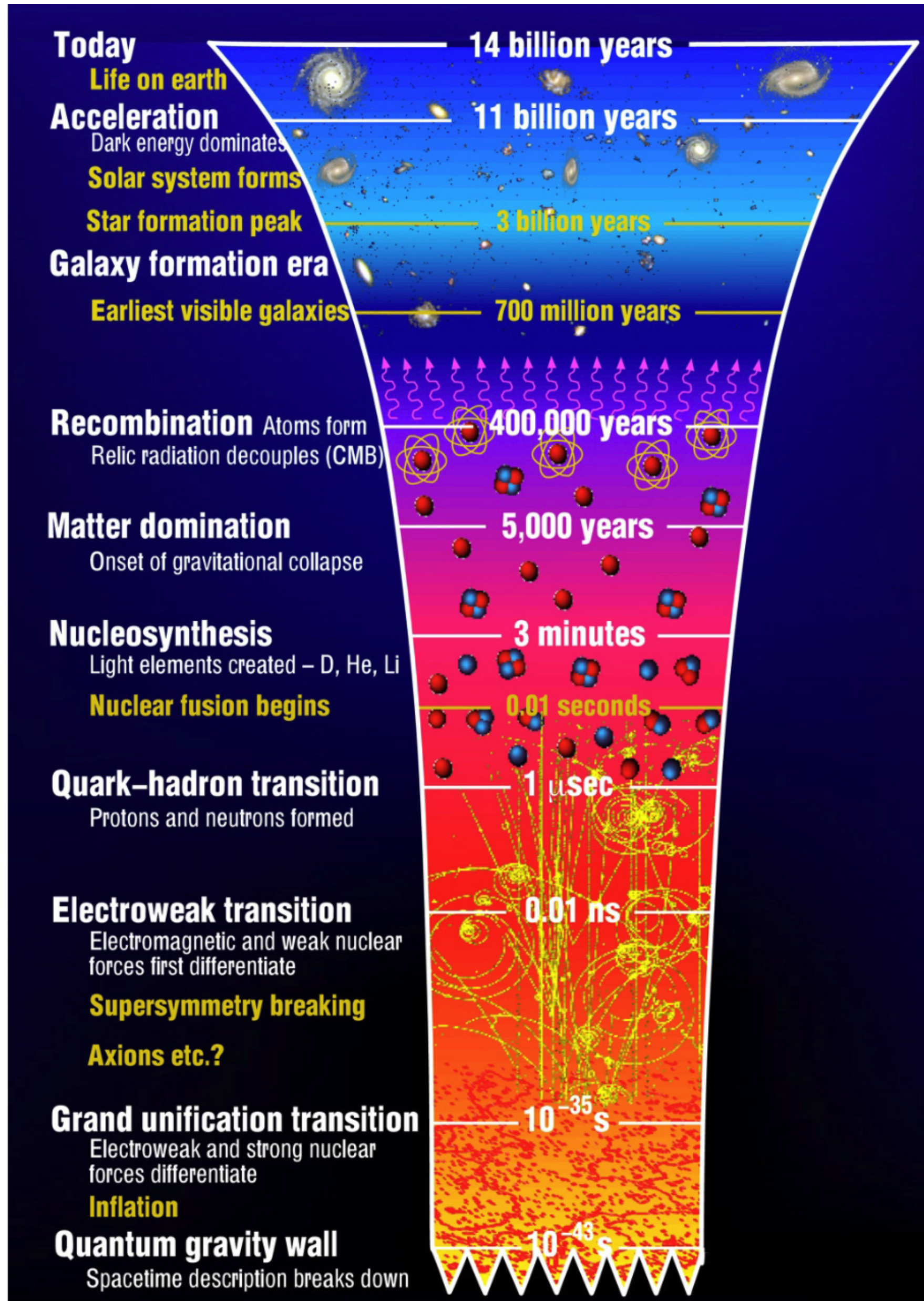
(10^{16} GeV)⁻¹ ~ 10⁻³² m)

$$H^{-1} \simeq \frac{M_{\text{Pl}}}{\sqrt{\rho}} \simeq \frac{M_{\text{Pl}}}{T^2}$$

(during radiation domination)



The early Universe



Temperature

10⁻⁴ eV

1 eV

1 MeV

1 GeV

100 GeV

(LHC energies)

10¹⁶ GeV

f_a

Hubble horizon:

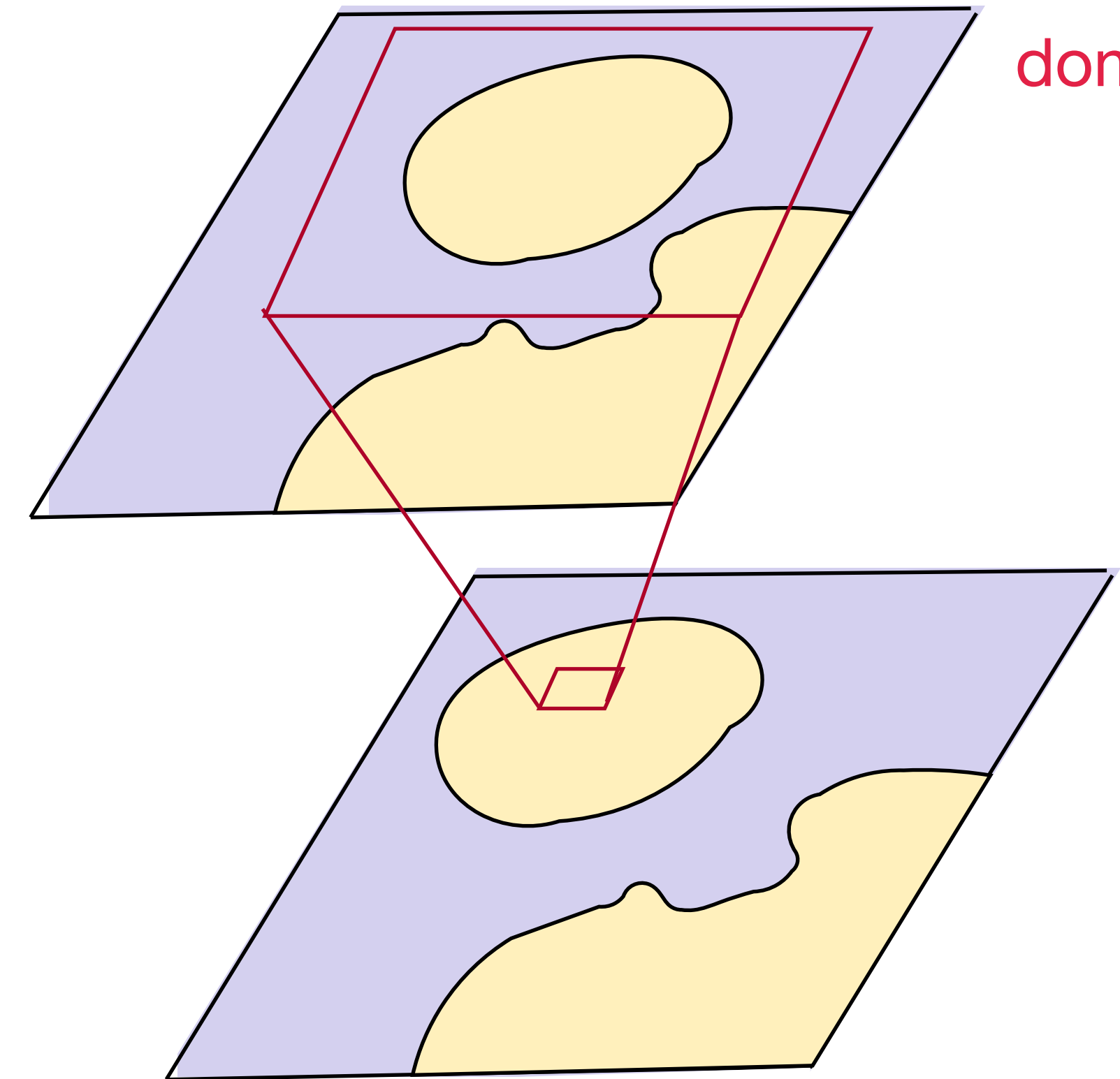
10²⁶ m

100 m

10⁻³⁰ m

$$H^{-1} \simeq \frac{M_{\text{Pl}}}{\sqrt{\rho}} \simeq \frac{M_{\text{Pl}}}{T^2}$$

(during radiation domination)



((eV)⁻¹ ~ 10⁻⁷ m

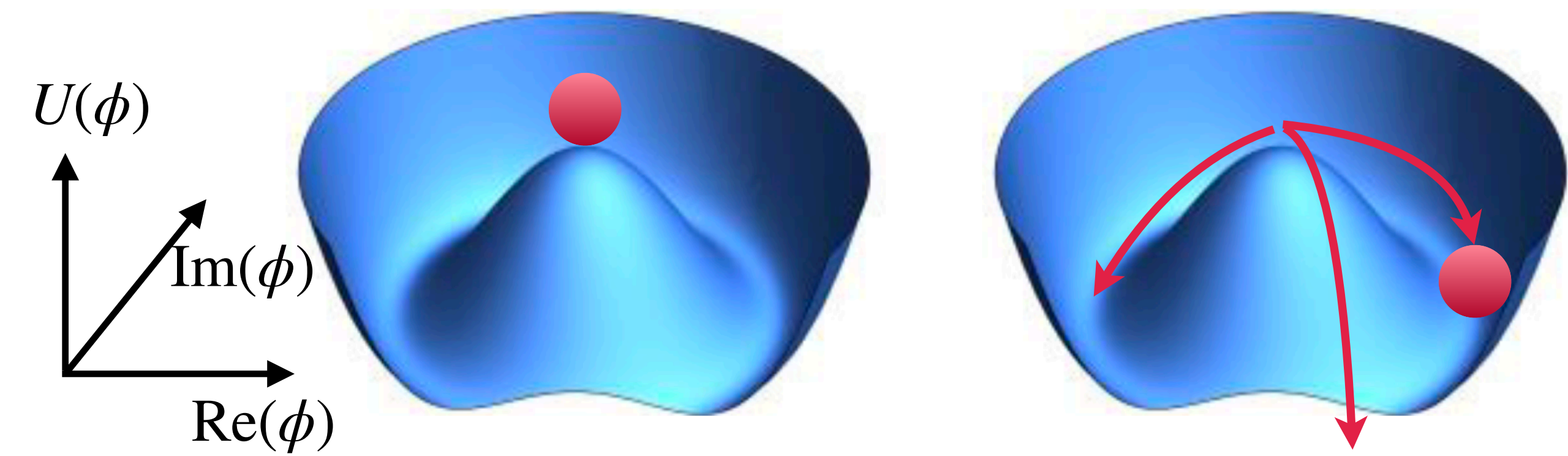
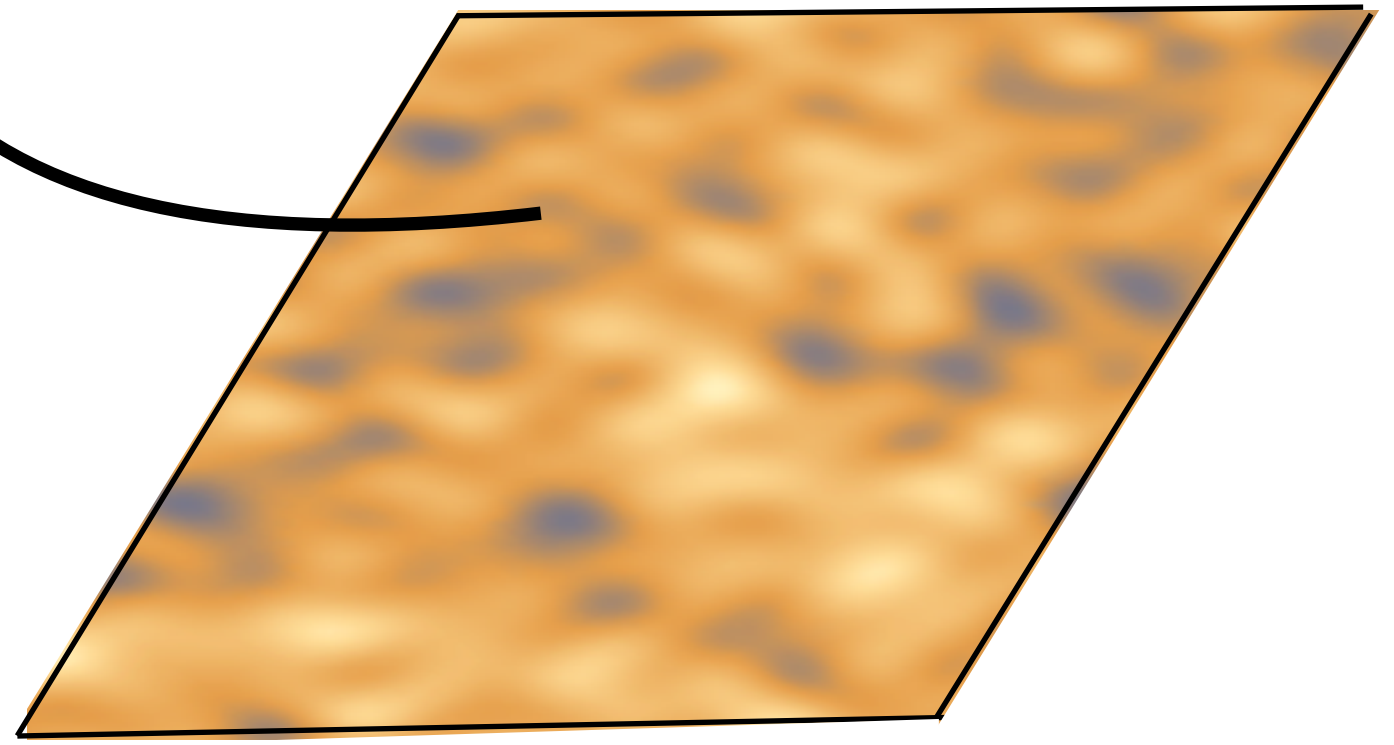
(MeV)⁻¹ ~ 10⁻¹³ m

(10¹⁶ GeV)⁻¹ ~ 10⁻³² m)

Strings

$$U(\phi) = (|\phi|^2 - f_a^2)^2$$

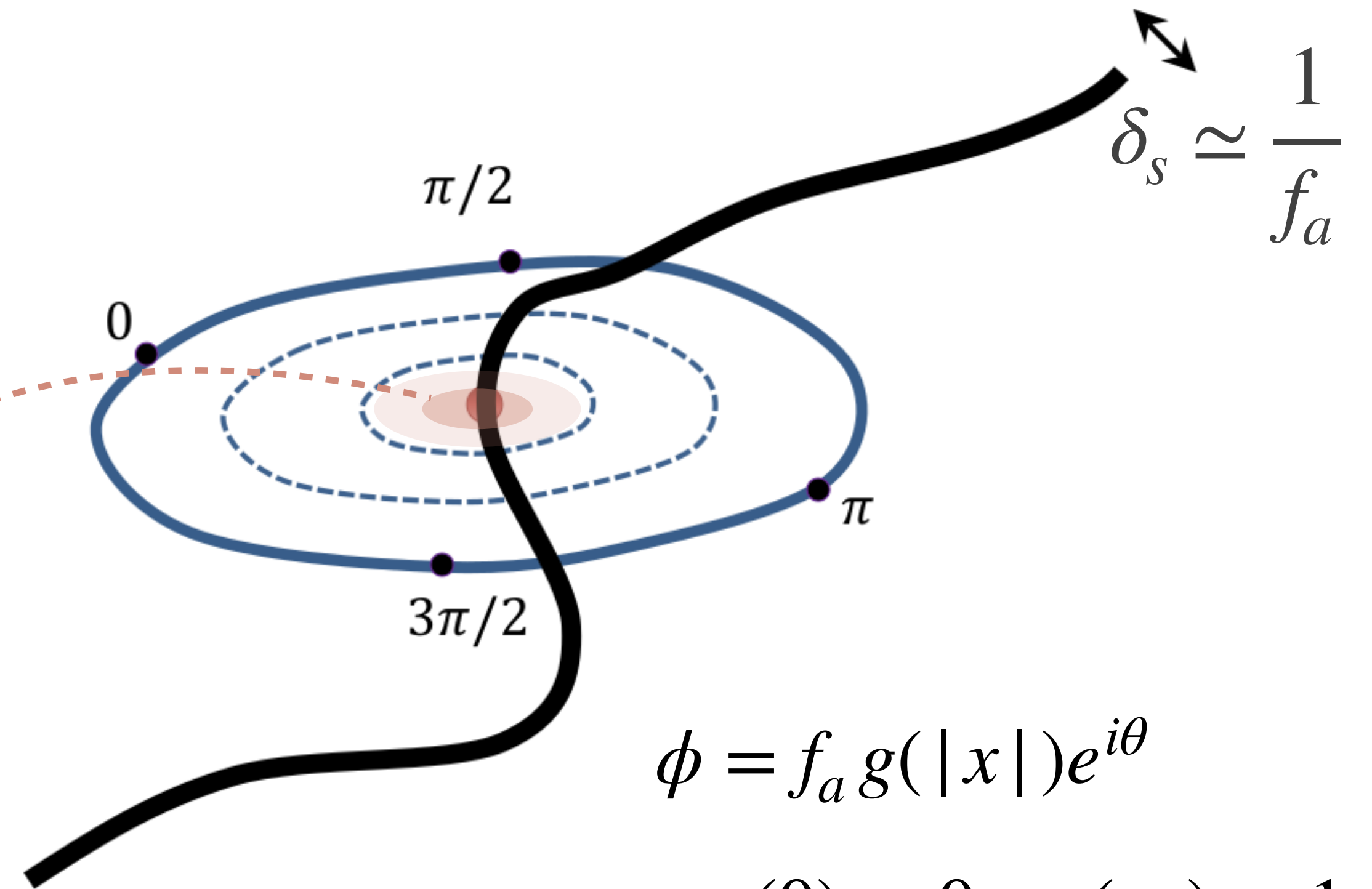
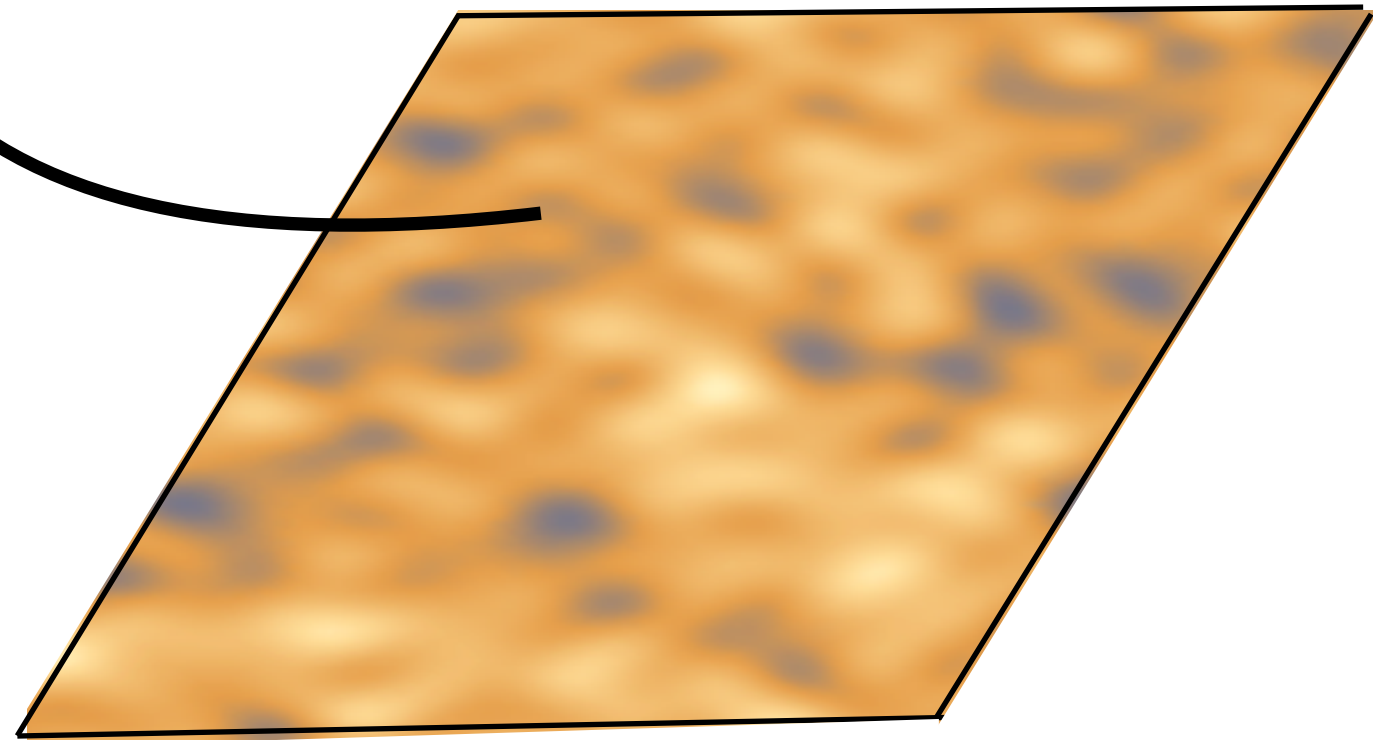
$$\phi(t, \vec{x}) \in \mathbb{C}$$



Strings

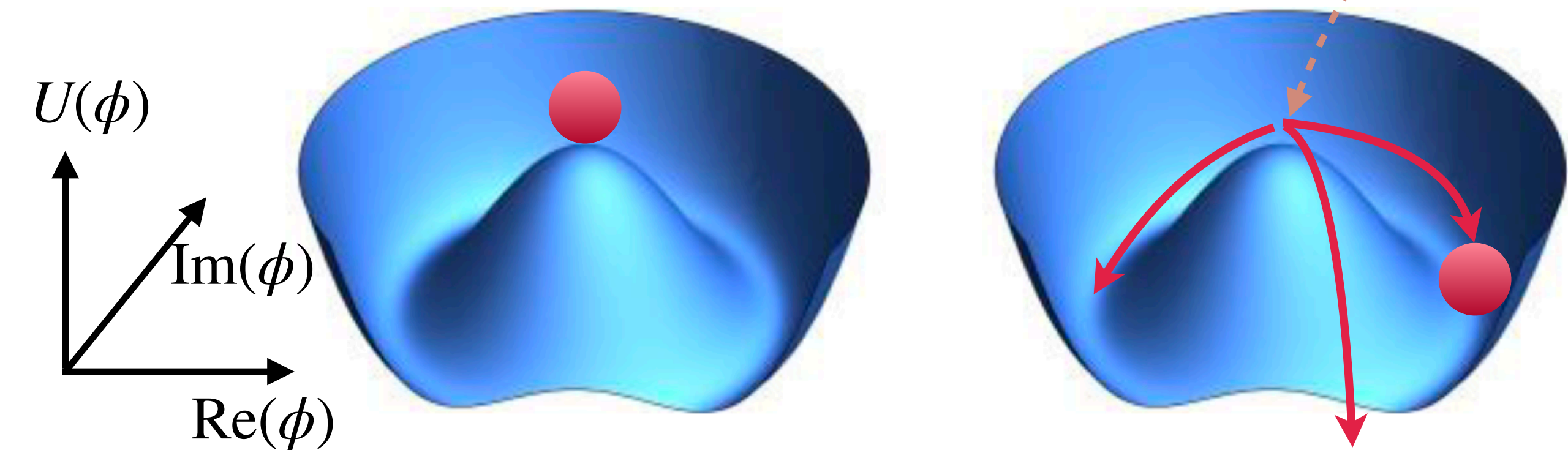
$$U(\phi) = (|\phi|^2 - f_a^2)^2$$

$$\phi(t, \vec{x}) \in \mathbb{C}$$



$$\phi = f_a g(|x|) e^{i\theta}$$

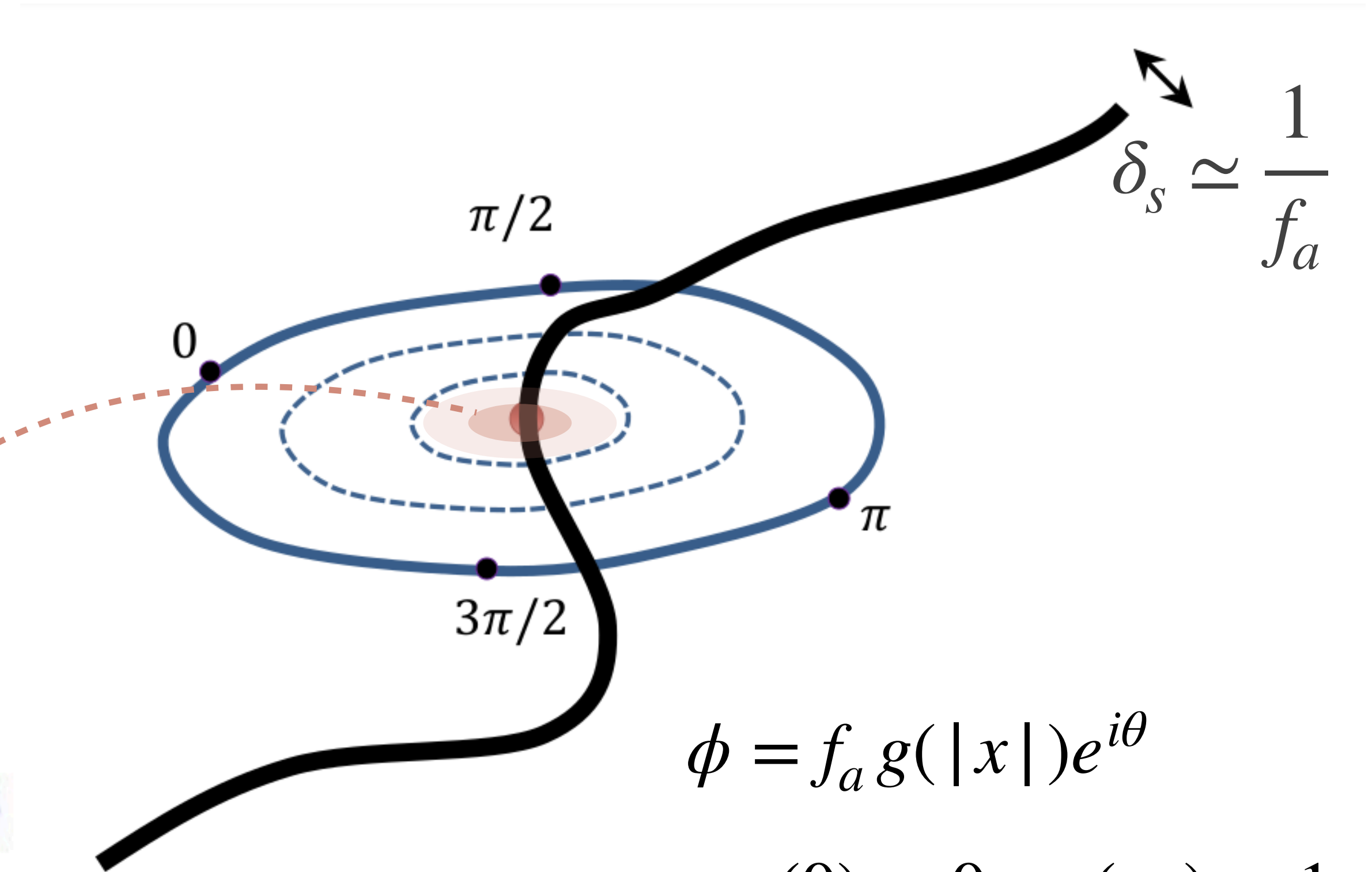
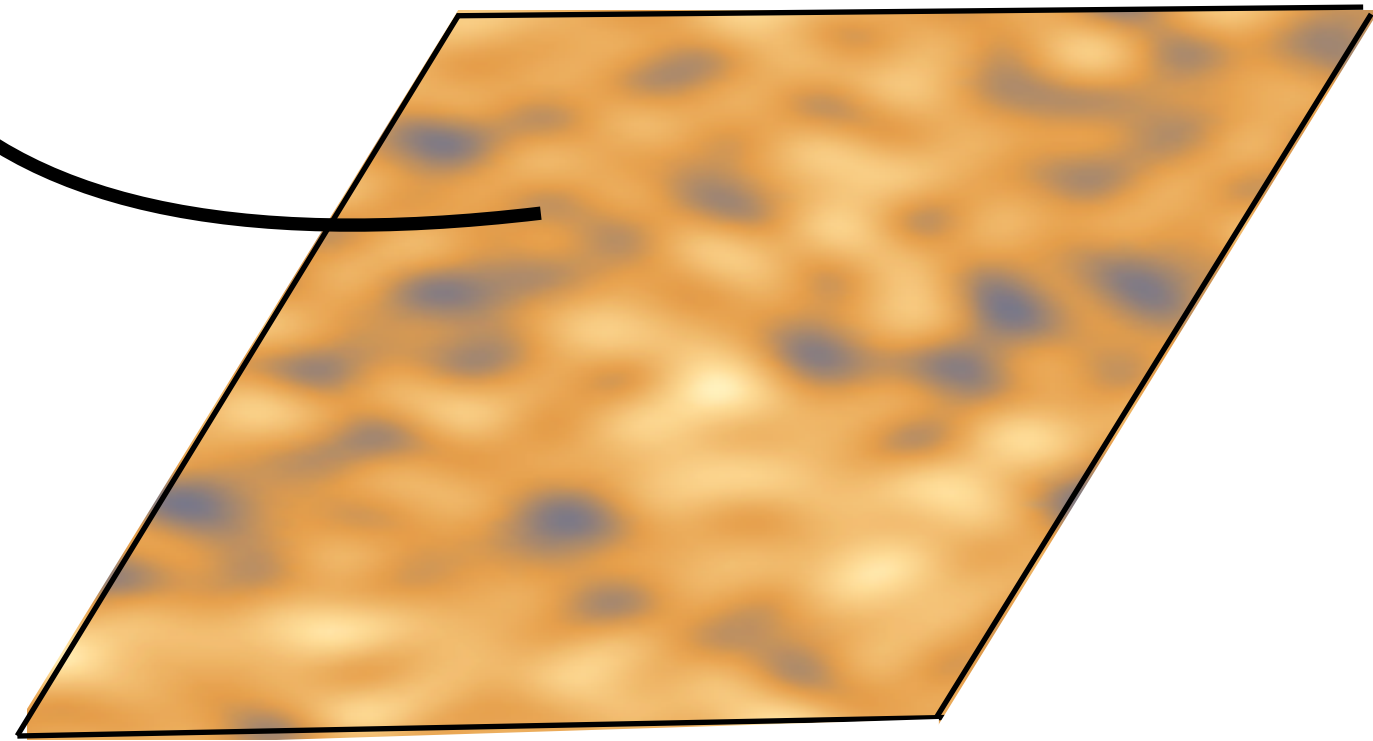
$$g(0) = 0 \quad g(\infty) = 1$$



Strings

$$U(\phi) = (|\phi|^2 - f_a^2)^2$$

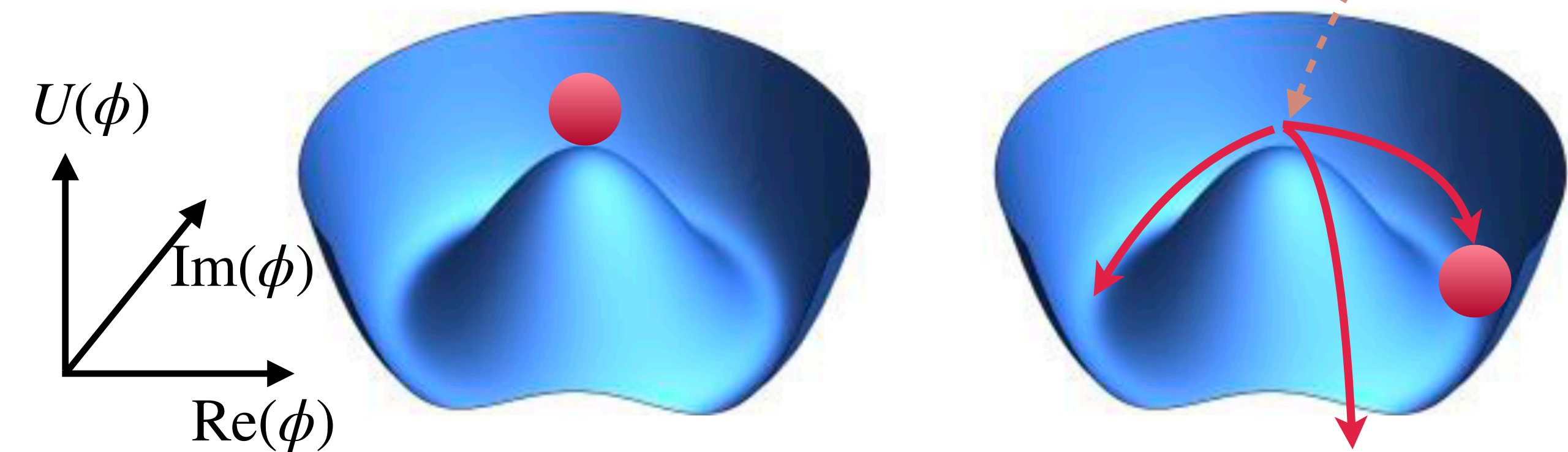
$$\phi(t, \vec{x}) \in \mathbb{C}$$



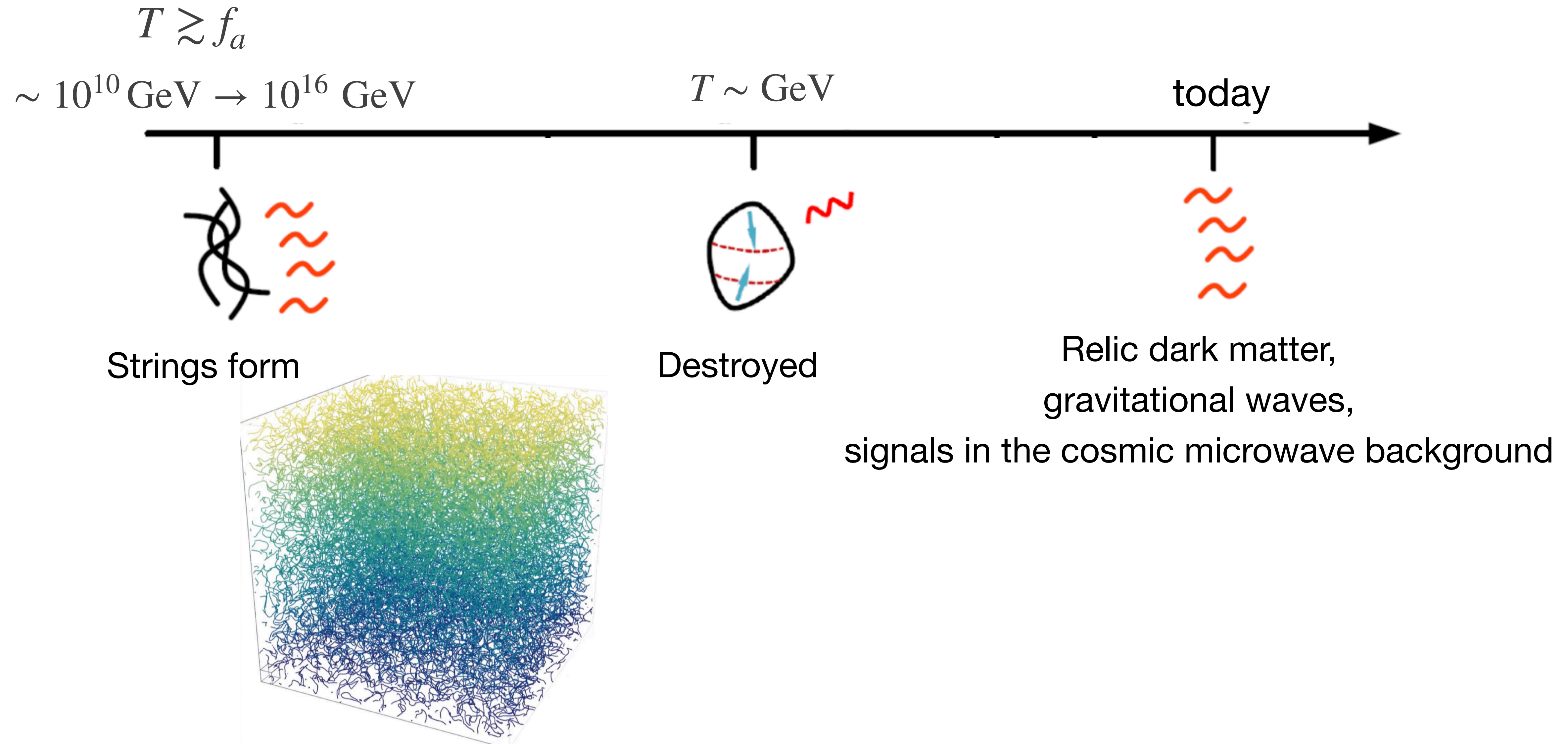
$$\phi = f_a g(|x|) e^{i\theta}$$

$$g(0) = 0 \quad g(\infty) = 1$$

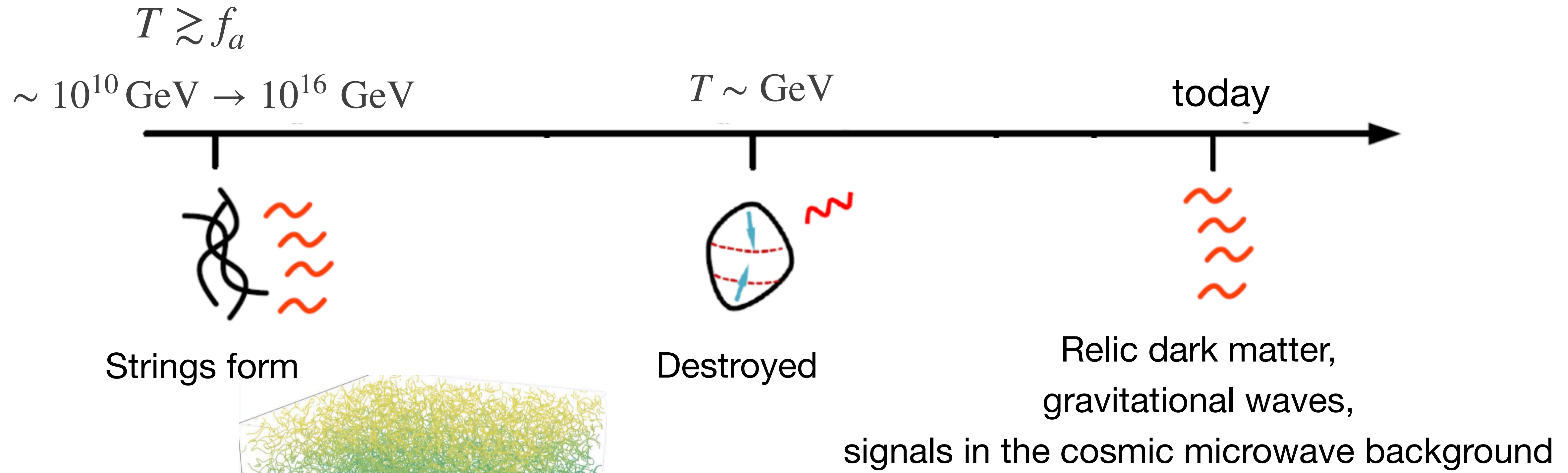
String tension $E/L \sim \pi f_a^2$



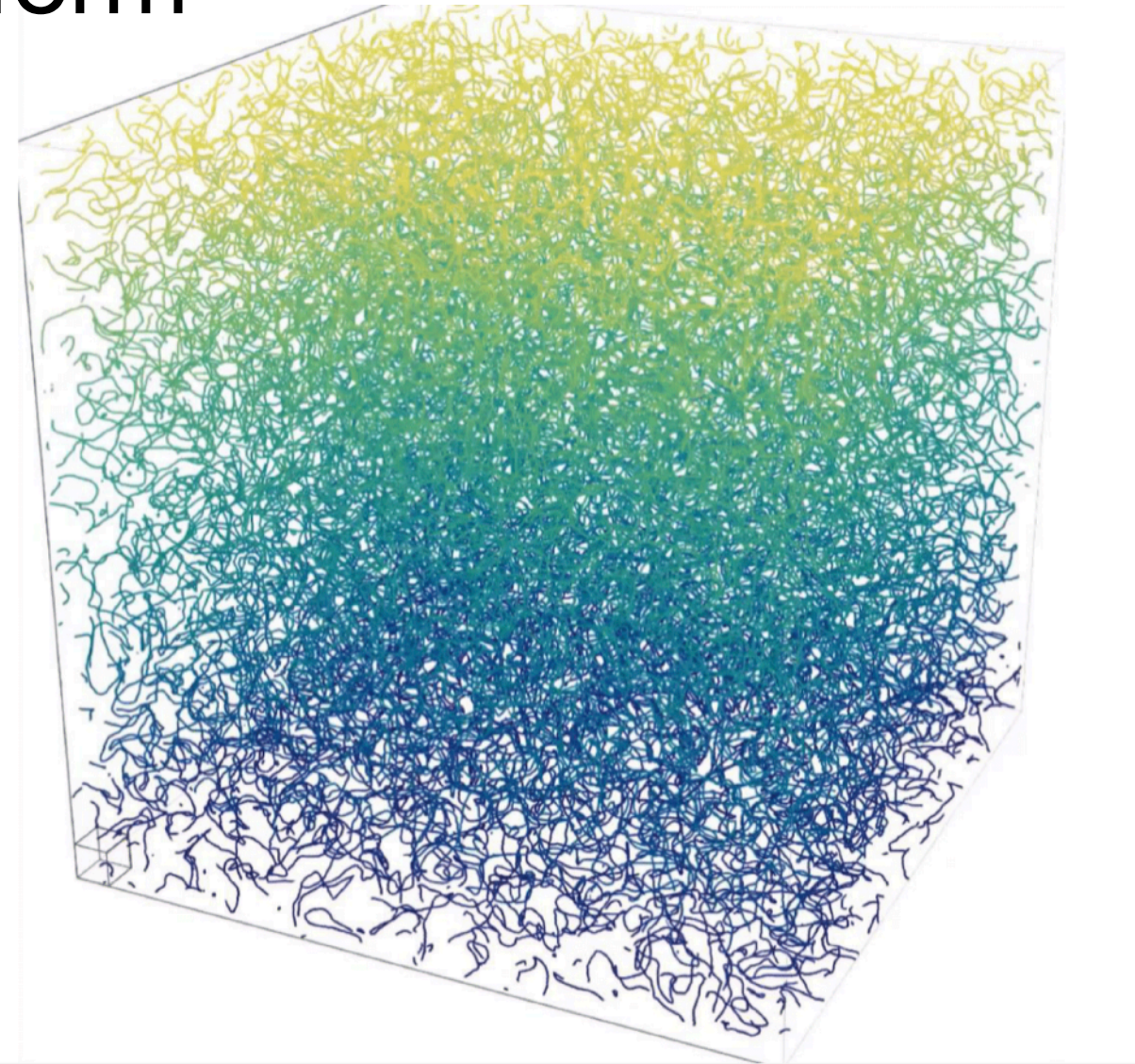
Cosmological evolution



Cosmological evolution



Strings form



Destroyed

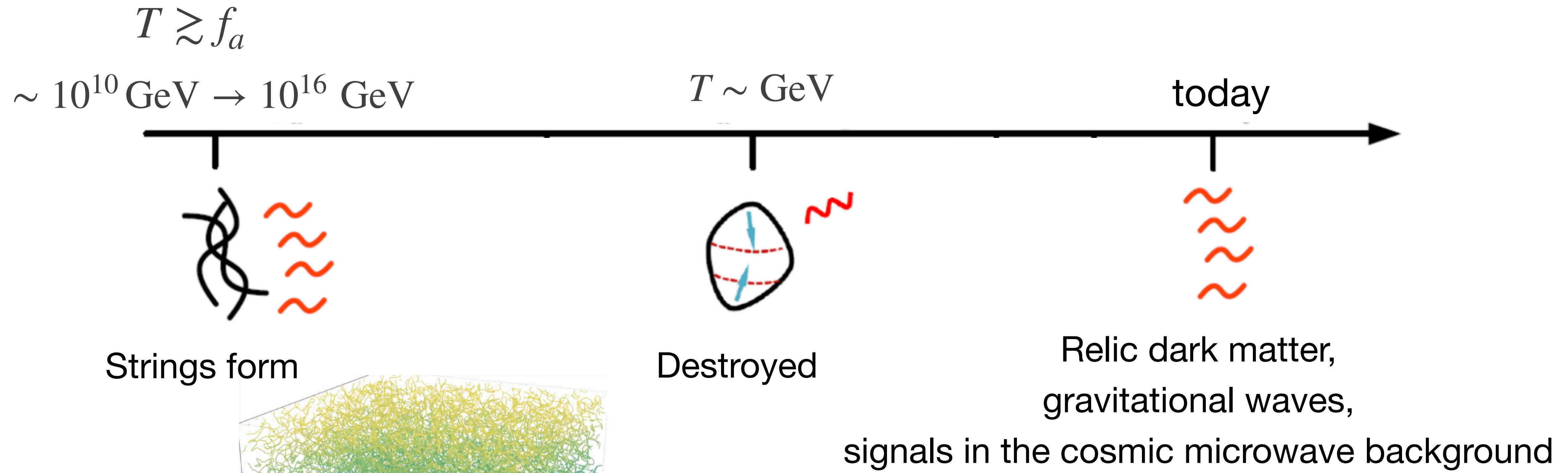
Relic dark matter,
gravitational waves,
signals in the cosmic microwave background

Dynamics:

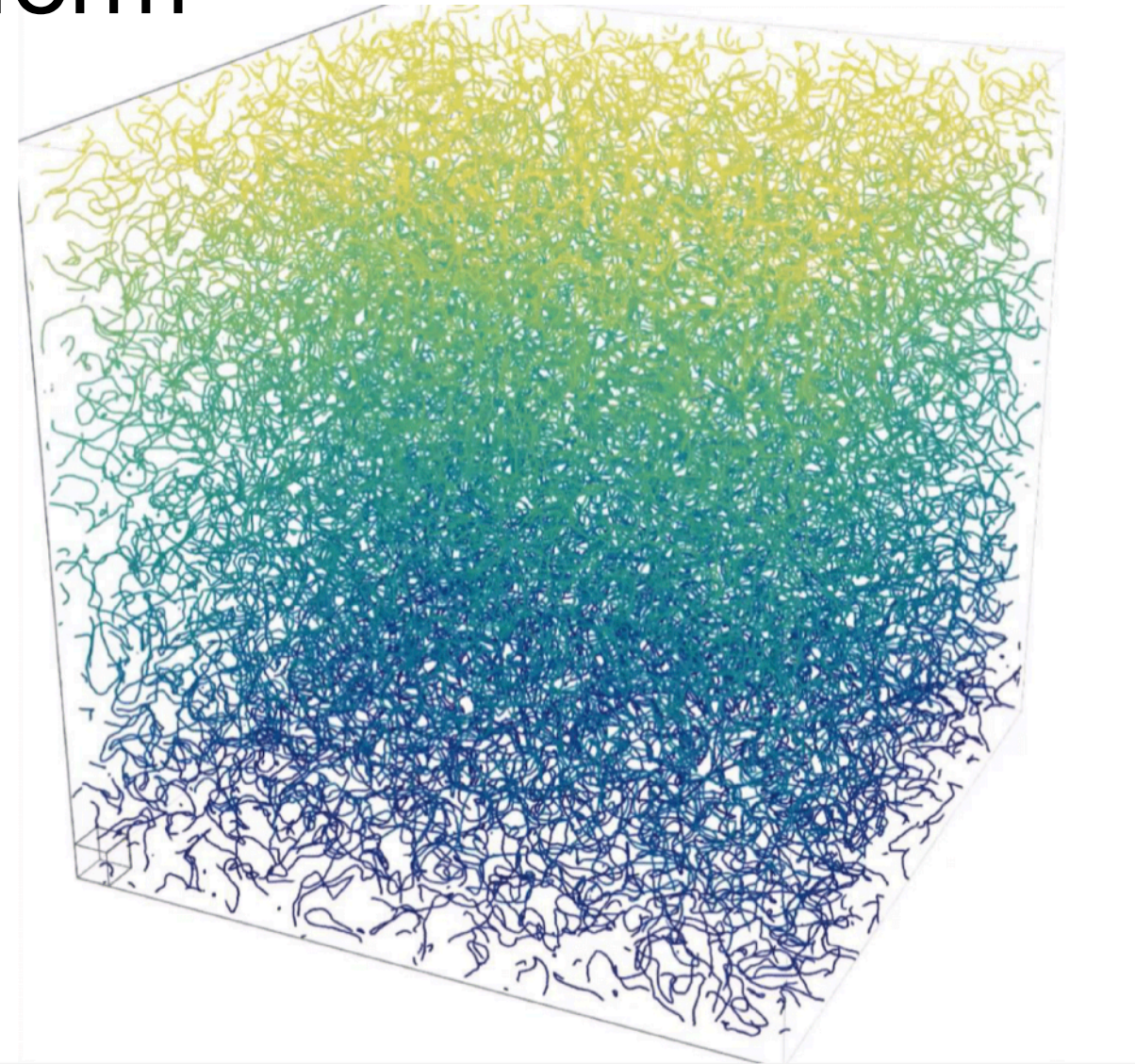
- *nonlinear*

- *large scale separation*

Cosmological evolution



Strings form



Destroyed

Relic dark matter,
gravitational waves,
signals in the cosmic microwave background

Dynamics:

- *nonlinear*

- *large scale separation*

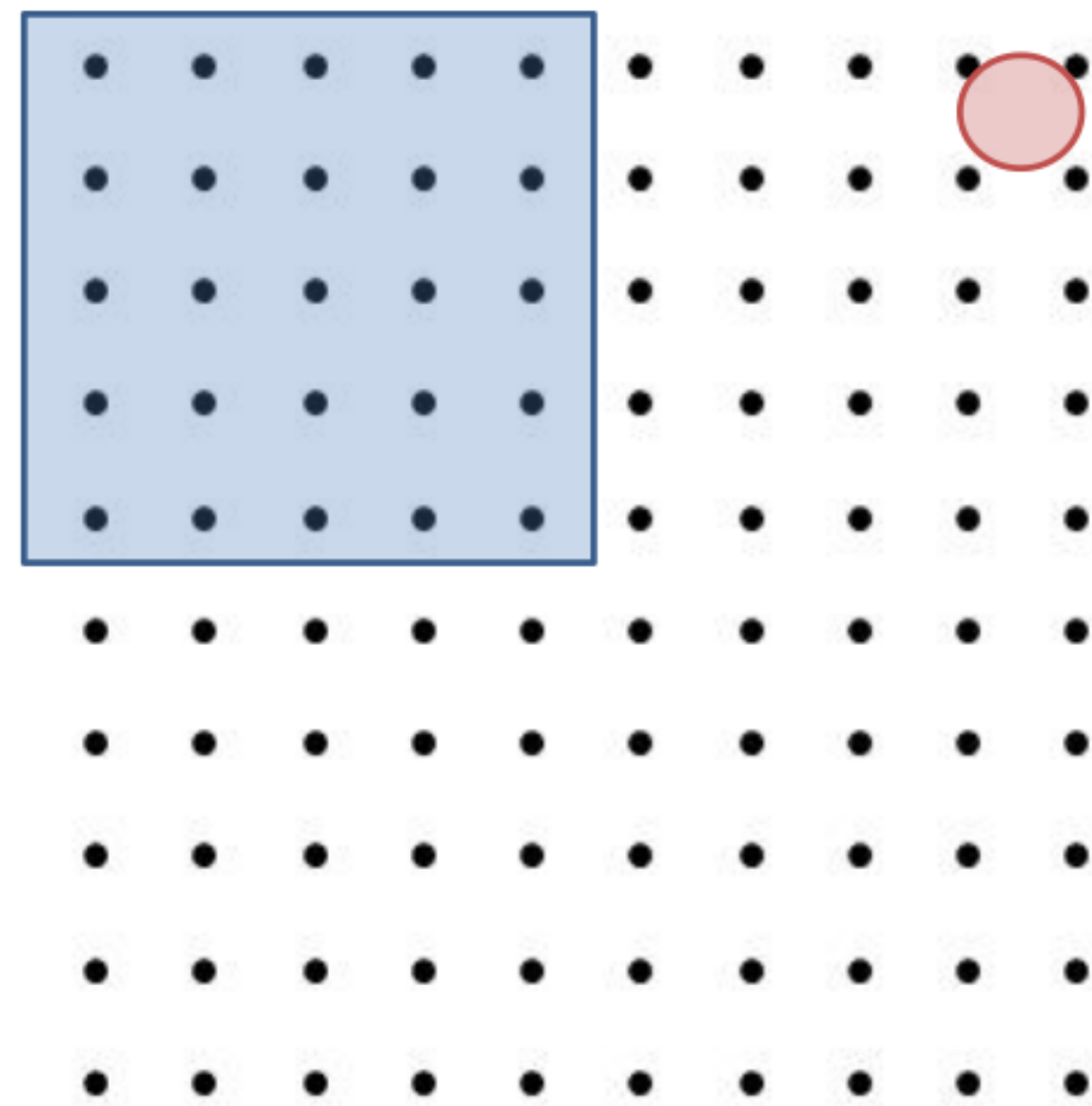
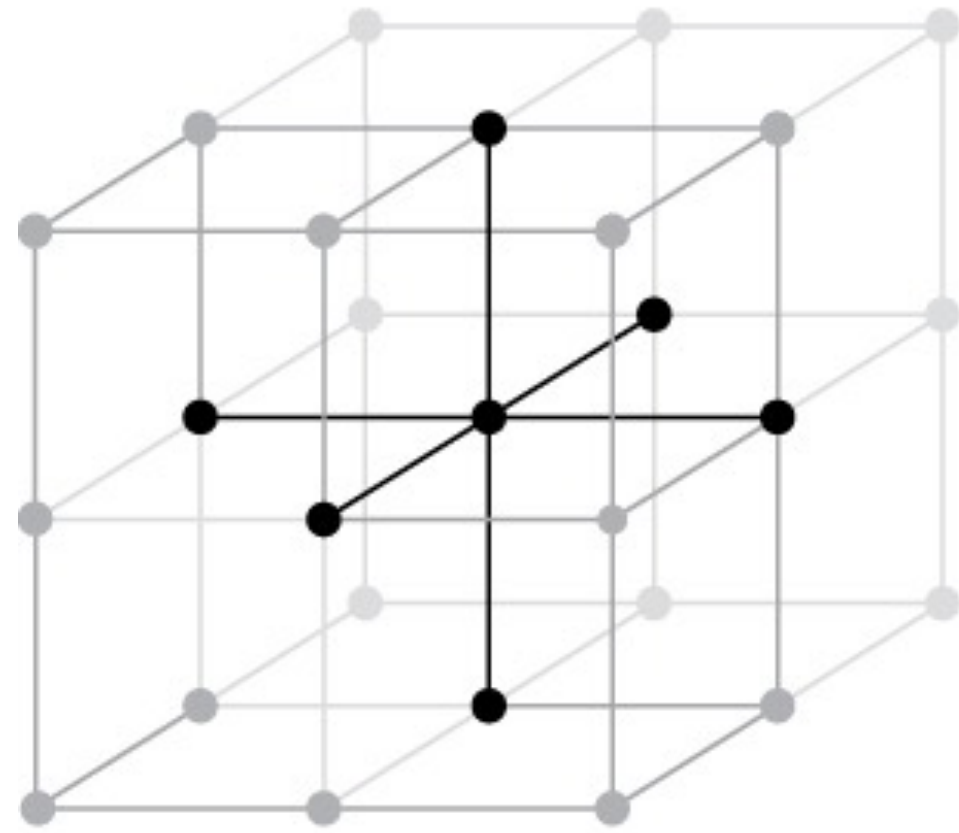
analytics



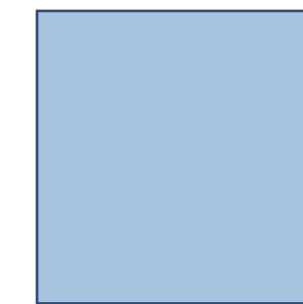
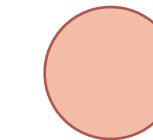
numerics



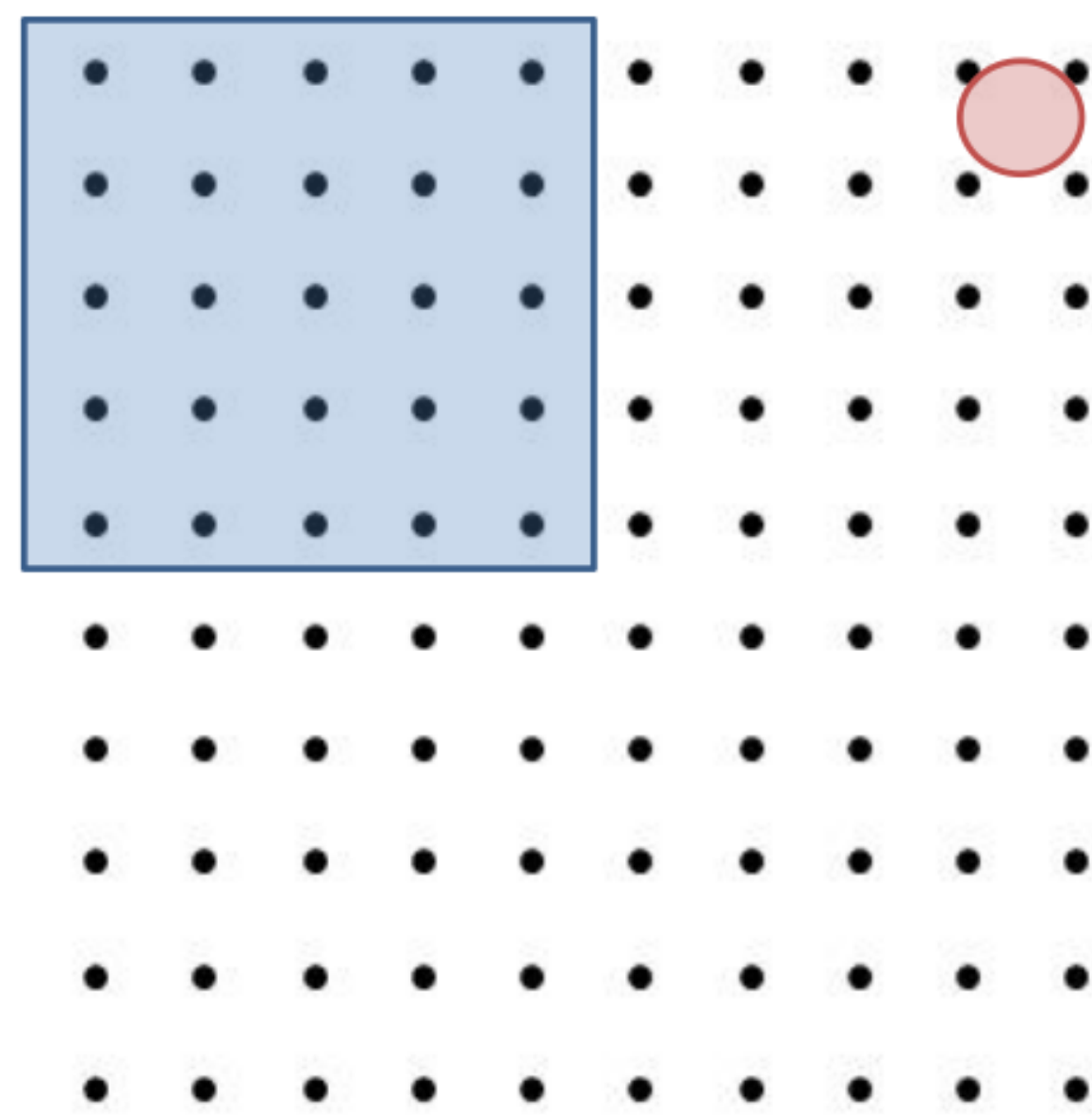
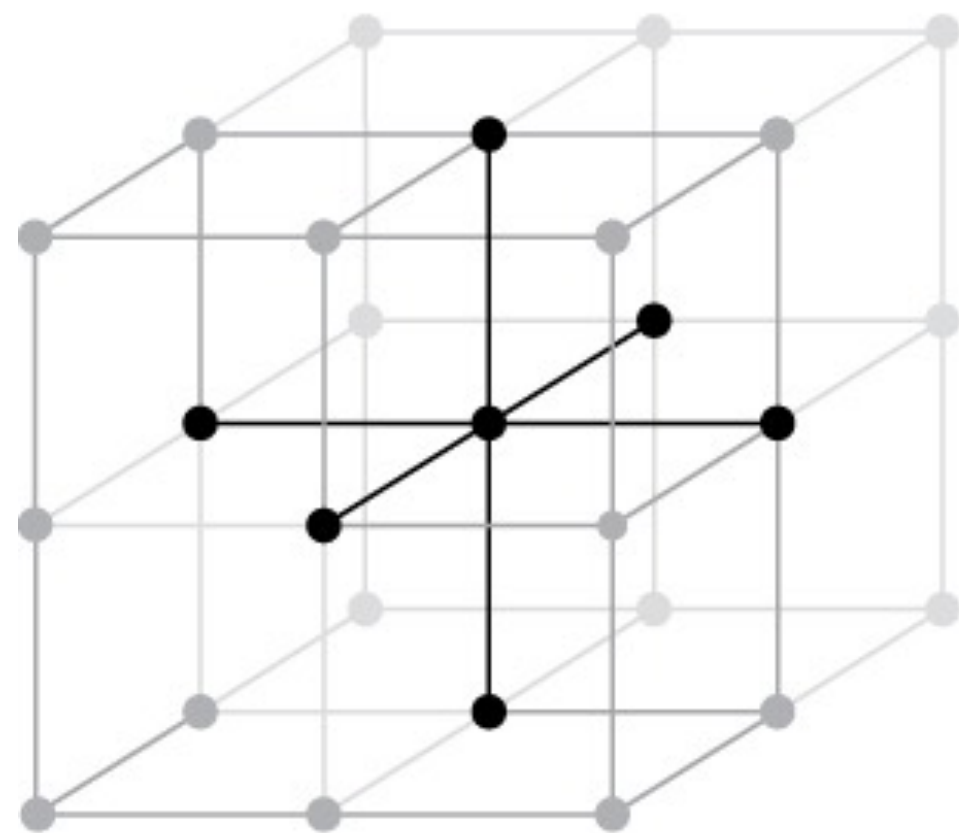
Simulations

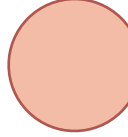




- a few lattice points per string core
- a few Hubble patches



Simulations

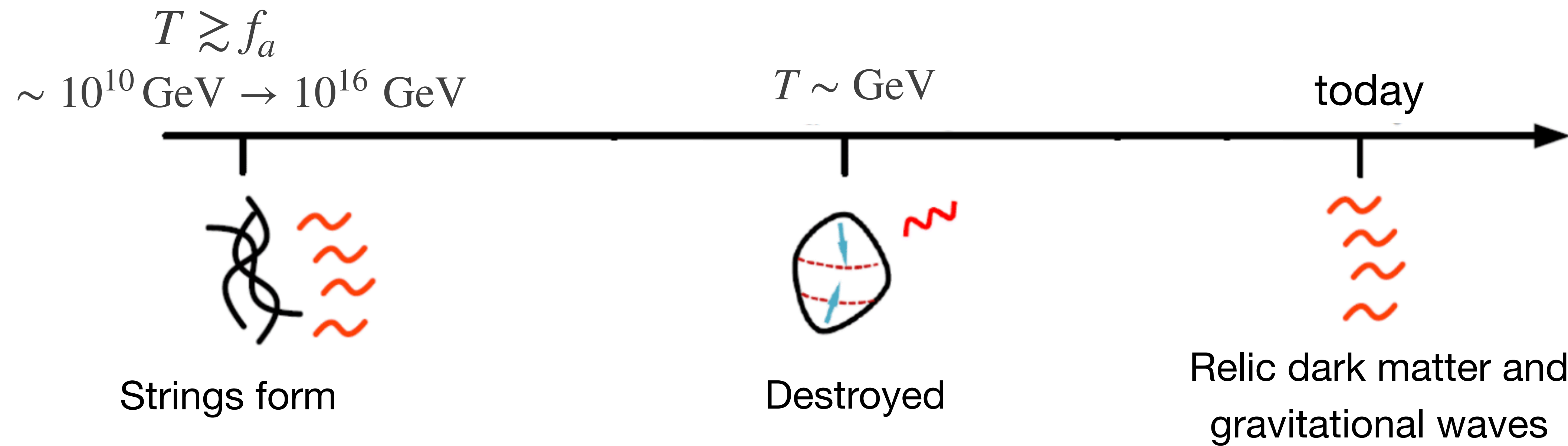


- a few lattice points per string core 
- a few Hubble patches 

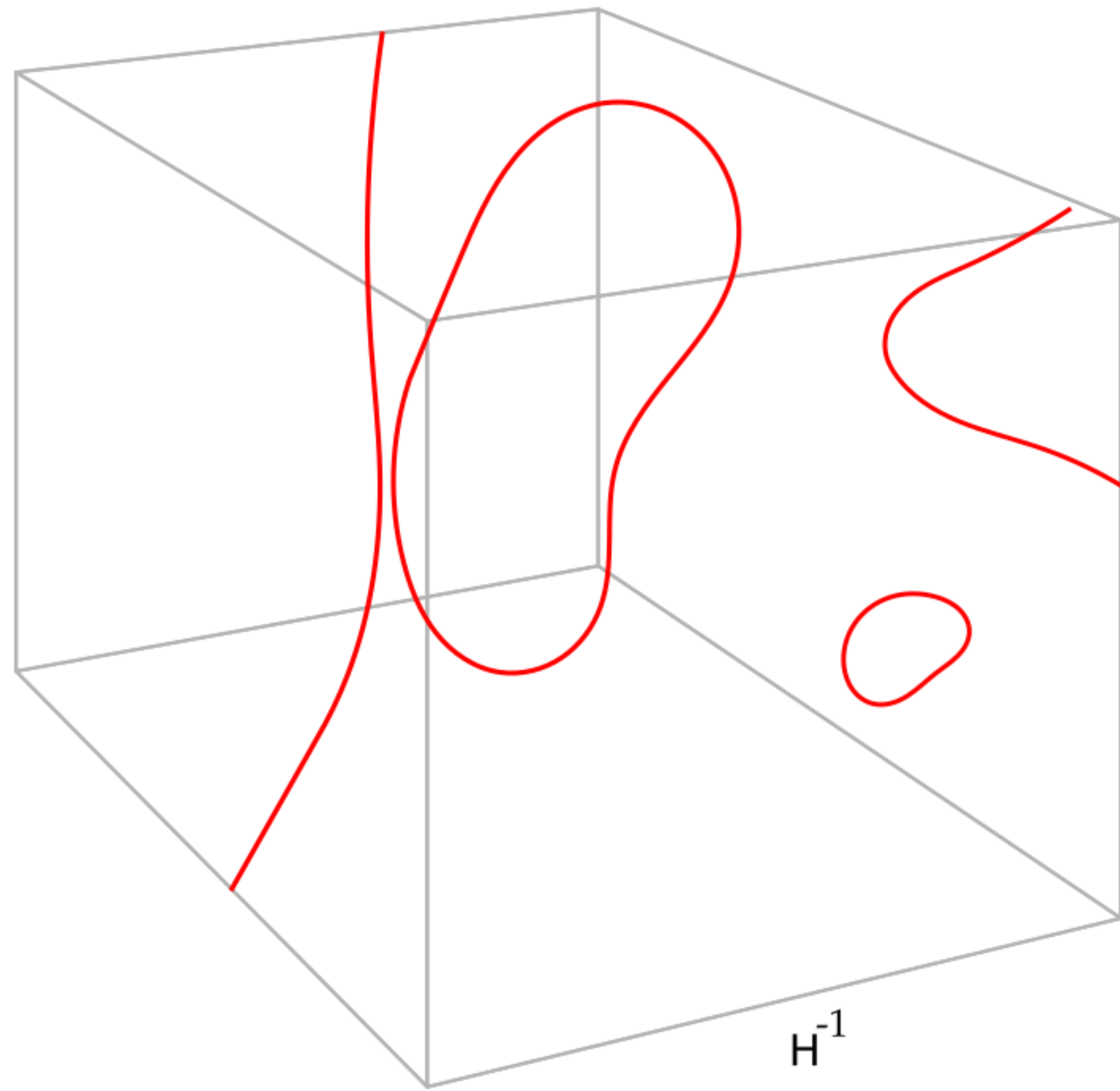
Memory constraints  max 5000^3 grid points

Simulations $\frac{f_a}{H} \lesssim \frac{\text{blue square}}{\text{red circle}} \lesssim 1000$

Physical $\frac{f_a}{H} \sim 10^{30}$



Scaling solution



~ one Hubble length of string per Hubble patch

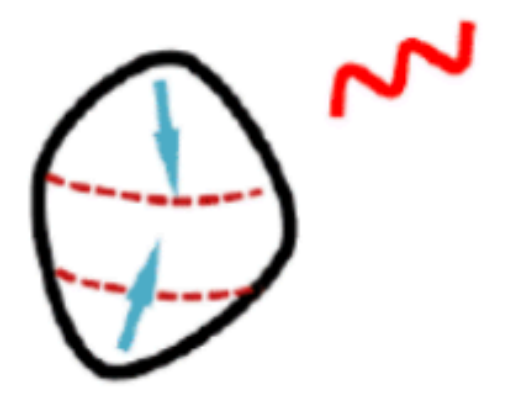
$T \gtrsim f_a$
 $\sim 10^{10} \text{ GeV} \rightarrow 10^{16} \text{ GeV}$

$T \sim \text{GeV}$

today



Strings form



Destroyed

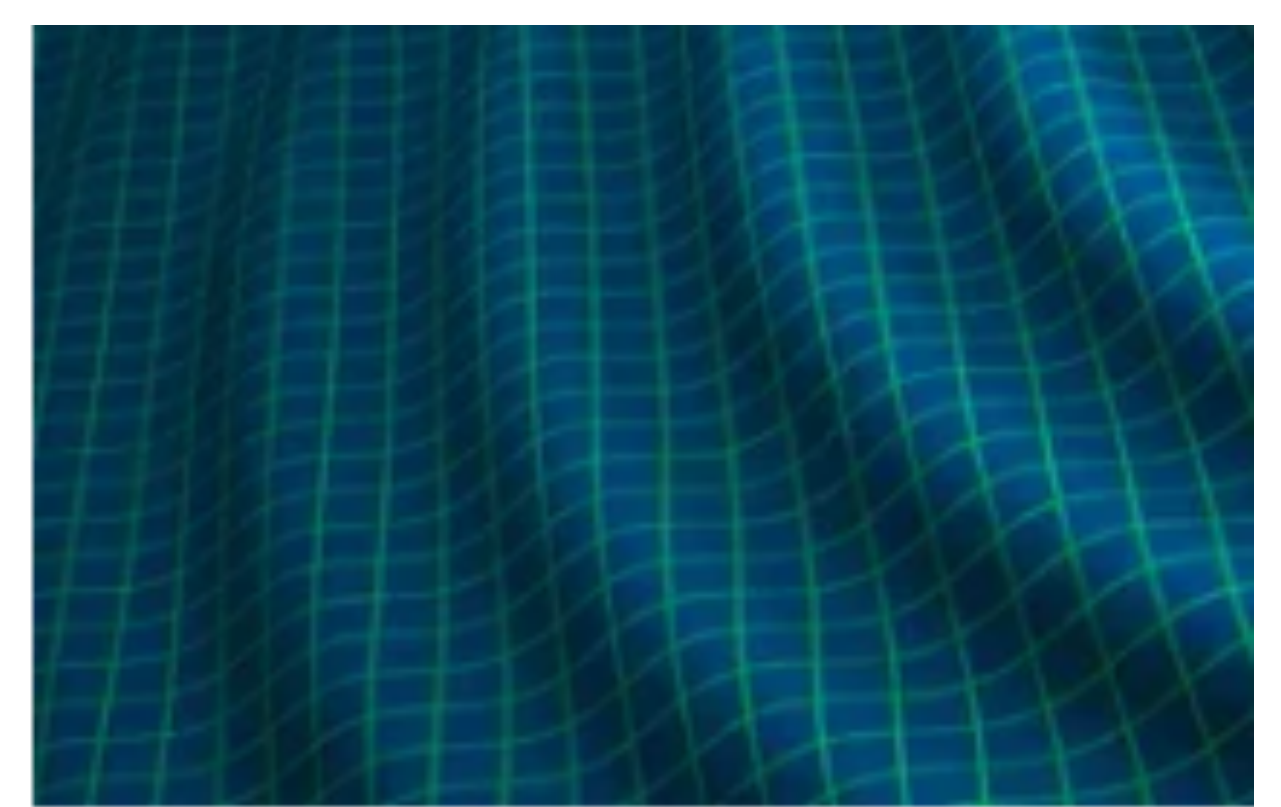
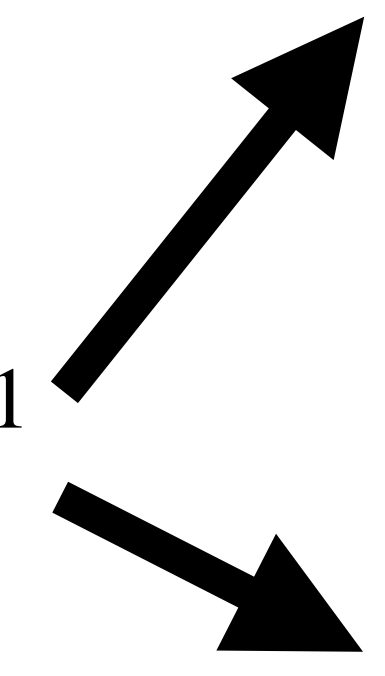


Energy emitted
per Hubble time
and Hubble
volume

$$\simeq \pi f_a^2 \times H^{-1}$$

Dark matter

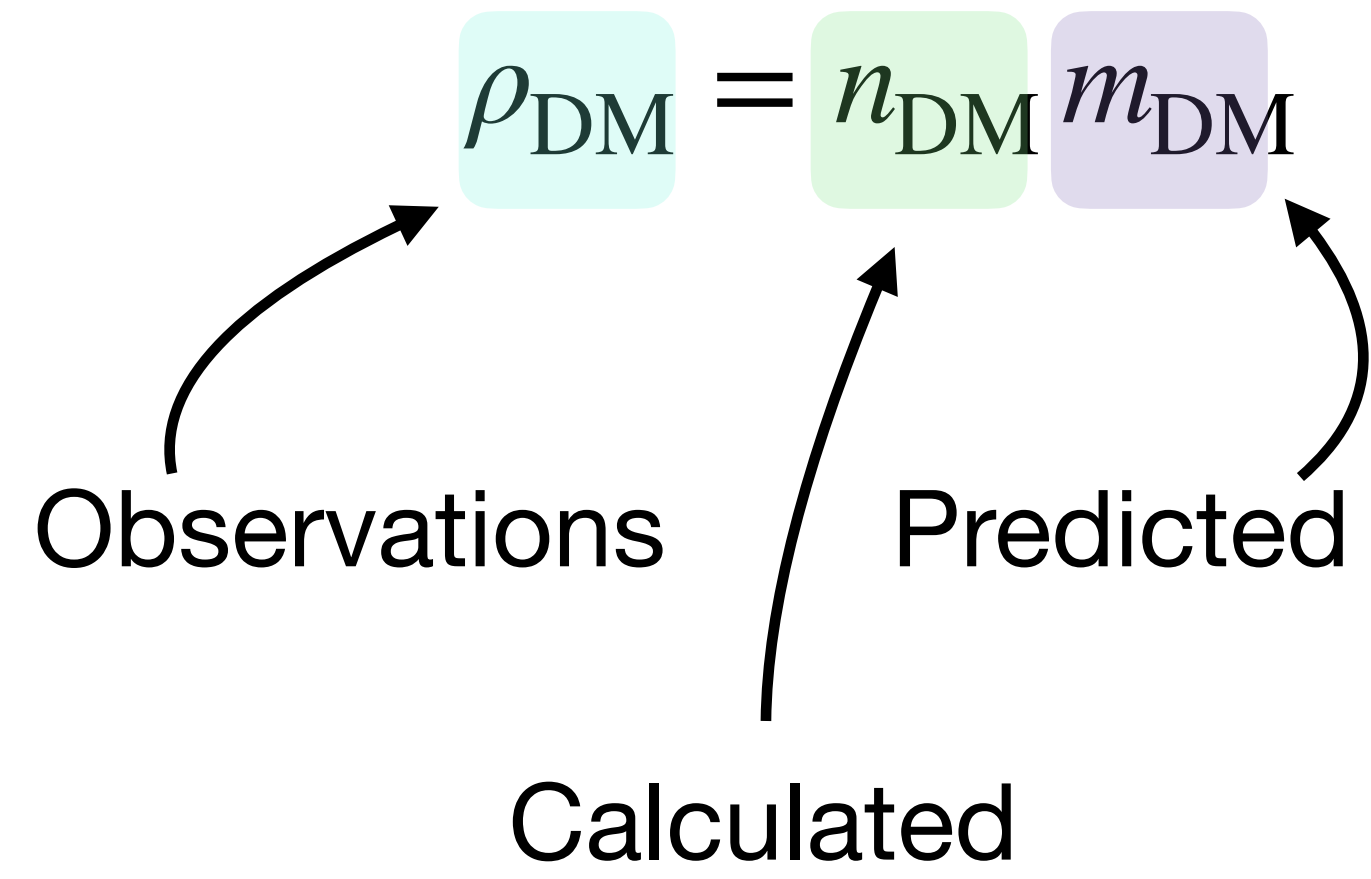
Gravitational waves



Dark matter abundance

$$\rho_{\text{DM}} = n_{\text{DM}} m_{\text{DM}}$$

Dark matter abundance



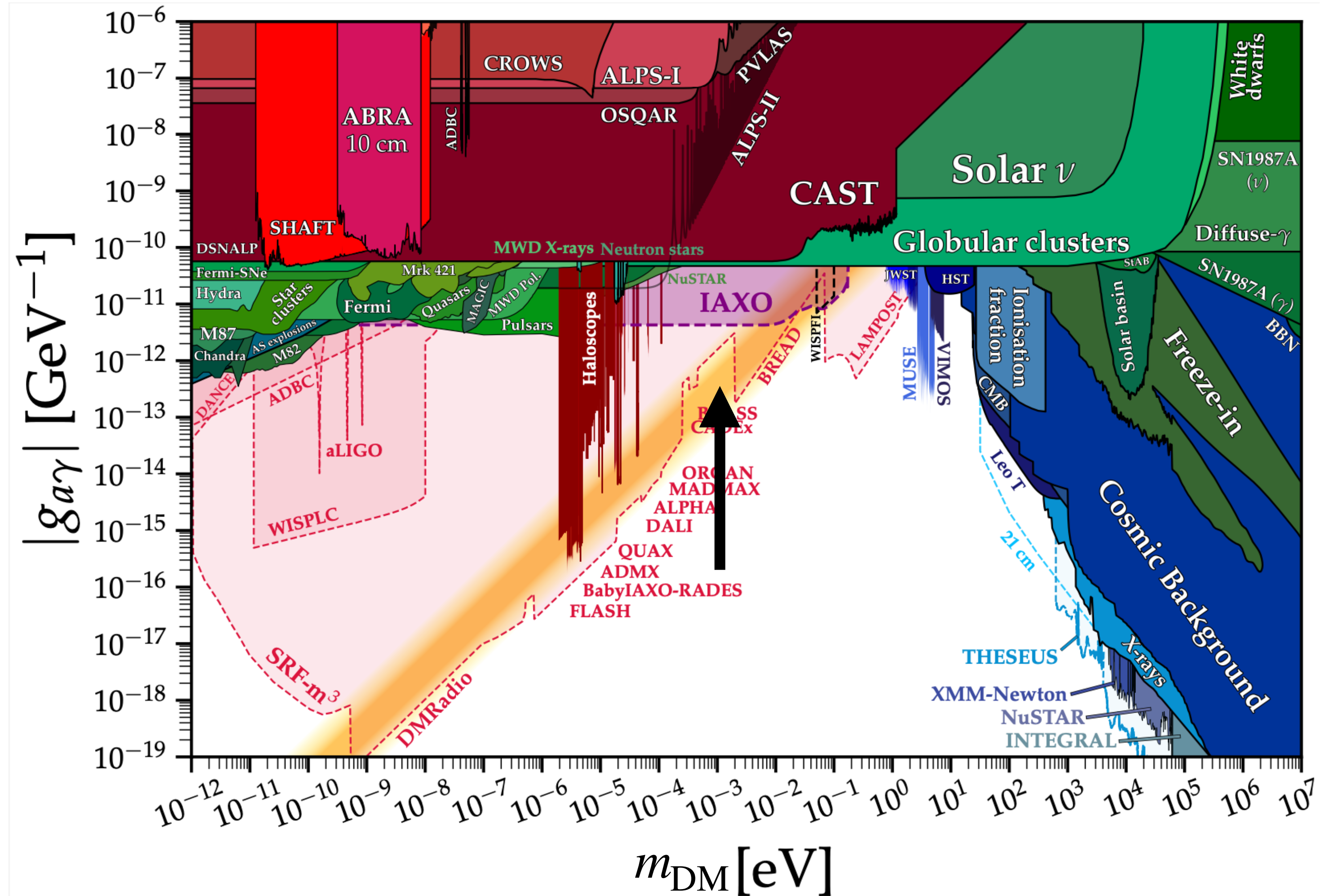
Dark matter abundance

$$\rho_{\text{DM}} = n_{\text{DM}} m_{\text{DM}}$$

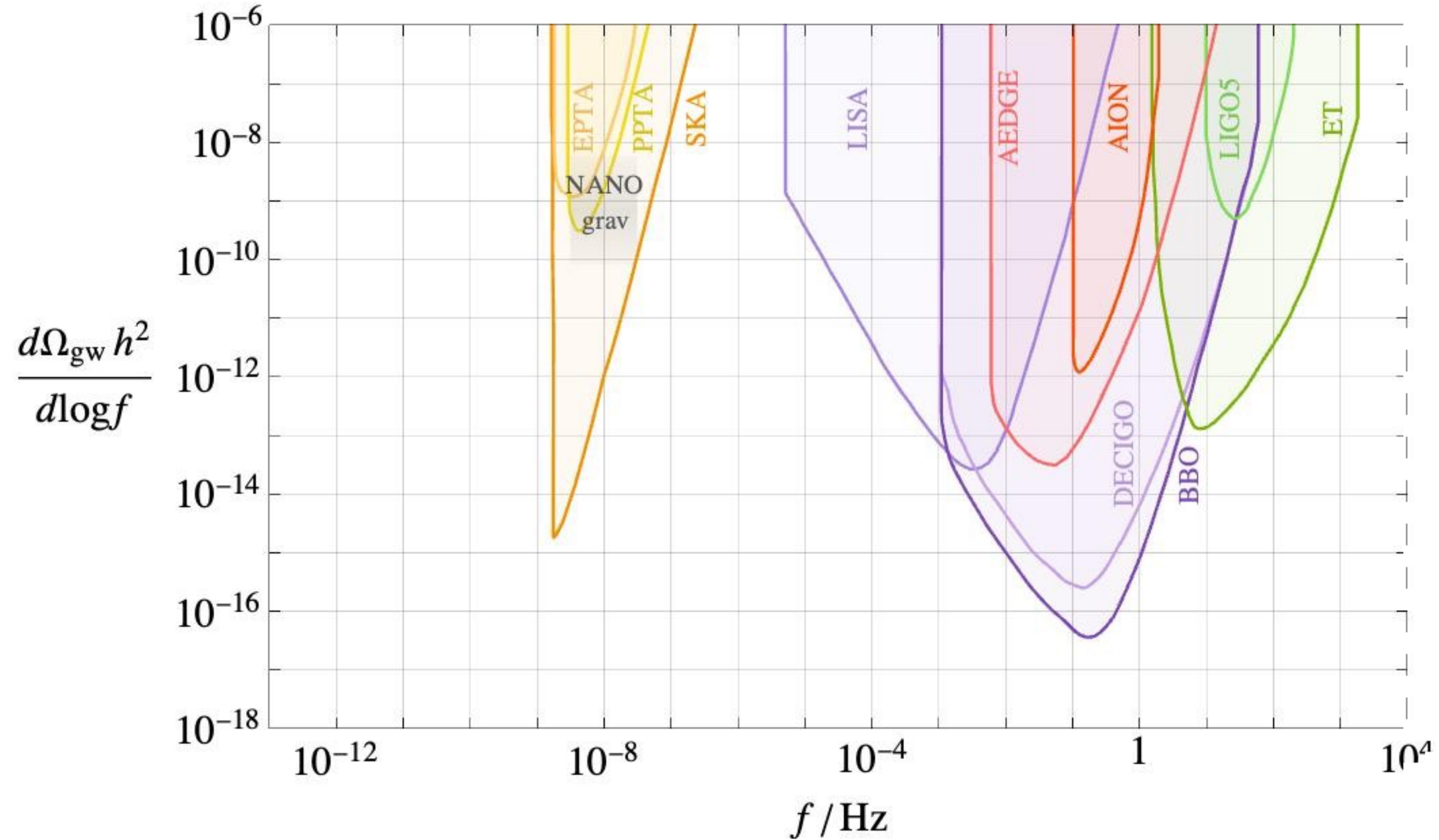
Observations → ρ_{DM}

Calculated → n_{DM}

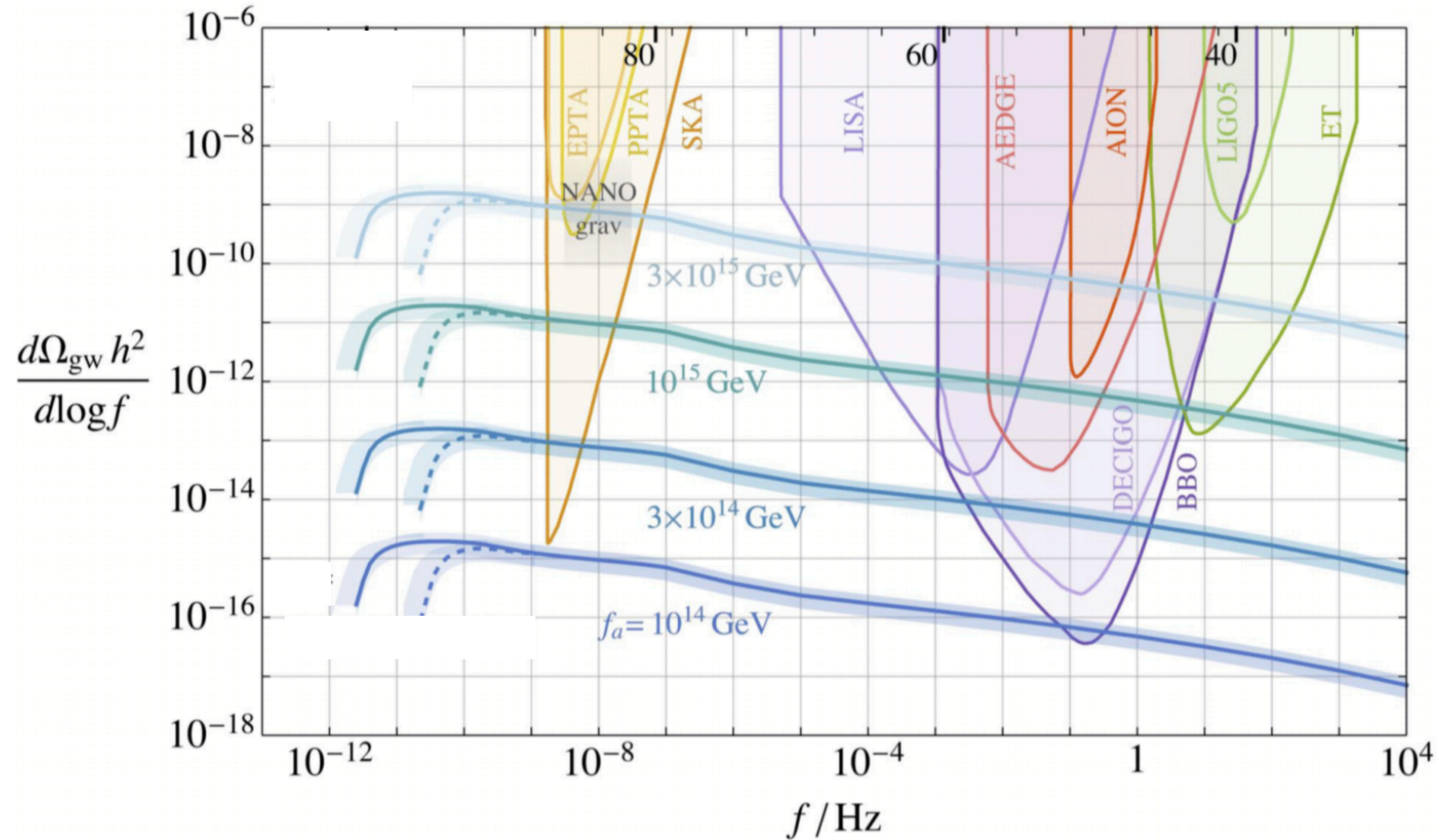
Predicted → m_{DM}



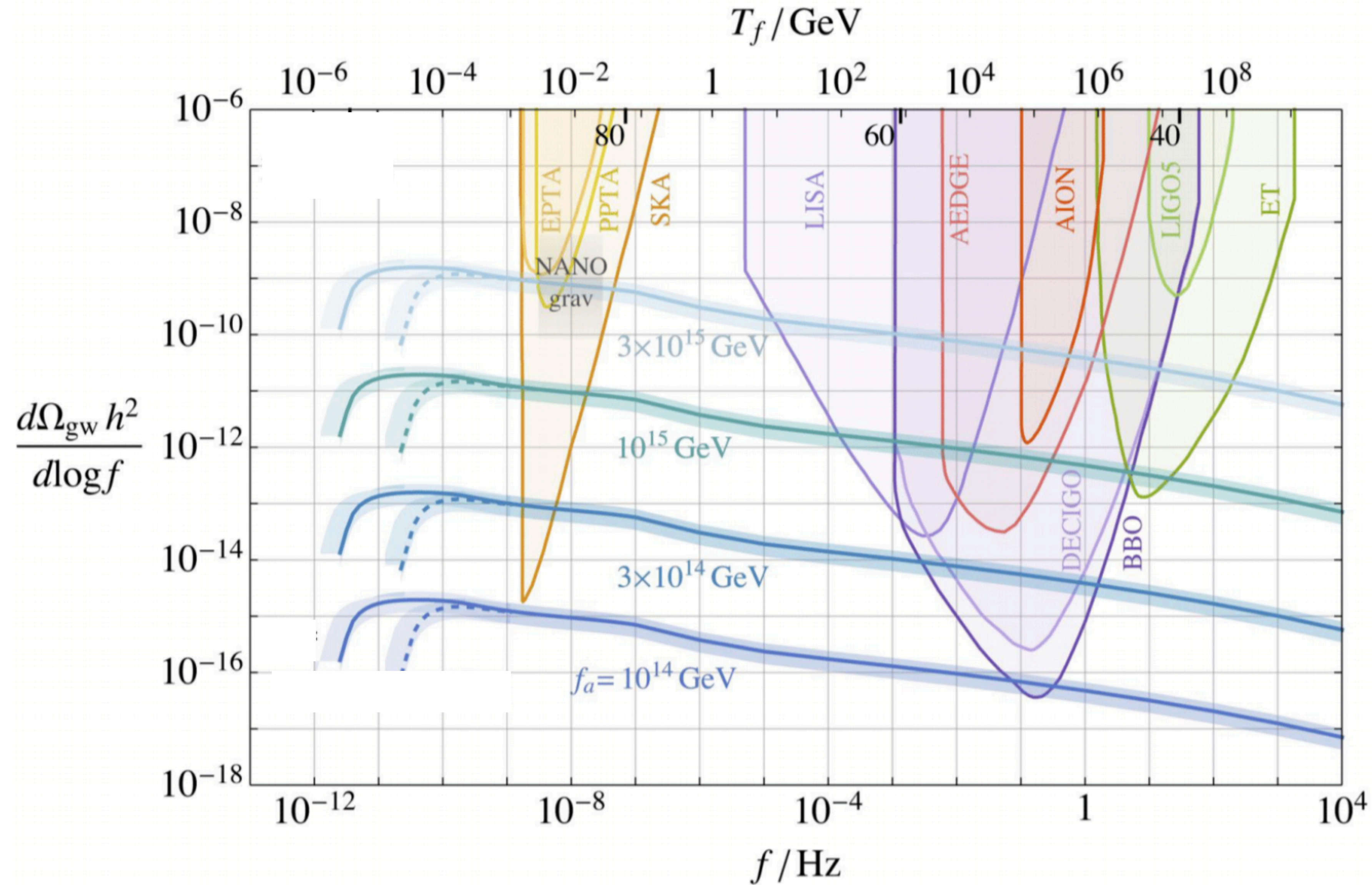
Gravitational waves



Gravitational waves

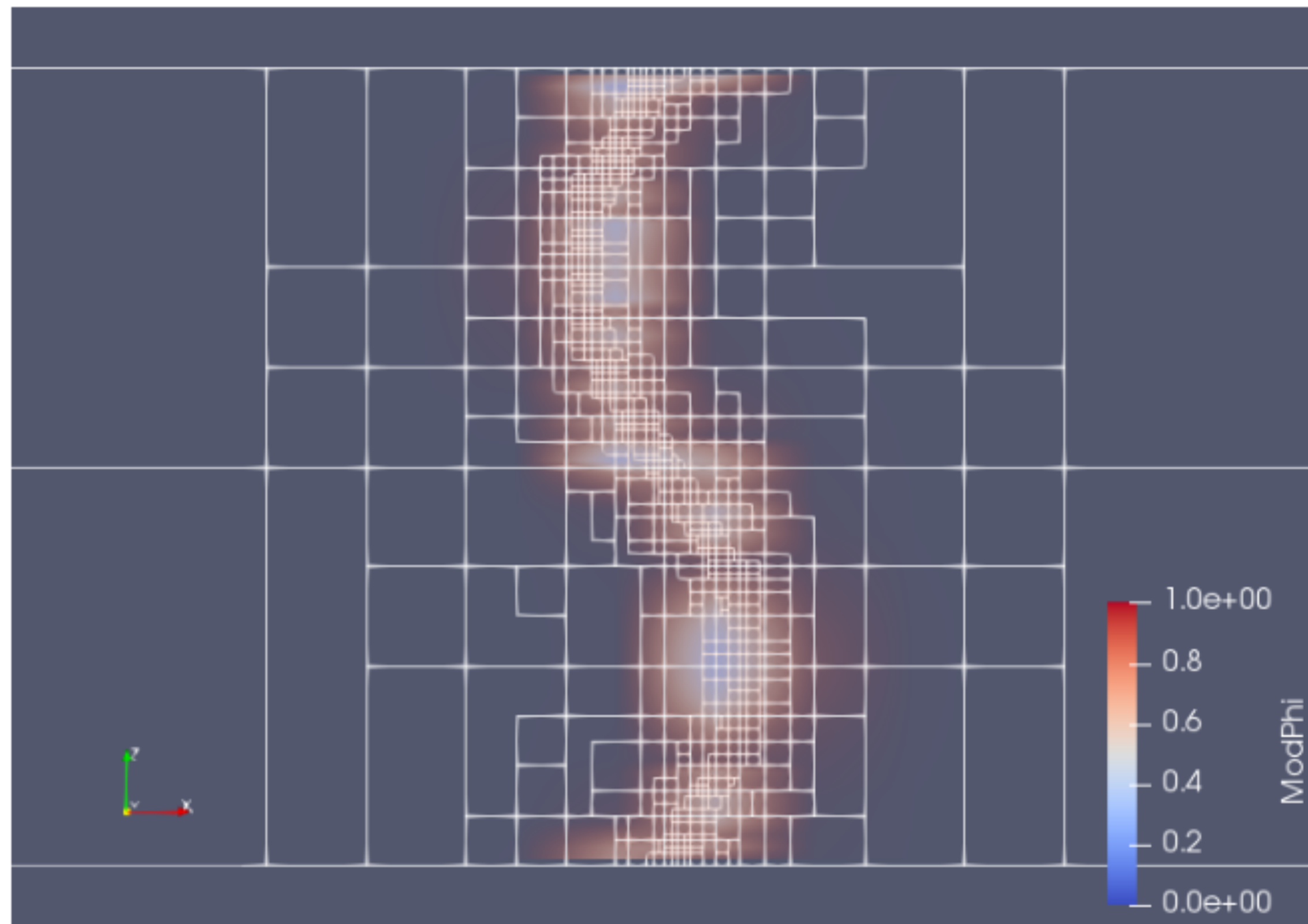


Gravitational waves

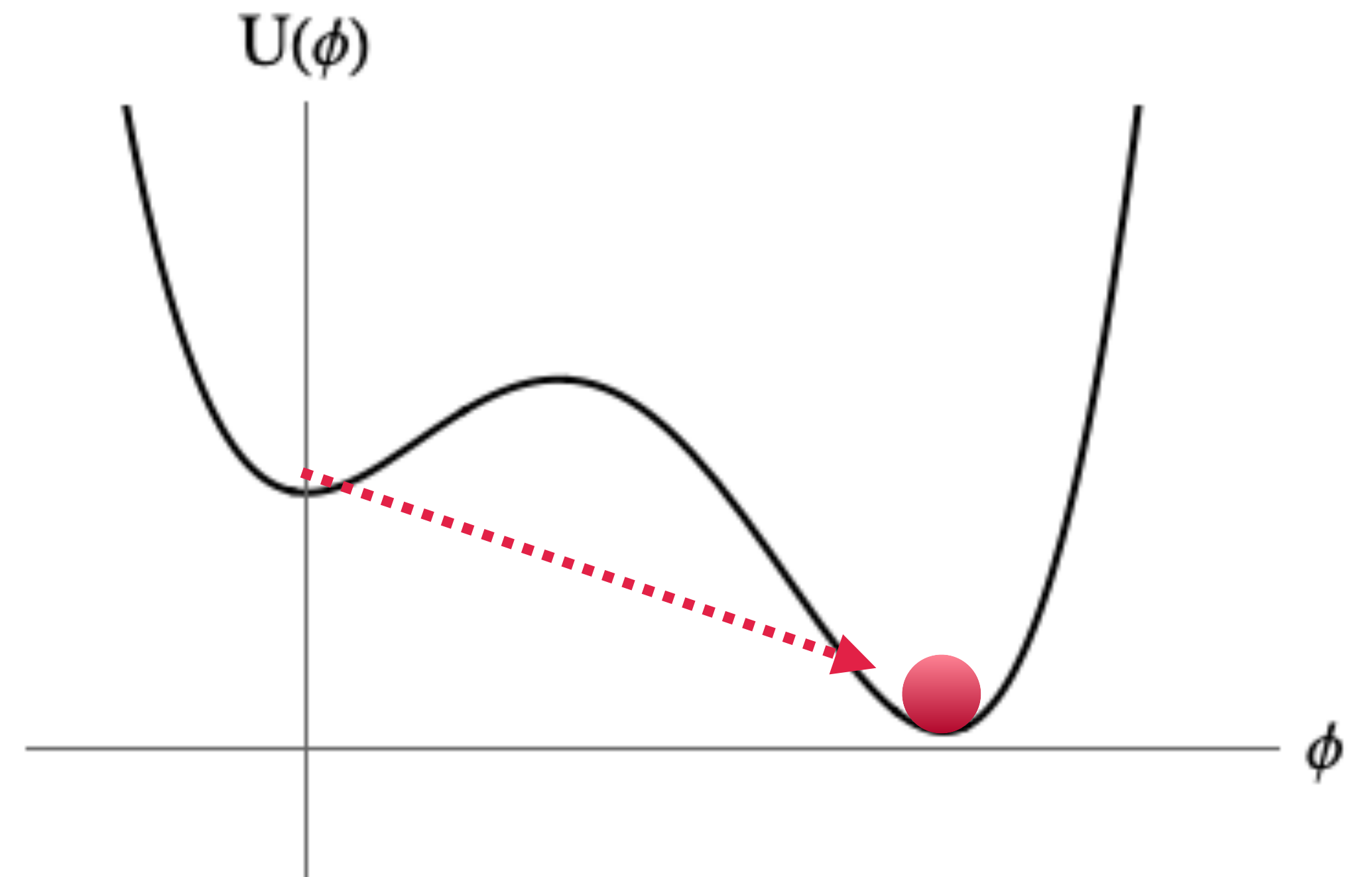


Work in progress

Adaptive meshing



1st order phase transitions



Summary

- Spontaneous symmetry breaking $\xrightarrow{\text{(often)}}$ “Topological defects”
- Persist from the early universe
- Access to ultra-high energy scales $\sim 10^{16}$ GeV
- Also the very early Universe: $T \sim 10^8$ GeV $\implies t_{\text{universe}} \sim 10^{-22}$ second
- Ongoing experimental and theoretical effort

Summary

- Spontaneous symmetry breaking $\xrightarrow{\text{(often)}}$ “Topological defects”
- Persist from the early universe
- Access to ultra-high energy scales $\sim 10^{16}$ GeV
- Also the very early Universe: $T \sim 10^8$ GeV $\implies t_{\text{universe}} \sim 10^{-22}$ second
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Thanks

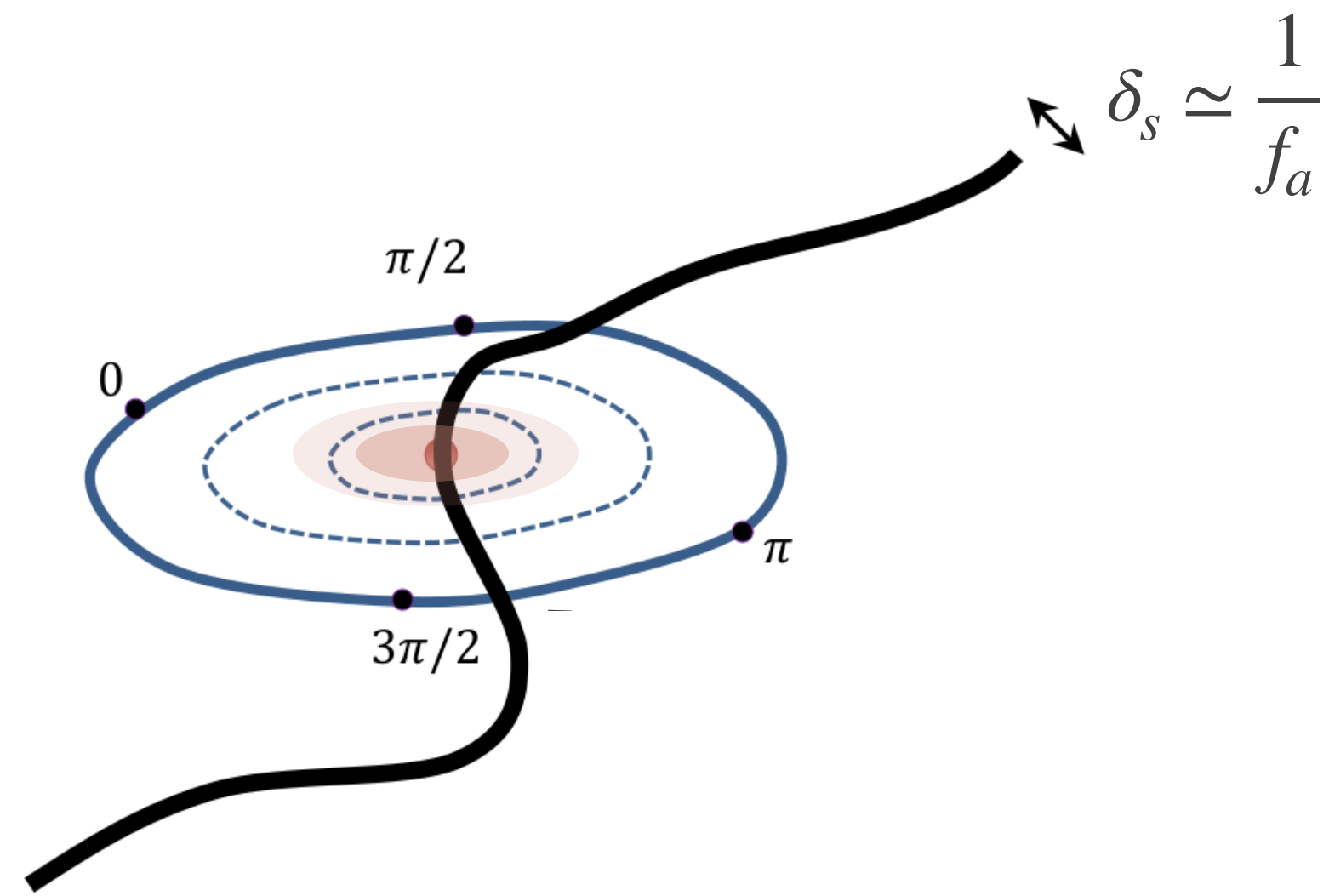
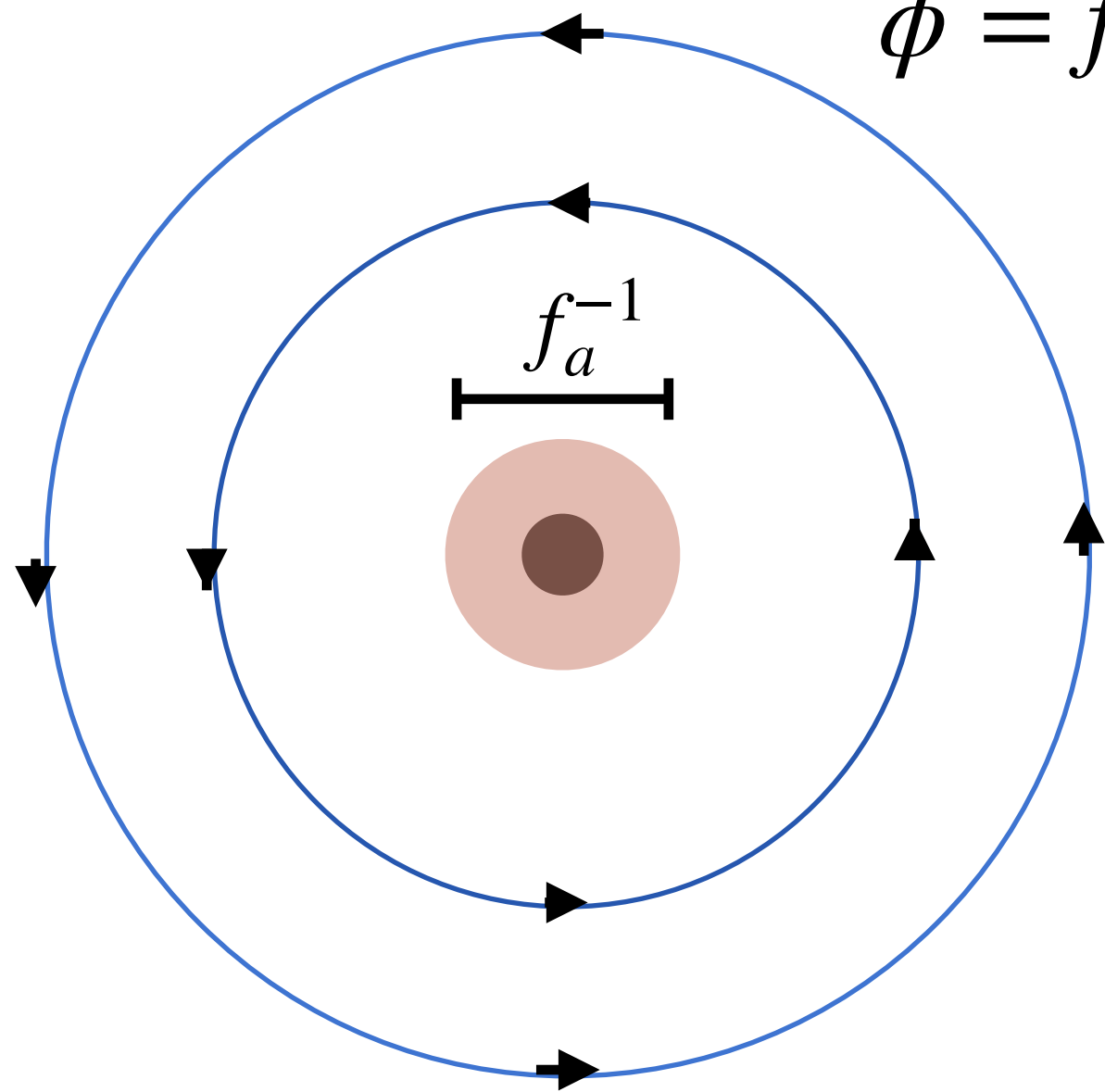
Backup

$$U(\phi) = (|\phi|^2 - f_a^2)^2$$

String tension $\mu = E/L$

$$\mathcal{H} \sim (\partial_i \phi)^2 + U(\phi)$$

$$\phi = f_a g(|x|) e^{i\theta}$$

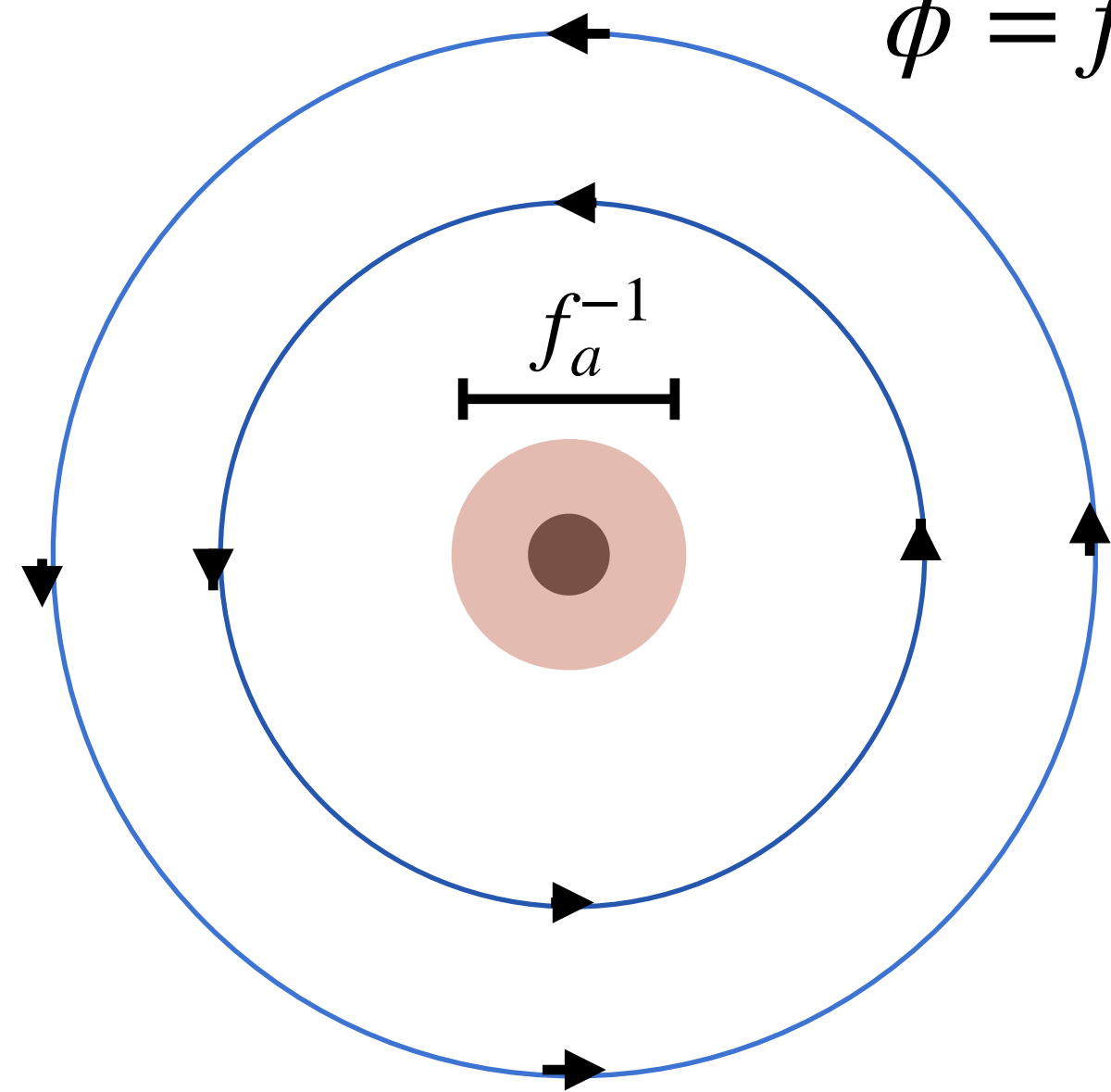


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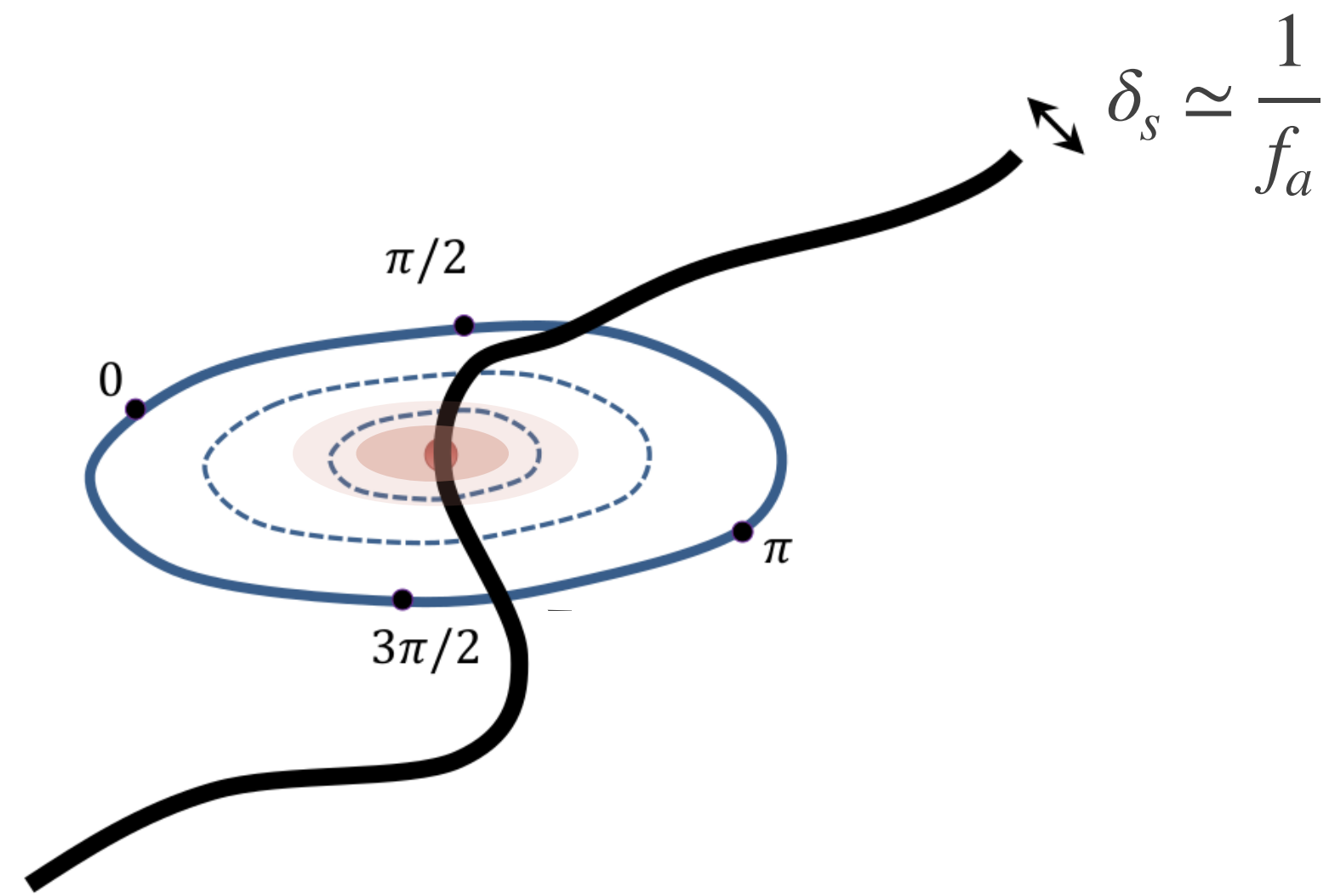
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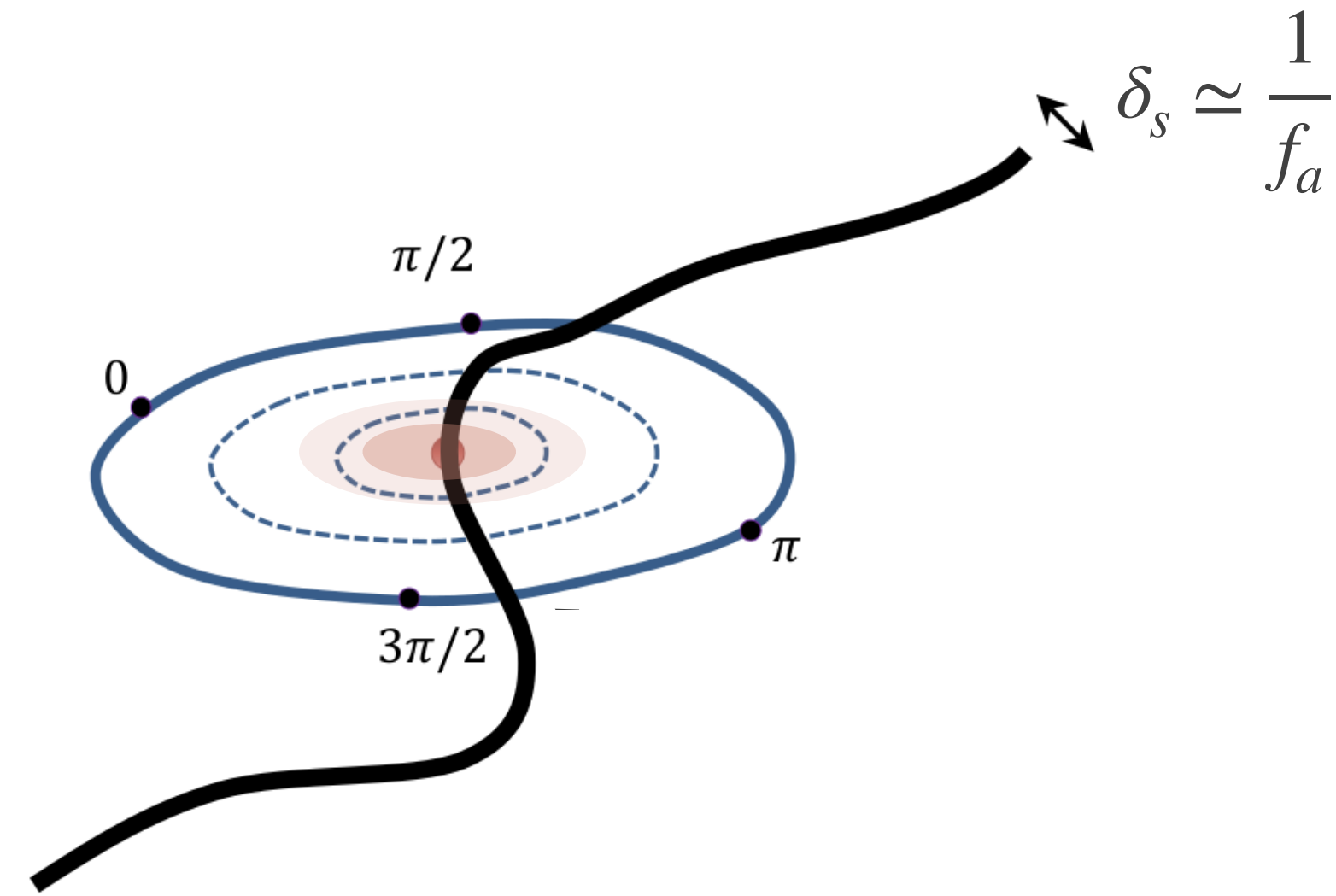
Core : πf_a^2



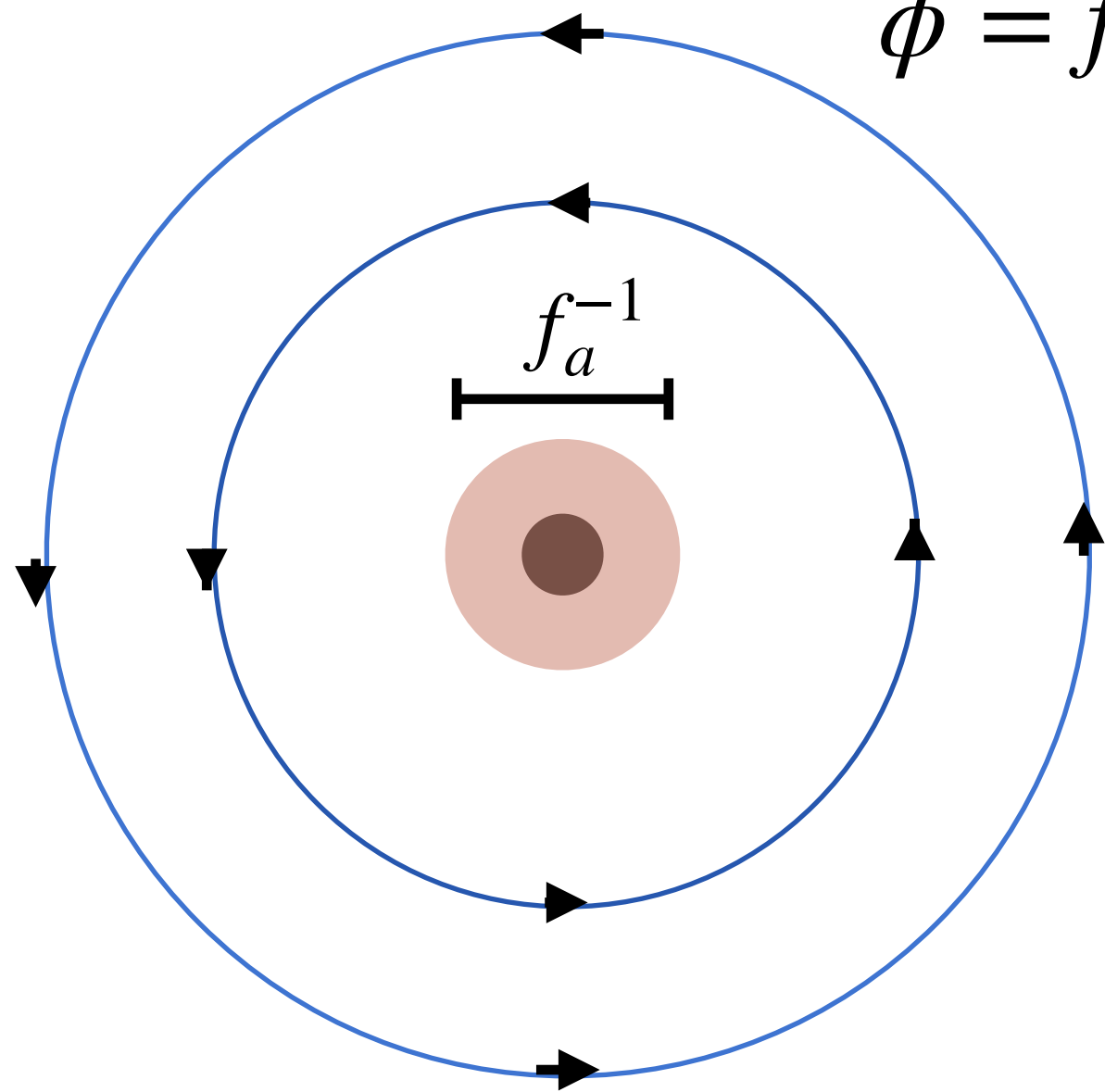
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Core : πf_a^2

$$\text{Gradient : } \int_{f_a^{-1}}^{r_{\max}} dr 2\pi r \left| \frac{\partial_\theta \phi}{r} \right|^2$$

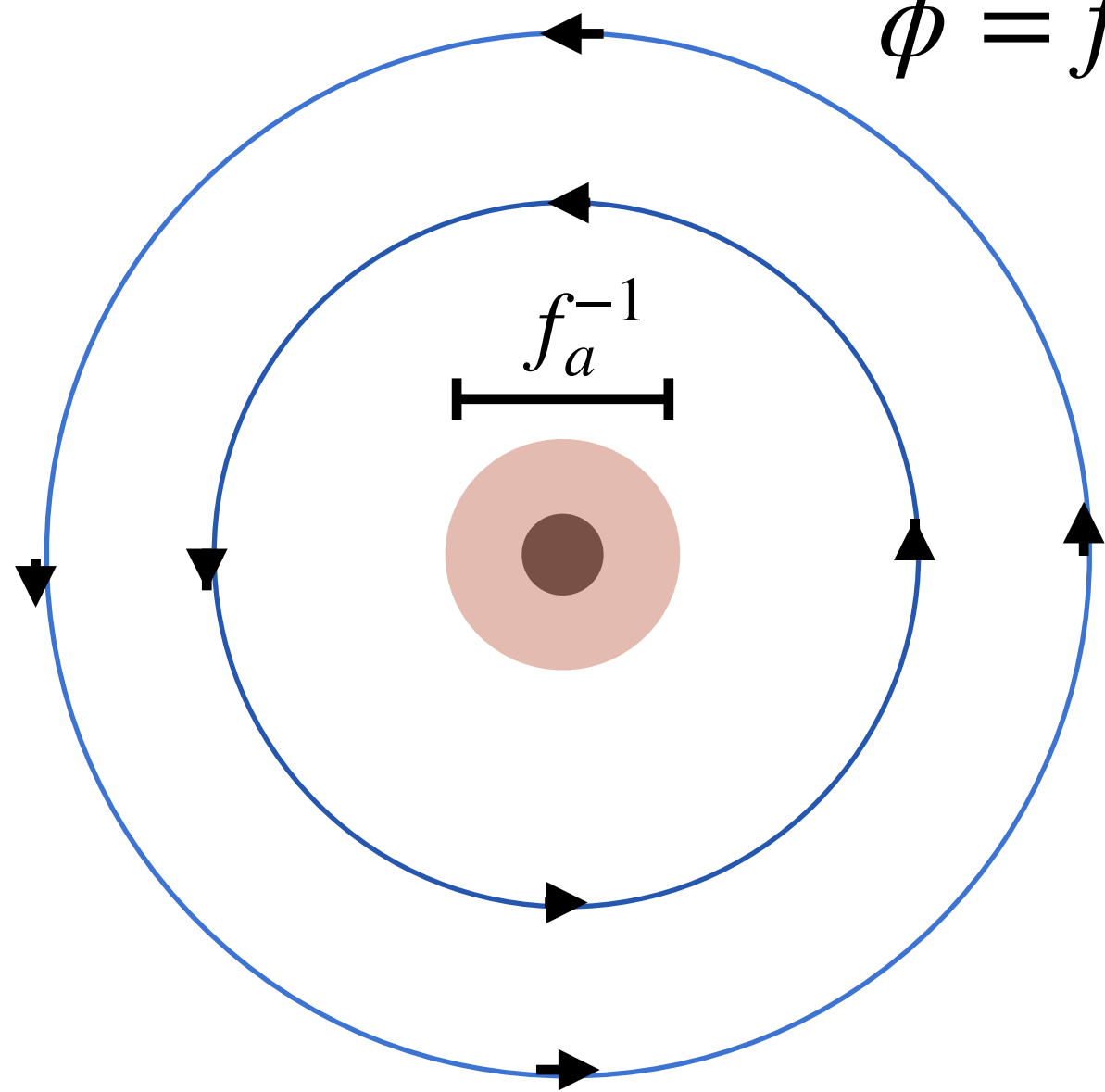
$$\sim \pi f_a^2 \log \left(\frac{r_{\max}}{f_a^{-1}} \right)$$

$$U(\phi) = (|\phi|^2 - f_a^2)^2$$

String tension $\mu = E/L$

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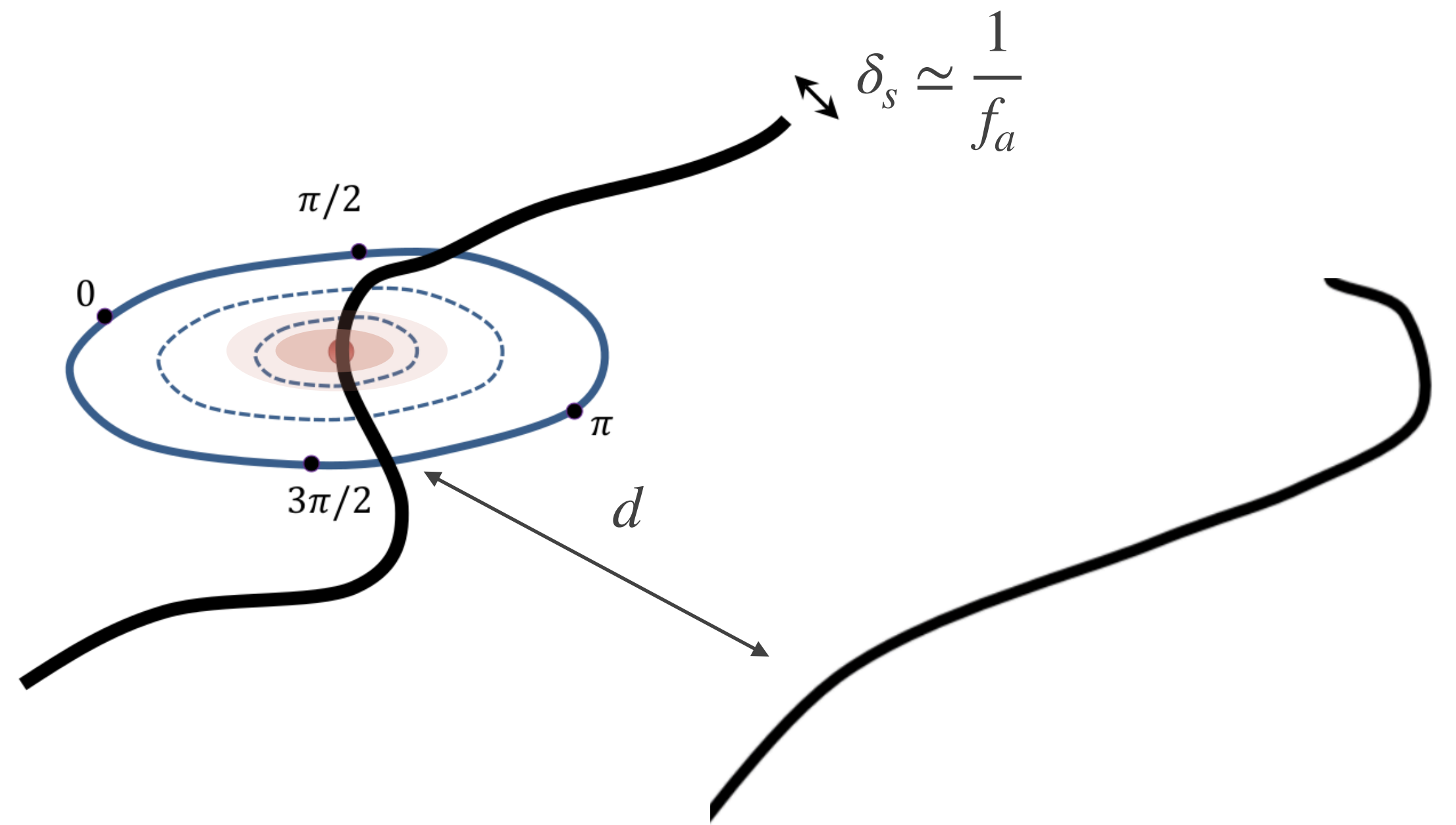
$$\phi = f_a g(|x|) e^{i\theta}$$



Core : πf_a^2

$$\text{Gradient : } \int_{f_a^{-1}}^{r_{\max}} dr 2\pi r \left| \frac{\partial_\theta \phi}{r} \right|^2$$

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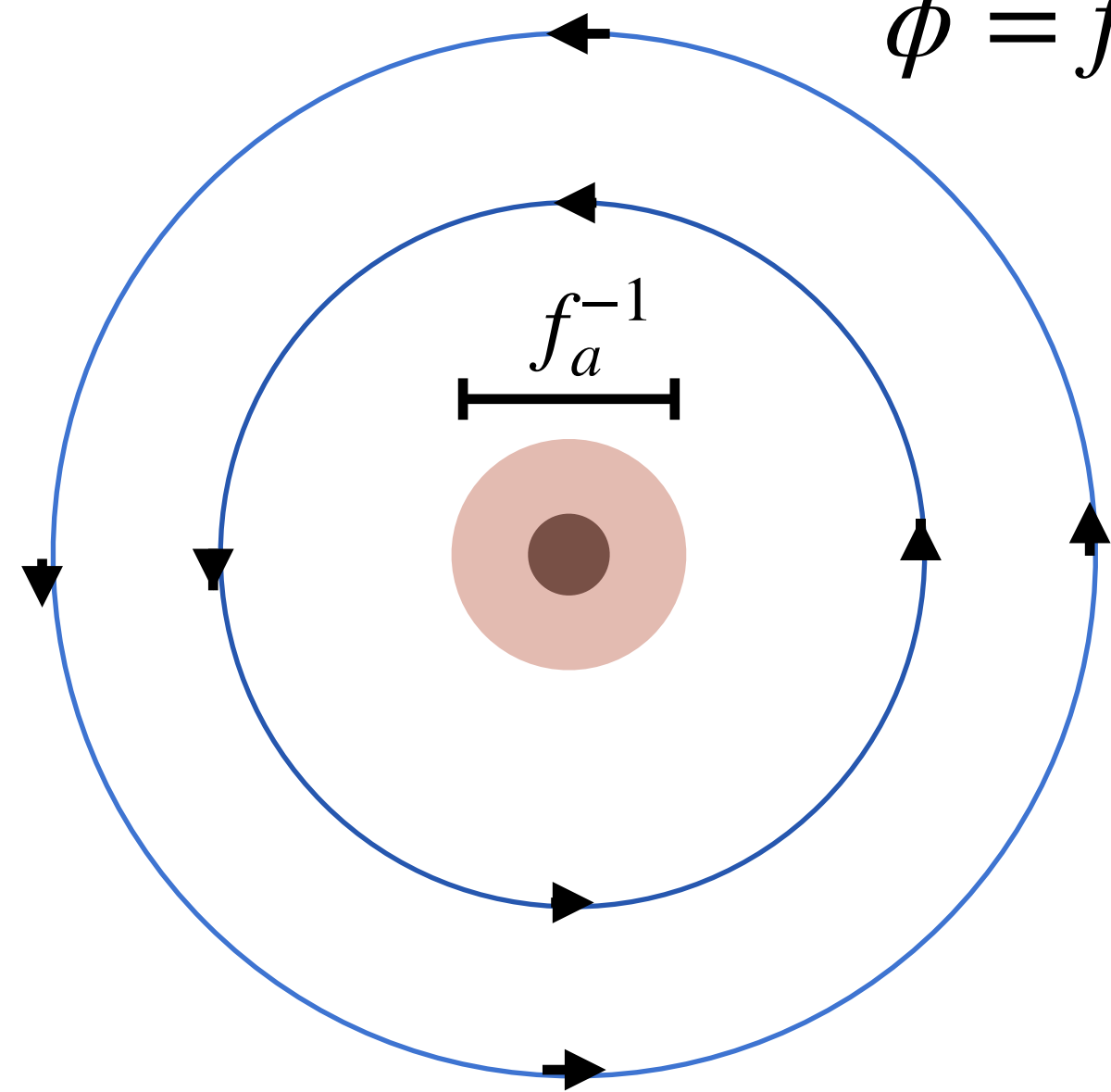
$$\mu \simeq \pi f_a^2 \log \left(\frac{d}{f_a^{-1}} \right)$$

$$U(\phi) = (|\phi|^2 - f_a^2)^2$$

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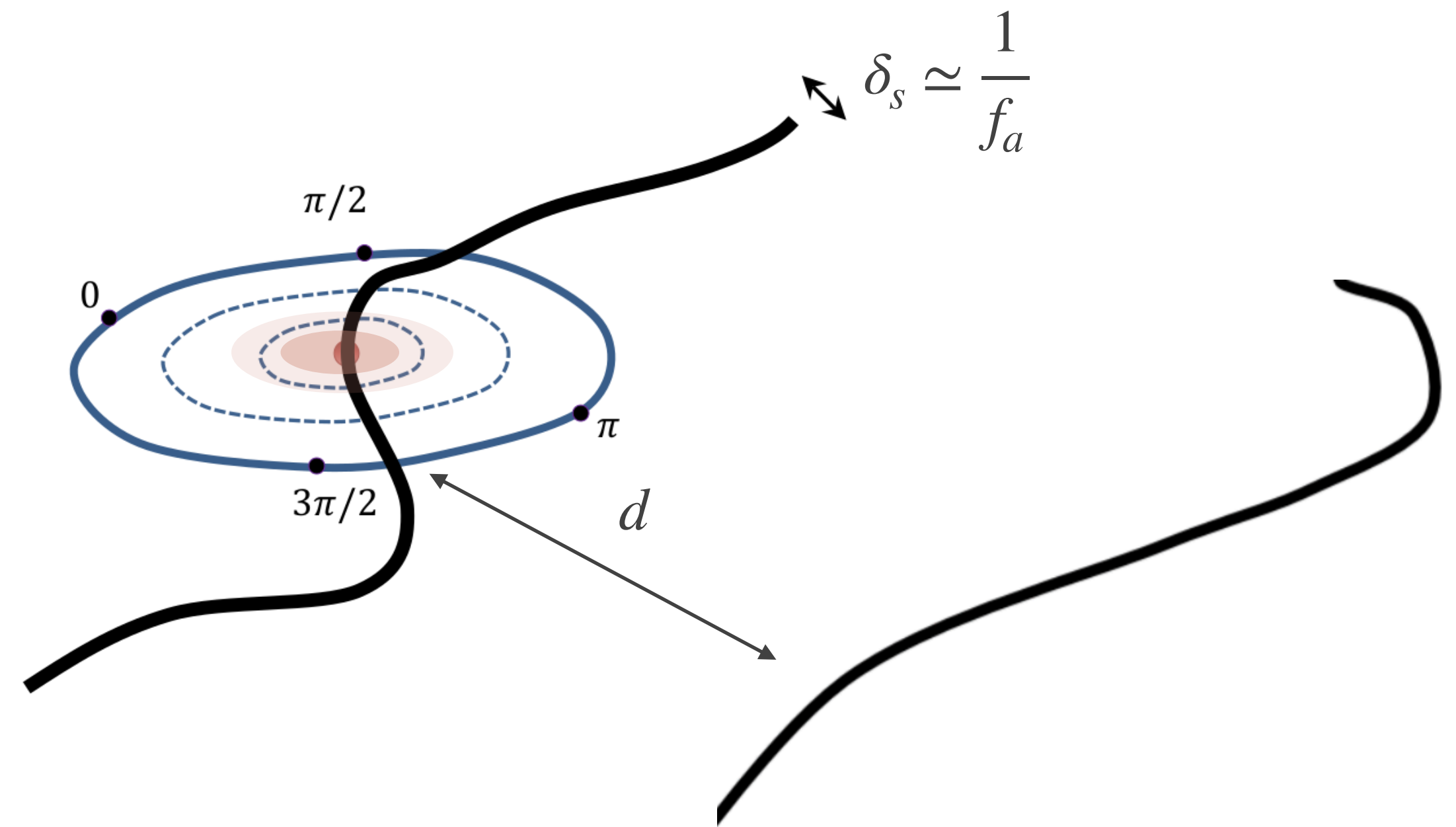
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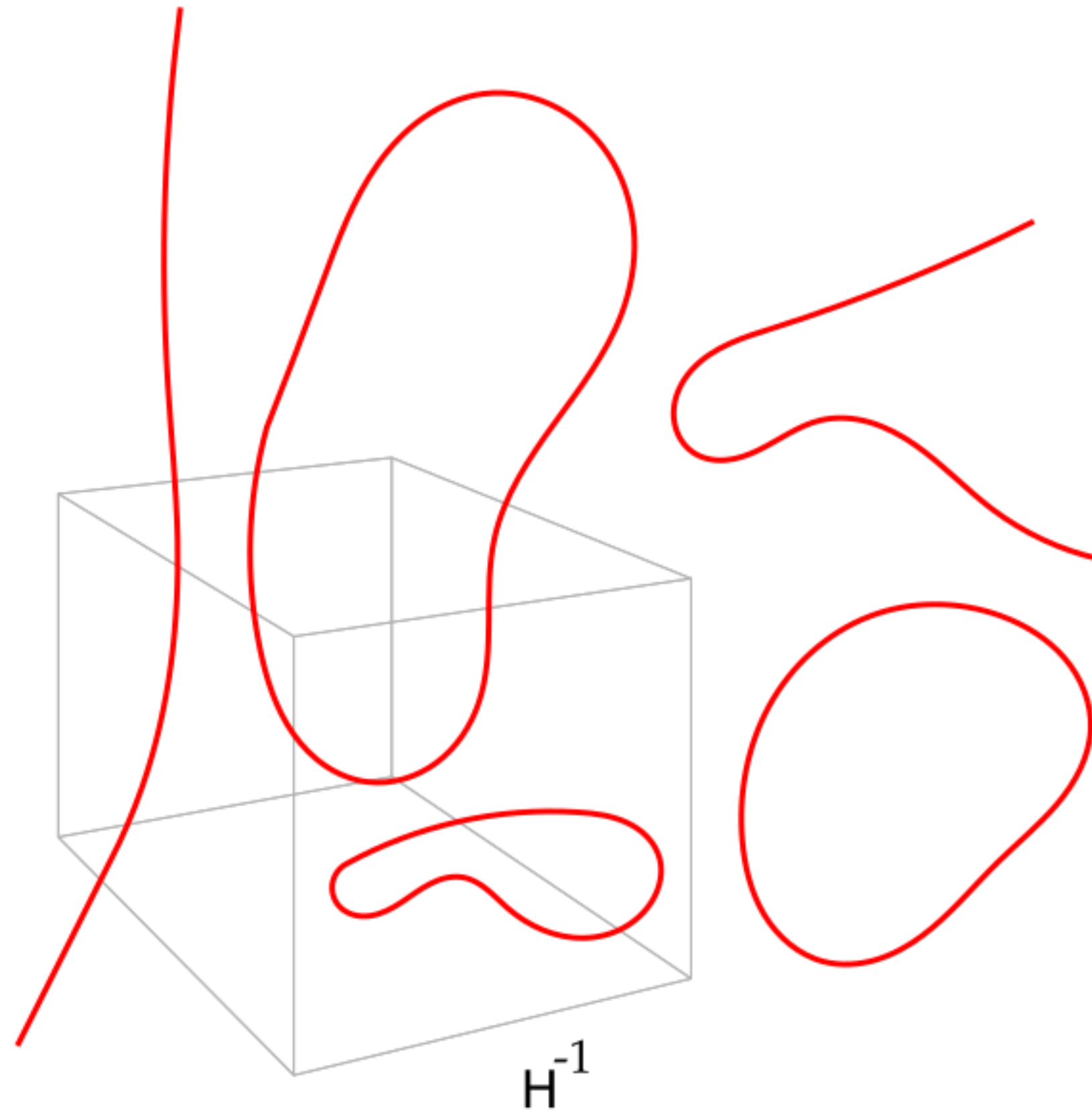
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Early universe:

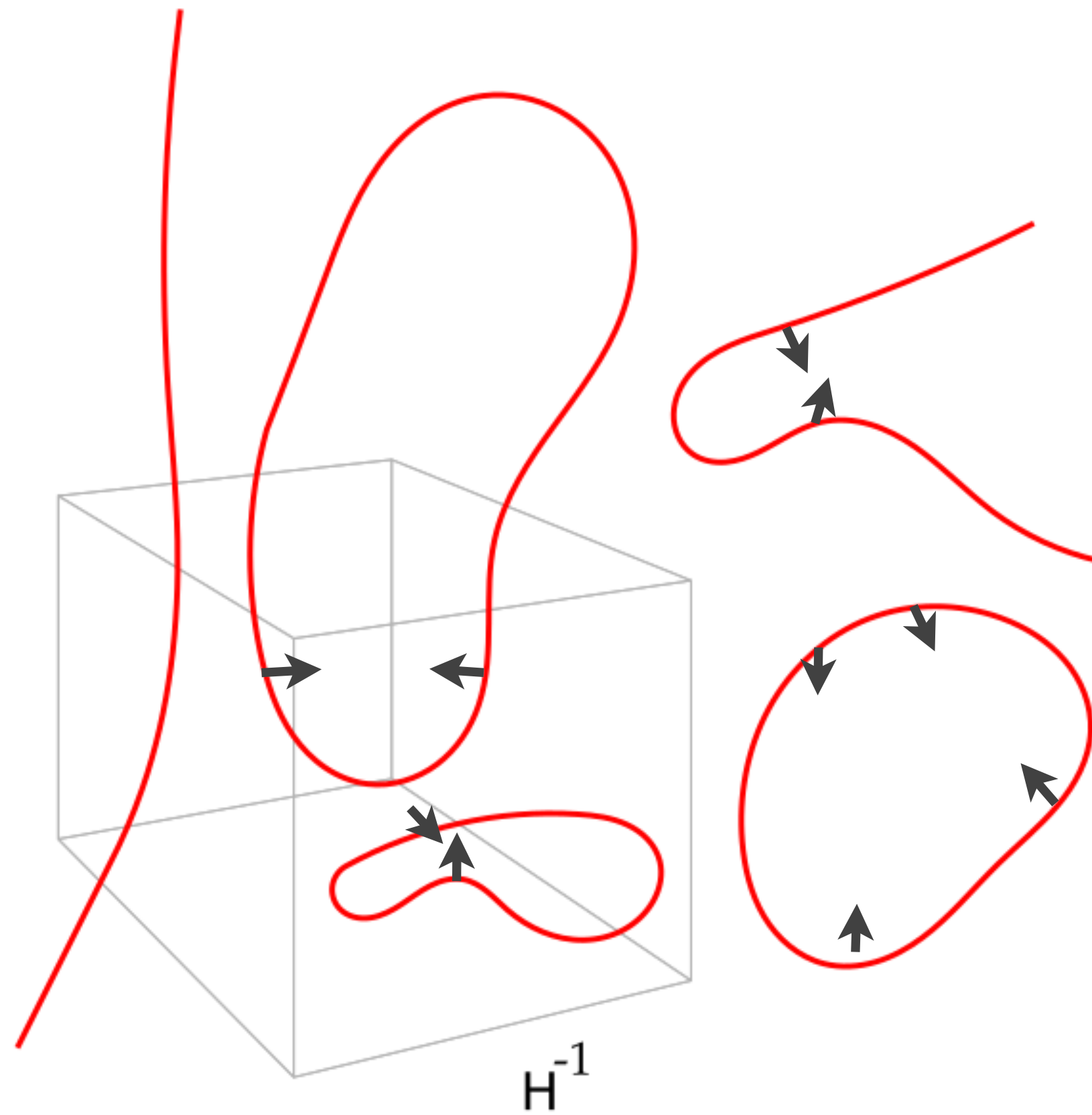
$$\mu \simeq \pi f_a^2 \log \left(\frac{f_a}{H} \right)$$

↑
Grows
with
time

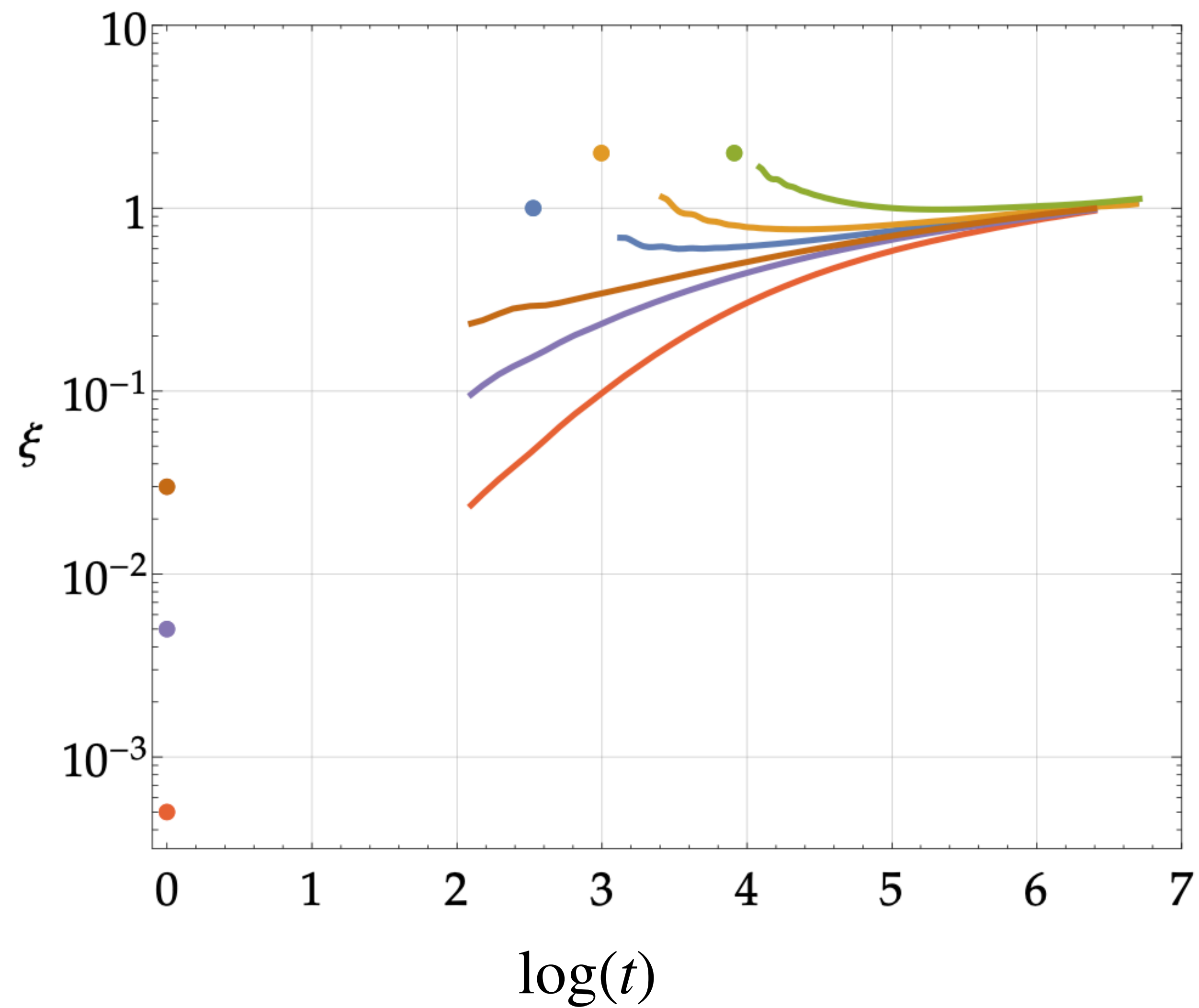
Scaling solution



Scaling solution



Scaling solution



Small violations of scaling

